

Radio morphology of southern narrow-line Seyfert 1 galaxies with Very Large Array observations*

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ABSTRACT

We present the results of new radio observations carried out with the Karl G. Jansky Very Large Array C-configuration at 5.5 GHz for a sample of southern narrow-line Seyfert 1 galaxies (NLS1s). This work increases the number of known radio-detected NLS1s in the Southern hemisphere, and confirms that the radio emission of NLS1s is mainly concentrated in a central region at kpc-scale and only a few sources show diffuse emission. In radio-quiet NLS1s, the radio luminosity tends to be higher in steep-spectrum sources and be lower in flat-spectrum sources, which is opposite to radio-loud NLS1s. This may be because the radio emission of steep NLS1s is dominated by misaligned jets, active galactic nucleus driven outflows, or star formation superposing on a compact core. Instead the radio emission of flat NLS1s may be produced by a central core that has not yet developed radio jets and outflows. We discover new NLS1s harbouring kpc-scale radio jets and confirm that a powerful jet does not require a large-mass black hole to be generated. We also find sources dominated by star formation. These NLS1s could be new candidates in investigating the radio emission of different mechanisms.

Key words: galaxies: active – galaxies: jets – galaxies: nuclei – galaxies: Seyfert – radio continuum: galaxies.

1 INTRODUCTION

Narrow-line Seyfert 1 galaxies (NLS1s), which are a subclass of active galactic nuclei (AGNs), are identified by their unusual optical spectral properties. They exhibit Balmer lines from the broad-line region (BLR) with a full width at half-maximum of $\text{FWHM}(\text{H}\beta) < 2000 \text{ km s}^{-1}$ that is only slightly broader than forbidden lines, and a flux ratio of $[\text{O III}] \lambda 5007/\text{H}\beta < 3$ (Osterbrock & Pogge 1985; Goodrich 1989). Besides, NLS1s show several extreme optical and X-ray properties, such as strong Fe II multiplets emission (Boroson & Green 1992), frequently observed blueshifted line profiles (Zamanov et al. 2002; Leighly & Moore 2004; Boroson 2005), soft X-ray spectra (Boller, Brandt & Fink 1996; Wang, Brinkmann & Bergeron 1996; Leighly 1999a), and rapid X-ray variability (Pounds, Done &

Osborne 1995; Leighly 1999b). Many studies suggest that NLS1s are powered by a relatively low-mass central black hole with $M_{\text{BH}} \sim 10^{6-8} M_{\odot}$, with respect to broad-line Seyfert 1 galaxies (BLS1s) with $M_{\text{BH}} \sim 10^{7-9} M_{\odot}$ (Järvelä, Lähteenmäki & León-Tavares 2015; Cracco et al. 2016; Chen et al. 2018). Since the bolometric luminosity of NLS1s is comparable to that of BLS1s, a low black hole mass corresponds to a high Eddington ratio (Boroson & Green 1992; Collin & Kawaguchi 2004). These properties suggest that NLS1s are likely to be a young and fast-growing phase of AGNs (Grupe 2000; Mathur 2000). Another possibility is that the low black hole mass of NLS1s is just an inclination effect, due to the pole-on orientation of a disc-like BLR that prevents us from seeing any Doppler broadening (Decarli et al. 2008). However, several studies reveal that this hypothesis is likely not true in the majority of NLS1s and the differences between BLS1 and NLS1 populations should be intrinsic rather than just due to orientation effects (Kollatschny & Zetzl 2011, 2013; Järvelä et al. 2017; Komossa 2018; Vietri et al. 2018; Berton et al. 2020a).

* Tables 1, 2, and 3 are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qc> at?J/MNRAS/vol/page

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The majority of NLS1 population is radio-quiet (RQ) with a radio loudness of $R < 10$.¹ The radio-loud (RL) fraction of NLS1s is very small, only ~ 7 per cent with $R > 10$ and ~ 2.5 per cent with $R > 100$ (Komossa et al. 2006), compared with that of BLS1s and quasars about ~ 10 – 15 per cent (Ivezić et al. 2002). However, the definition of radio loudness could be misleading, especially for NLS1s (Foschini 2011b; Berton et al. 2020b). For instance, star formation can produce strong radio emission at low frequencies in late-type galaxies. AGNs with such kind of host galaxies may be classified as RL even without launching a relativistic jet (Caccianiga et al. 2015; Ganci et al. 2019). On the contrary, if the relativistic jet is faint, or misaligned, or absorbed by ionized material at low frequencies, the radio emission is weak thus the AGN may be classified as RQ (Foschini 2011a, 2012). Moreover, Sikora, Stawarz & Lasota (2007) found a similar dependence of the radio loudness increases with the decreasing Eddington ratio in both spiral-hosted and elliptical-hosted AGNs, which implies that the radio loudness alone has little meaning if the brightness and accretion rate are not considered. Indeed, radio observations found some jet-like structures in RQ Seyfert galaxies (Giroletti & Panessa 2009; Mundell et al. 2009; Doi et al. 2013). Furthermore, Lähteenmäki et al. (2018) found strong radio emission at ~ 1 Jy level at 37 GHz from sources formerly classified as RQ. Such strong radio emission at such a high frequency cannot be explained by radiation from supernova remnants and can only be produced by a relativistic jet. Indeed, Padovani (2017) suggested that AGNs should be classified as jetted and non-jetted AGNs, instead of using the R parameter. Jetted sources are characterized by powerful relativistic jets. Instead, non-jetted sources do not launch a jet but may have outflows with considerably less power compared to jetted sources. This classification is based on physical differences rather than observational phenomena.

RL NLS1s are likely to harbour powerful relativistic jets, which are the predominant origin of the radio, optical, partially X-ray, and in some cases, γ -ray radiation. The discovery of γ -ray emission from RL NLS1s detected by the *Fermi* Large Area Telescope (LAT) confirmed the presence of relativistic beamed jets in this type of AGNs along with blazars and radio galaxies (Abdo et al. 2009a, b, c; Foschini et al. 2011, 2015; D’Ammando et al. 2015, 2016; Liao et al. 2015; Yao et al. 2015; Paliya & Stalin 2016; Berton et al. 2017; Lähteenmäki et al. 2018; Paliya et al. 2018; Yao et al. 2019). In addition, radio observations of several RL NLS1s with the Very Long Baseline Array (VLBA) showed that they exhibit pc-scale radio jets, compact radio morphologies, flat or inverted radio spectra, high brightness temperatures, and a significant degree of polarization (Abdo et al. 2009a, b; Gu et al. 2015; Lister et al. 2018). Some radio monitoring campaigns also revealed the presence of fast and intense variability which is a common feature of relativistic jets (Foschini et al. 2012; Angelakis et al. 2015; Fuhrmann et al. 2016). In some cases, an apparent superluminal motion was observed as well (Lister et al. 2013, 2016, 2019). These findings pose a challenging question about the generation and evolution of the jet system in NLS1s harbouring a relatively undermassive black hole.

The origin of the radio emission in RQ NLS1s is not established yet. The possible mechanisms might include a low-power jet, an AGN-driven outflow, an accretion disc corona, star formation, and a combination of them (Panessa et al. 2019). Radio observations can partly distinguish these physical scenarios. A compact radio core at pc-scale with a flat or inverted spectral slope is probably associated

with the corona (Guedel & Benz 1993; Laor & Behar 2008), the jet-base (Blandford & Königl 1979; Reynolds 1982), or the outflow-base (Lena et al. 2015). This central core is thought to be driven by the non-thermal processes with a high brightness temperature of $T_b > 10^7$ K (Blundell & Beasley 1998). The jet/outflow base may physically coincide with the corona (Merloni & Fabian 2002; King, Lohfink & Kara 2017), and they are indistinguishable in the radio images. The linear extended radio emission resolved at various scales from pc to kpc could be related to different components of the jet and/or outflow and be characterized by steep spectral indexes. The difference between a jet and an outflow depends on whether the ejected plasma is collimated or not (Begelman, Blandford & Rees 1984). At present, it is hard to separate them based on radio information only. Star formation activity can produce both free-free emission and synchrotron emission (Condon 1992). The radio images exhibit host-like distributed diffuse emission surrounding the central core with a steep spectral slope and a low surface brightness of $T_b < 10^5$ K (Oriente et al. 2015). However, these radio properties are limited by the resolution and sensitivity of the radio observations.

In order to study the radio properties of NLS1s and to identify sources with large radio jets, we proposed new observations with the Karl G. Jansky Very Large Array (VLA, Proposal ID: VLA/18B-126). This paper presents the main results of this survey and is organized as follows. In Section 2, we describe the data reduction, in Section 3 we present the data analysis, in Section 4 we discuss the main results, and in Section 5 we provide the summary. Throughout this work, we adopt a standard Λ cold dark matter cosmology with a Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.73$, and $\Omega_M = 0.27$ (Komatsu et al. 2011). We assume the flux density and spectral index convention at observed frequency is $S_\nu \propto \nu^{-\alpha}$.

2 DATA REDUCTION

The targets were selected from Chen et al. (2018), who presented a catalogue of NLS1s in the Southern hemisphere. It includes 168² NLS1s which were classified according to their optical spectra from the Six-degree Field Galaxy Survey (6dFGS) (Jones et al. 2009). We observed 62 sources that are higher than the declination limit of the VLA (-30°). The redshifts of these targets range from 0.01 to 0.44.

We were granted 55 h observing time in C-configuration centred at 5.5 GHz with a bandwidth of 2 GHz and an angular resolution of 3.5 arcsec for this NLS1 sample. These observations were carried out between November 2018 and February 2019 with an exposure time of 30 min on each target, yielding an image sensitivity of $\sim 7 \mu\text{Jy beam}^{-1}$. Such a resolution and sensitivity enable us to see diffuse emission that is missing in A-configuration and discover extended structure. We reduced the data using the VLA calibration pipeline version 5.4.0 and the Common Astronomy Software Applications (CASA) version 5.5.0. A standard flux density calibrator, 3C 147 or 3C 286, was used for every target. The measurement set is split into 16 spectral windows (centred at 4.55, 4.68, 4.81, 4.94, 5.06, 5.19, 5.32, 5.45, 5.55, 5.68, 5.81, 5.94, 6.06, 6.19, 6.32, and 6.45 GHz).

To produce the radio maps, we used a cell size of 0.5 arcsec to properly sample the beam that has an FWHM of 3.5 arcsec. The maps were created in a region of 2048×2048 pixels centred on the source coordinates to check for the presence of nearby sources. We modelled the main target along with the nearby sources using the

¹The radio loudness is defined as the ratio of 5 GHz radio flux density and optical B -band flux density $R = S_{5 \text{ GHz}}/S_{4400 \text{ \AA}}$ (Kellermann et al. 1989).

²There are 167 NLS1s in Chen et al. (2018). We included one more source IRAS 13224–3809 as it was classified as NLS1 in Boller et al. (1993) and also detected with the 6dFGS.

CLEAN algorithm in all spectral windows to avoid the contamination from sidelobes. If a bright source is outside the mapped region and its sidelobes do affect the target, we enlarged the image to include the bright source. In this way, we can model the extra source to reduce its effect. We used a natural weighting to create the first tentative image. For some faint sources, the noise level approaches the predicted image sensitivity of $\sim 7 \mu\text{Jy beam}^{-1}$ after the first cleaning, and we did not proceed any further. Conversely, the noise level of some bright sources remained high after the first cleaning, and we applied iterative cycles of phase-only self-calibration with a Briggs weighting, intermediate between natural and uniform, on the visibilities to improve the dynamic range of the final maps.

We modelled the source with a Gaussian fit on the image plane and deconvolved it from the beam, to recover the radio position, the core size and its position angle, and the integrated and peak flux densities, S_{int} and S_p , centred at 5.5 GHz. If the core size is too small to be determined compared to the beam size, we adopted a half beam size as an upper limit for the core size. We only modelled the central core if the source is resolved in more than one component. Thus, we remark that our flux densities could be underestimated if the source displays some resolved structures. For a few sources, the integrated flux density is less than the peak flux density due to either the small core size with respect to the beam size or errors in the estimation of the integrated flux density. In this case, we treated S_p as the upper limit of S_{int} . The root mean square (RMS) was estimated in a source-free region. The errors of each measurement were produced by CASA. We obtained the integrated flux densities ranging from 0.05 to 21.28 mJy and the average RMS of $10 \mu\text{Jy beam}^{-1}$.

In addition, we split the data in 0–7 and 8–15 spectral windows with a bandwidth of 1 GHz and created another two maps centred at 5 and 6 GHz following the reduction procedures mentioned above. We did not use a narrower bandwidth in order to minimize the uncertainty of each measurement. We then smoothed the 6 GHz image with a Gaussian kernel to match the resolution of the 5 GHz image using the task IMSMOOTH, with the aim of evaluating the in-band spectral indexes.

In total, 49 sources have a detection in these observations. Thirteen sources are not detected or only detected at 3σ level because (a) the target has either weak radio emission or high noise background, or (b) the image is affected by strong sidelobes of a nearby bright source. The detection rate at 5.5 GHz of this NLS1 sample is 79 per cent (49/62), which suggests that most NLS1s are radio emitters when observed with an adequate sensitivity. The radio maps of southern NLS1s observed with the VLA are shown in Fig. A1 and Appendix B.

3 DATA ANALYSIS

3.1 5.5 GHz

We calculated the radio loudness according to the definition

$$R = S_{5\text{GHz}}/S_{4400\text{\AA}}, \quad (1)$$

where $S_{4400\text{\AA}}$ is the flux density in optical B band derived from the optical spectra (Chen et al. 2018). Despite the boundary of $R = 10$ being arbitrary, it can give us an idea of the fraction of radio flux with respect to optical flux and make our results comparable to other studies using this parameter. There are 45 RQ NLS1s in the sample. Only four sources (J0122–2646, J0452–2953, J0846–1214, and J2021–2235) have $R > 10$, but not very high with $R \sim 14$ –38. The RL fraction is 6.5 per cent (4/62) in the whole sample. This is consistent with Komossa et al. (2006) who found that the fraction of

RL NLS1s is about 7 per cent. However, the radio loudness strongly depends on whether the optical and radio observations include the whole galaxy (both AGN and host) or only the nucleus (Ho & Peng 2001; Kharb et al. 2014). Besides, the RL fraction tends to be lower in the nearby universe and be higher with an increasing large-scale environment density (Cracco et al. 2016; Järvelä et al. 2017).

To study the concentration of radio emission, we estimated the fraction of peak to integrated flux density following Berton et al. (2018)

$$f = S_p/S_{\text{int}}. \quad (2)$$

We considered that the radio emission is concentrated (C) in a compact core at the centre if $f \geq 0.5$ and is diffuse (D) in a region slightly larger than the angular resolution if $f < 0.5$. In the case of the upper limit of S_{int} equal to S_p , we get $f = 1$ which means that all the flux densities are concentrated in a central beam. These fractions range from 0.3 to 1. For the majority of NLS1s (45 objects), the radio emission is mainly concentrated in a central core. Only in four RQ sources (J0447–0403, J0622–2317, J0845–0732, and J1937–0613), the radio emission spreads over a region somewhat larger than the core. This could be because Seyfert galaxies and low-luminosity AGNs are typically unresolved at arcsec-scale (Ho & Ulvestad 2001) and confirms that RL sources do not show a large fraction of diffuse emission.

In addition, we measured the in-band spectral index between 5 and 6 GHz $\alpha_{\text{in-band}}$ by modelling the spectrum with a power law and giving the definition

$$\alpha = -\frac{\log(S_2/S_1)}{\log(\nu_2/\nu_1)}, \quad (3)$$

where S_1 and S_2 are the integrated flux densities at the observing frequencies $\nu_1 = 5$ GHz and $\nu_2 = 6$ GHz, respectively. There are 18 sources having a flat (F) radio spectrum with $\alpha < 0.5$ and 31 sources having a steep (S) radio spectrum with $\alpha \geq 0.5$. The in-band spectral indexes range from -1.91 to 3.11 with a mean value of $\alpha_{\text{in-band}} = 0.53$. Nevertheless, we remark that the in-band spectral indexes have large errors due to the narrow bandwidth between 5 and 6 GHz.

The integrated and peak luminosities, L_{int} and L_p , at the observing frequency $\nu = 5.5$ GHz were derived by

$$L_\nu = 4\pi D_L^2 \nu S_\nu (1+z)^{(\alpha-1)}, \quad (4)$$

where D_L is the luminosity distance at a cosmological redshift and α is the in-band spectral index. Throughout this work, the luminosity L is measured in unit erg s^{-1} . The integrated luminosities at 5.5 GHz have a range of $\log L_{\text{int}} \sim 37.5$ –40.7 with an average value of $\log L_{\text{int}} = 39.0$.

Besides, we also collected the central black hole masses from Chen et al. (2018) who estimated the virial black hole mass via the BLR size and the $H\beta$ line dispersion. The BLR size was calculated by the relation between the BLR size and the 5100 Å continuum luminosity (Bentz et al. 2013), and this relation was derived from reverberation mapping method (Peterson et al. 2004; Peterson 2014). The black hole masses have a range of $\log(M_{\text{BH}}/M_\odot) \sim 6.0$ –7.4 with a mean value of $\log(M_{\text{BH}}/M_\odot) = 6.8$, which are typical for NLS1s. During this estimation, the errors of the black hole mass are 0.1–0.2 dex. Without taking into account the radiation pressure, which is likely to be important in highly accreting AGNs such as NLS1s, there will be an additional average scatter of 0.2 dex (Marconi et al. 2008). Thus, the uncertainty of the black hole mass is about 0.5 dex.

The distributions of redshift, radio loudness, flux concentration, in-band spectral index, integrated luminosity at 5.5 GHz, and black

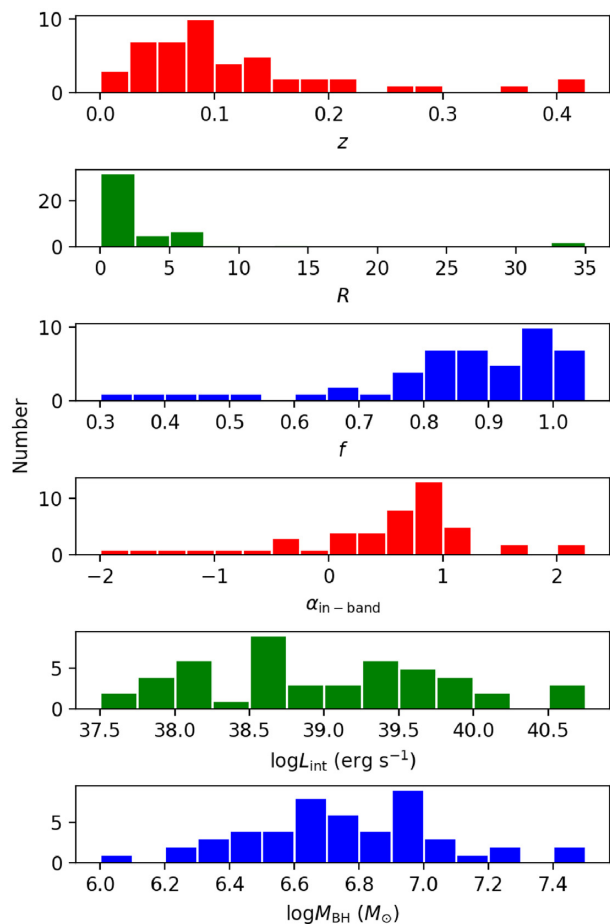


Figure 1. The distributions of redshift, radio loudness, flux concentration, in-band spectral index, integrated luminosity at 5.5 GHz, and black hole mass (from top to bottom panels) using 0.025, 2.5, 0.05, 0.25, 0.25, and 0.1 bin width, respectively.

hole mass are plotted in Fig. 1. The coordinates, redshifts, scales, core sizes, position angles, and integrated and peak flux densities and luminosities at 5.5 GHz of the sample are listed in Table 1. The radio loudness, flux concentration, in-band spectral index between 5 and 6 GHz, different classifications, flux density in optical B band, and black hole mass are reported in Table 2.

3.2 1.4 GHz

For comparison of the classifications derived at different frequencies, we collected the radio data from other sky surveys. There are 21 NLS1s detected at 1.4 GHz, including five sources from the Faint Images of the Radio Sky at Twenty-Centimeters (FIRST) with a resolution of 5 arcsec (Helfand, White & Becker 2015) and 16 sources from the NRAO VLA Sky Survey (NVSS) with a large restoring beam of 45 arcsec (Condon et al. 1998). If a source was detected with both surveys, we applied the flux density from the FIRST as it has a better angular resolution. Their flux densities range from 1.2 to 59.8 mJy.

We computed the radio loudness for these 21 sources using equation (1). The radio flux density at 5 GHz is derived from that at 1.4 GHz under the hypothesis of a power-law spectrum. As the spectral slope at 5.5 GHz cannot precisely represent that at 1.4 GHz, we applied a conservative spectral index of $\alpha = 0.5$ that is consistent

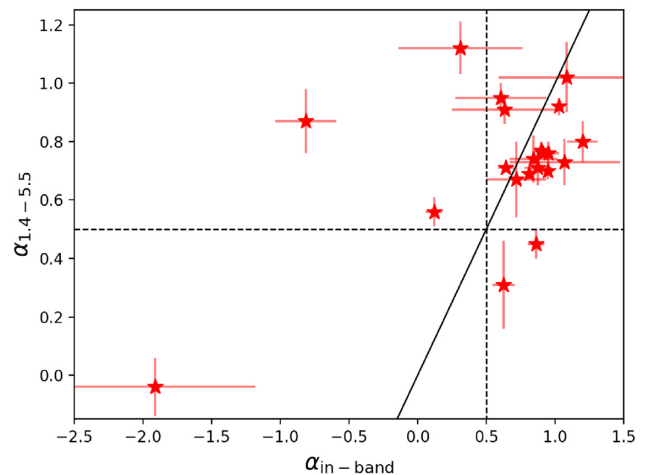


Figure 2. The spectral index between 1.4 and 5.5 GHz versus the in-band spectral index. The solid black line indicates the 1:1 ratio. The dashed lines are $\alpha_{1.4-5.5} = 0.5$ and $\alpha_{\text{in-band}} = 0.5$.

with the average in-band spectral index, to avoid the large scatters for individual sources. The RL and RQ classifications at the two frequencies appear to be robust, as they are consistent in 20 objects and differ in just one case, J1511–2119, a source lying close to the $R \sim 10$ threshold, which is classified as RQ using the 5.5 GHz flux but becomes RL using the 1.4 GHz flux.

We also measured the spectral index between 1.4 and 5.5 GHz $\alpha_{1.4-5.5}$ using equation (3). They range from -0.04 to 1.12 and have an average of 0.72 . As a comparison, we plotted $\alpha_{1.4-5.5}$ versus $\alpha_{\text{in-band}}$ for 21 NLS1s shown in Fig. 2. We note that 16 sources keep the same spectral class, while five sources (J0354–1340, J0400–2500, J0549–2425, J0622–2317, and J1638–2055) show a difference in the spectral class. Besides, 14 and 7 objects stay below and above the 1:1 ratio line indicating that the spectra become steeper and flatter towards higher frequencies, respectively. Nevertheless, we remark that the spectral index between 1.4 and 5.5 GHz could be biased due to flux density variability and different beam sizes during these non-simultaneous observations.

The luminosity at 1.4 GHz of 21 sources was computed using equation (4) as well, where the spectral index between 1.4 and 5.5 GHz was used. They come within a range of $\log L_{1.4} \sim 37.0$ – 40.5 and have a mean value of $\log L_{1.4} = 39.2$. The flux density, luminosity, radio loudness, spectral index between 1.4 and 5.5 GHz, and different types based on the 1.4 GHz detections are listed in Table 3.

4 DISCUSSION

4.1 Radio emission concentration

We plotted the distribution of the in-band spectral index $\alpha_{\text{in-band}}$ and the concentration of radio emission f shown in Fig. 3. As expected, those sources whose flux densities are not centrally concentrated, basically have steep spectral indexes indicating that the diffuse emission is mainly from an optically thin region. For the majority of NLS1s, their flux densities are concentrated in a central core and their spectral indexes have a wide range from flat to steep. This suggests that the radio emission of NLS1s originated from either the core or the extended structure, is mainly concentrated in a region at arcsec-scale. In addition, Berton et al. (2018) found that RQ NLS1s usually

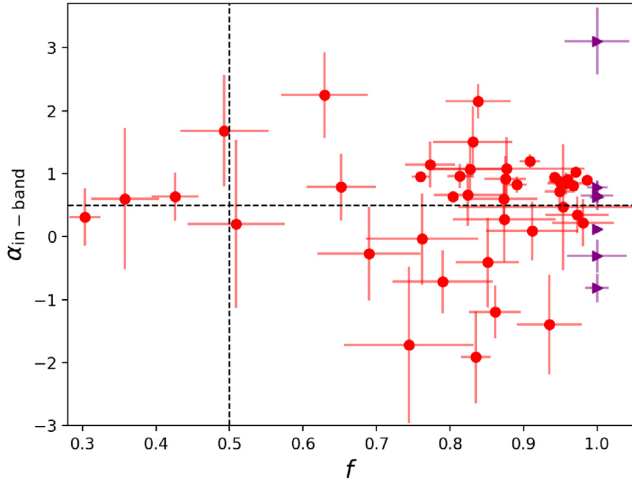


Figure 3. The distribution of the in-band spectral index and the concentration of radio emission. Sources with $f = 1$ are marked with right triangles in purple, the others are marked with circles in red. The dashed lines are $\alpha_{\text{in-band}} = 0.5$ and $f = 0.5$.

show diffuse emission surrounding a central core at the resolution of the VLA A-configuration. At mas-scale, Orienti & Prieto (2010) found that the radio emission of Seyfert galaxies is concentrated in the central core if a flat spectrum is present, and is extended or diffuse in a large region if a steep spectrum is present. Hence it may reveal that the core emission is mainly concentrated at pc-scale, and the extended or diffuse emission can spread over a region at kpc-scale.

4.2 Radio spectral slope

We divided this sample into two subsamples: flat and steep spectral slopes based on the in-band spectral index. The distributions of radio luminosity and black hole mass for these two subgroups are shown in Fig. 4. We applied the Kolmogorov–Smirnov (K–S) test on these parameters to verify if the flat and steep sources are actually statistically different. The null hypothesis is that two distributions originate from the same parent population. We applied the rejection of the null hypothesis at a 99 per cent confidence level corresponding to a value of $p \leq 0.01$. The K–S test of radio luminosity ($p \sim 7.6 \times 10^{-3}$) suggests that the flat and steep objects have different origins. Conversely, the K–S test of black hole mass ($p \sim 0.86$) suggests that these two subsamples originate from the same parent population. We further applied the K–S test on the redshift and found that the flat and steep sources have similar distributions ($p \sim 0.77$), which suggests that the different origins of these two subgroups on the radio luminosity are not a selection effect and could be intrinsic.

On one hand, these distributions imply that the radio luminosity may be not related to the black hole mass. On the other hand, unlike RL NLS1s, flat-spectrum objects have a higher radio luminosity compared to steep-spectrum objects (Berton et al. 2018). The majority of NLS1s in our sample are RQ; the flat sources tend to have a lower radio luminosity than the steep sources. This may be related to the different radio mechanisms. The steep-spectrum radio emission may be originated from an optically thin synchrotron region, where a misaligned jet, an AGN-driven outflow, or star-forming activity may be present and superpose on an optically thick core. The flat-spectrum radio emission may be produced in a central compact core, which has not yet developed a radio jet or outflow. Therefore, the steep sources tend to have a higher radio luminosity than the flat sources

in the RQ NLS1 population. However, these results should be taken with caution, since the separation of the flat and steep sources is somewhat arbitrary, due to a large fraction of sources cluster around the threshold of $\alpha = 0.5$ with relatively large uncertainties.

4.3 Pieces of evidence for radio jets

We find one source J0354–1340 that is a jetted NLS1 (details seen in Appendix A3). Its morphology favours an Fanaroff–Riley type II (FR II) radio galaxy (Fanaroff & Riley 1974). The central core is compact and bright at radio frequencies. Two components located to the south and north of the central core, are probably the radio lobes showing diffuse emission. Additionally, a hotspot where the relativistic jet terminates in, may be present in the end of the southern lobe. The projected linear sizes of the radio lobes are around a hundred kpc, which are consistent with those of FR II radio galaxies (Hardcastle et al. 1998; Kharb et al. 2008). The luminosity of the southern lobe is brighter than that of the northern one, indicating that the southern component is the approaching jet and the northern component is the receding jet. The NVSS map shows an elongated structure which is consistent with the VLA C-configuration map exhibiting a core-jet/lobe morphology. In the 5.5 GHz image, the radio lobes are resolved, and the central compact core is probably the jet-base having a flat in-band spectral index. In the 1.4 GHz image, the radio lobes are unresolved, and the radio emission from both the core and the lobes is included giving a slightly steep slope between 1.4 and 5.5 GHz.

The jet speed can be calculated based on the assumptions that the morphology and the physical parameters of the FR II radio galaxy are intrinsically symmetric. The only difference is caused by different Doppler enhancements, leading to an amplification of the approaching jet with an observed flux of f_+ and a dimming of the receding jet with an observed flux of f_- . The flux ratio can be derived from the Doppler factors (Beckmann & Shrader 2012)

$$\frac{f_+}{f_-} = \left(\frac{1 + \beta \cos \theta}{1 - \beta \cos \theta} \right)^{(2+\alpha)}, \quad (5)$$

where $\beta = v/c$, θ is the angle of the jet with respect to the line of sight, and α is the spectral index of the continuum emission. If we assume that the jet propagates close to the speed of light ($\beta \sim 1$), this gives the radial velocity of $v \cos \theta \sim 0.25c$ and the projected velocity of $v \sin \theta \sim 0.97c$. The age of the approaching jet can also be evaluated using the projected linear size over the projected velocity, resulting in a lower limit for the age of $\sim 3 \times 10^5$ yr. This lies at the low tail of the age range of 10^5 – 10^7 yr found for other RL NLS1s (Doi et al. 2012; Richards & Lister 2015; Rakshit et al. 2018).

We further estimated the jet power following two methods. The first one is adopting the relationships based on the radio core luminosity at 15 GHz in Foschini (2014)

$$\log P_{\text{rad}} = (12 \pm 2) + (0.75 \pm 0.04) \log L_{15} \quad (6)$$

and

$$\log P_{\text{kin}} = (6 \pm 2) + (0.90 \pm 0.04) \log L_{15}, \quad (7)$$

where L_{15} was derived from the radio core flux at 5.5 GHz and the in-band spectral index assuming a power-law radio spectrum. We obtained the same radiative and kinetic powers of $\log P_{\text{rad}} = \log P_{\text{kin}} = 42.0 \pm 2.6$ (in unit erg s^{-1}), revealing that the particles and radiation have equivalent contributions to the jet power. The total jet power was calculated via the sum of the kinetic and radiative powers, resulting in $\log P_{\text{jet}} = 42.27 \pm 1.81$ (in unit erg s^{-1}). The

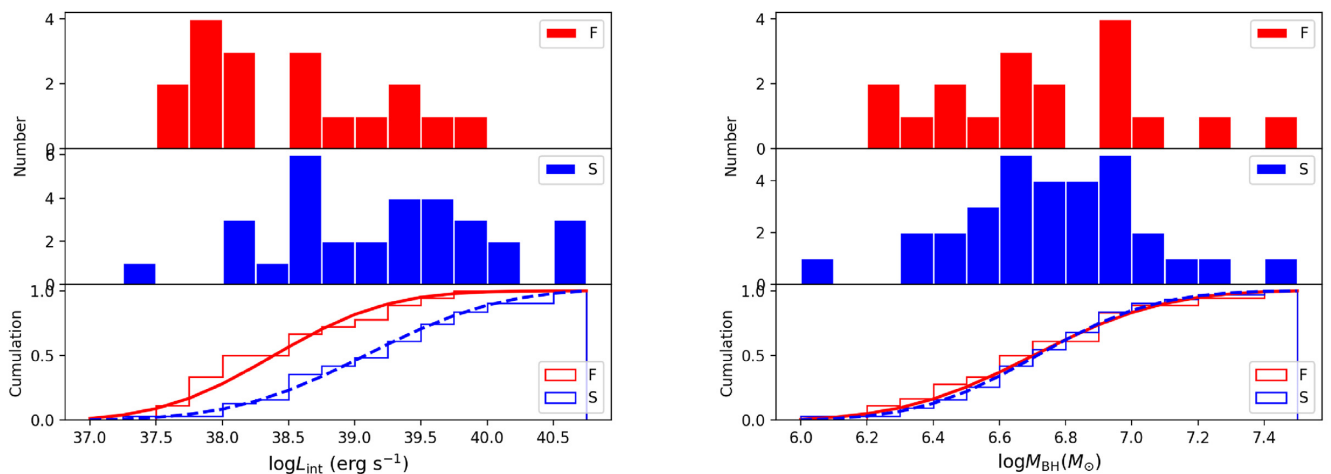


Figure 4. The numerical distributions of radio luminosity (*left-hand panel*) and black hole mass (*right-hand panel*) for the flat (*top panels*) and steep (*middle panels*) subsamples. *Bottom panels*: the red solid and blue dashed lines represent the cumulative distributions of the flat and steep subsamples, respectively.

second one is using the relationship based on the radio power at 1.4 GHz in Cavagnolo et al. (2010)

$$\log P_{\text{cav}} = (0.75 \pm 0.14) \log P_{1.4} + (1.91 \pm 0.18), \quad (8)$$

where $P_{1.4}$ is in unit $10^{40} \text{ erg s}^{-1}$ and P_{cav} is the cavity power in unit $10^{42} \text{ erg s}^{-1}$. Assuming $P_{\text{cav}} = P_{\text{jet}}$, this can give an estimate of the jet power with $\log P_{\text{jet}} = 43.51 \pm 0.20$ (in unit erg s^{-1}). The jet powers based on two measurements are consistent within errors and generally less powerful than flat-spectrum RL NLS1s (Foschini et al. 2015).

In addition, J1615–0936 may also be a jetted NLS1 (details seen in A10). It displays a core-jet/lobe structure though the radio emission is not so powerful. The western component and the eastern diffuse emission may be the radio jets or lobes spreading to a projected linear size of ~ 8 kpc. At such a scale, the radio emission is probably produced by a misaligned jet. A steep spectral slope indicates that the source is dominated by extended emission.

Two sources, J0952–0136 and J2021–2235, are well studied in literature and suggested to be jetted NLS1s (details seen in A7 and A12). J0952–0136 exhibits a resolved and linear extended structure at pc-scale and a slightly elongated structure at kpc-scale. The radio spectrum is steep showing that the radiation is originated from an optically thin source. These indicate that the radio emission is dominated by a misaligned jet at pc-scale (Doi et al. 2013, 2015). J2021–2235 displays a compact core at kpc-scale and a steep spectral slope. It has a remarkably high star formation rate which only account for up to ~ 30 per cent of the total radio luminosity (Caccianiga et al. 2015). Thus the presence of a jet is required to explain the non-thermal synchrotron emission. These suggest that a pc-scale jet and star formation are co-existing in this source (Berton et al. 2019).

Furthermore, other sources resembling J0952–0136 and J2021–2235 (e.g. J0452–2953, J0846–1214, and J1511–2119), may be candidates of harbouring a radio jet at pc-scale. They are characterized by a compact radio core at kpc-scale with relatively high luminosities and steep spectral slopes, which reveal that the strong radio emission is generated from an optically thin source and centrally concentrated at kpc-scale. However, high-resolution observations to resolve the central core are necessary to verify this interpretation.

These jetted NLS1s, especially the FR II-like NLS1 which has a smaller black hole mass than a regular FR II radio galaxy, confirm our current knowledge that the generation of relativistic jets does not

require a massive black hole. Indeed, radio emission from kpc-scale jets was discovered in about a dozen NLS1s (Whalen et al. 2006; Antón, Browne & Marchã 2008; Gliozzi et al. 2010; Doi et al. 2012, 2015; Richards & Lister 2015; Berton et al. 2018; Rakshit et al. 2018). Other powerful radio jets were also found in disc galaxies that are believed to be late-type galaxies (Keel et al. 2006; Morganti et al. 2011; Kotilainen et al. 2016; Järvelä, Lähteenmäki & Berton 2018; Olguín-Iglesias, Kotilainen & Chavushyan 2020). How is a powerful jet generated in an AGN with a relatively low-mass black hole? Other factors, such as the spin of black hole, may also play an important role in the evolution of relativistic jets (Laor 2000; Foschini 2012). Further studies will be helpful to unveil this picture.

4.4 Peculiar sources

J0000–0541 is a radio bright source (details seen in A1). It is currently unclear if the northwestern source is connected to the central target since no optical counterpart is associated with it on the image from the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS) (Chambers et al. 2016; Flewelling et al. 2016), or if they are two independent sources since no diffuse emission is connecting them. A dedicated analysis is needed to clarify if this object is jetted or non-jetted NLS1.

In addition, we find two sources, J0436–1022 and J1937–0613, probably dominated by star formation (details seen in A4 and A11). Their radio emission is extended or diffuse and has similar morphologies to their host galaxies in the optical images. Other sources whose radio emission is not centrally concentrated at kpc-scale, may also have a star-forming origin. They are characterized by extended or diffuse emission with low luminosities and steep spectral slopes. Further studies at infrared frequencies may help to clarify this hypothesis.

Another interesting source may be J0850–0318 (details seen in A6). The radio luminosity is comparable to jetted NLS1s. It shows a slightly extended morphology and a very steep spectral index with the C-configuration. This may suggest that the radio emission is originated from a relic emission due to a previous AGN activity episode (Congiu et al. 2017). However, such extended emission could be out-resolved with the A-configuration, resulting in a compact core and a spectral slope that is broadly consistent with optically thin synchrotron emission. Besides, it is unclear whether the western

source and the central target are interacting or not. More research is needed to find an answer.

5 SUMMARY

In this work, we present a new VLA survey using the C-configuration at 5.5 GHz of a southern NLS1 sample. There are 49 sources detected in these observations. We reduced their radio maps, measured their flux densities, and calculated their radio loudness, radio flux concentration, spectral index, and luminosities. Other properties were also collected from literature. The main results of this work are summarized as follow.

(1) This work increases the number of radio-detected NLS1s especially in the southern sky that remains largely unexplored. Among this sample, 47 sources are observed at 5 GHz for the first time, and 27 sources are new radio-emitting NLS1s without a previous observation at any radio frequencies. It provides good candidates for future observations with new facilities located in the Southern hemisphere, such as the Atacama Large Millimeter/submillimeter Array (ALMA) and the Square Kilometre Array (SKA).

(2) This study confirms that the majority of NLS1s are unresolved at kpc-scale. The radio emission is generally concentrated in a central region, independent of whether the origin is from a compact core or from extended structures. Only a few sources show diffuse emission with steep spectral slopes.

(3) We find that the radio luminosity of RQ NLS1s tends to be higher in steep-spectrum sources and be lower in flat-spectrum sources, which is in contrast to RL NLS1s. This may be because the radio emission of steep NLS1s is dominated by misaligned jets, AGN-driven outflows, or star formation superposing on a compact core, while the radio emission of flat NLS1s is produced by a central core which has not yet developed large radio jets or outflows. However, we should take with caution since the distinction between the flat and steep spectral sources is blurred due to a large fraction of sources cluster around the threshold and the statistical uncertainty is relatively high.

(4) We discovered new NLS1s harbouring kpc-scale radio jets, such as J0354–1340 and J1615–0936. Such jetted NLS1s confirm that a powerful jet does not require a large-mass black hole to be generated. We also find sources dominated by star formation, such as J0436–1022 and J1937–0613. These NLS1s could be new candidates in investigating the radio emission of different mechanisms.

Future high-resolution observations will be necessary to assess whether the radio properties of this sample have been correctly interpreted in terms of intrinsic structure. Further broad-band observations will also play an important role in unveiling the nature of NLS1s.

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DATA AVAILABILITY

The data underlying this article are available in the NRAO Science Data Archive at <https://archive.nrao.edu/archive/advquery.jsp>, and can be accessed with the project codes of VLA/18B-126 and VLA/19B-285.

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APPENDIX A: INDIVIDUAL SOURCES

We present an investigation on individual sources exhibiting a resolved or slightly resolved morphology in the VLA maps or having well-studied radio properties in literature.

A1 J0000–0541

This target exhibits a compact core at kpc-scale (Fig. B1, left-hand panel) with a relatively large radio power but small black hole mass compared to the averages of this sample. The flux density, luminosity, and black hole mass are $S_{\text{int}} \sim 1.16$ mJy, $\log L_{\text{int}} \sim 39.16$, and $\log (M_{\text{BH}}/M_{\odot}) \sim 6.36$, respectively. It has a steep spectrum with $\alpha_{\text{in-band}} \sim 1.20$. This source was also detected at 1.4 GHz associating with a counterpart of FIRST J000040.2–054101. It has a flux density of $S_{1.4} \sim 3.5$ mJy, a luminosity of $\log L_{1.4} \sim 39.03$, and a steep slope of $\alpha_{1.4-5.5} \sim 0.80$.

A northwestern source (centred at RA 00:00:38.53 and Dec. –05:40:37.63) has a separation of 34.8 arcsec from the central target. We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C1, left-hand panel). No optical counterpart is associated with it and no redshift is available. This nearby source shows a compact core with a flux density of $S_{\text{int}} \sim 2.05$ mJy. It has an associated counterpart of FIRST J000038.5–054037 at 1.4 GHz with a flux density of $S_{1.4} \sim 7.8$ mJy. The spectral slope is steep with $\alpha_{\text{in-band}} \sim 1.30$ and $\alpha_{1.4-5.5} \sim 0.97$. We added a Gaussian taper with a radius of 10 kλ (the longest baseline is 60 kλ) to increase the sensitivity of weak extended emission possibly connecting these two sources. However, no diffuse emission is seen between them in the tapered map.

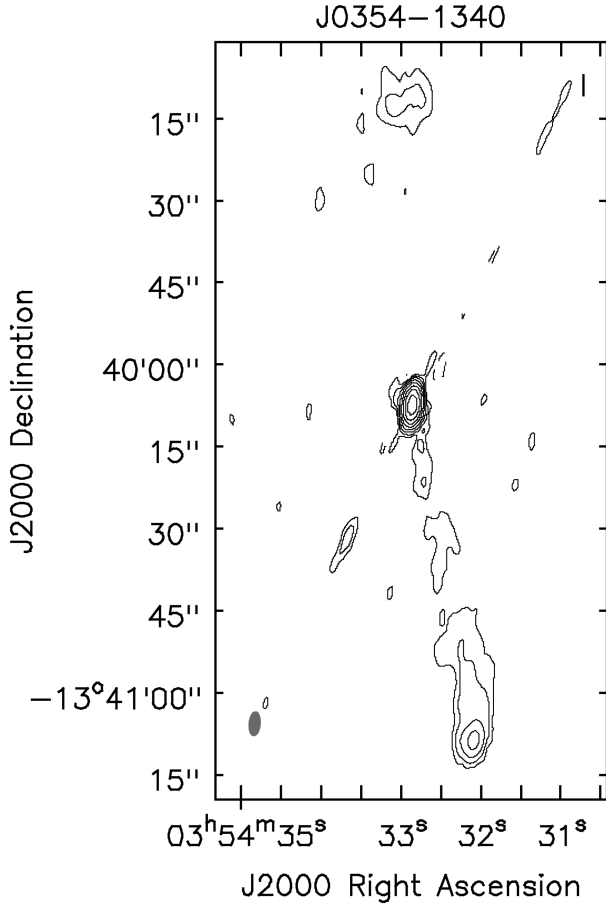


Figure A1. J0354–1340 with the VLA observation, rms = $8 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 7]$, beam size 7.38×3.47 kpc.

A2 J0350–1025

This source is the faintest one in the sample with a flux density of $S_{\text{int}} \sim 0.06$ mJy at redshift $z = 0.128$. It is unresolved at kpc-scale (Fig. B6, right-hand panel). The luminosity of $\log L_{\text{int}} \sim 38.27$ is below the average, while the black hole mass of $\log (M_{\text{BH}}/M_{\odot}) \sim 6.91$ is above the average. It has a very steep spectrum with $\alpha_{\text{in-band}} \sim 3.11$. This object was not detected at 1.4 GHz.

A nearby source to the west of the target, is associated with an optical counterpart of 2MASS J03505598–1025556 at redshift $z = 0.071$. We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C1, middle panel). It is weak at radio with a flux density of $S_{\text{int}} \sim 0.07$ mJy and a luminosity of $\log L_{\text{int}} \sim 37.66$. The spectrum is steep with $\alpha_{\text{in-band}} \sim 1.4$. No detection at 1.4 GHz is available.

A3 J0354–1340

This target has a very compact central core at kpc-scale with a flux density of $S_{\text{int}} \sim 5.09$ mJy. The VLA map is shown in Fig. A1. The luminosity and black hole mass are $\log L_{\text{int}} \sim 39.58$ and $\log (M_{\text{BH}}/M_{\odot}) \sim 6.99$, respectively, which are higher than the mean values of this sample. The in-band spectral slope is flat with $\alpha_{\text{in-band}} \sim 0.12$.

A southern component (centred at RA 03:54:32.09 and Dec. $-13:41:08.46$) has a separation of 62.2 arcsec from the central target. At redshift $z = 0.076$, the projected linear size is 93.5 kpc. The flux

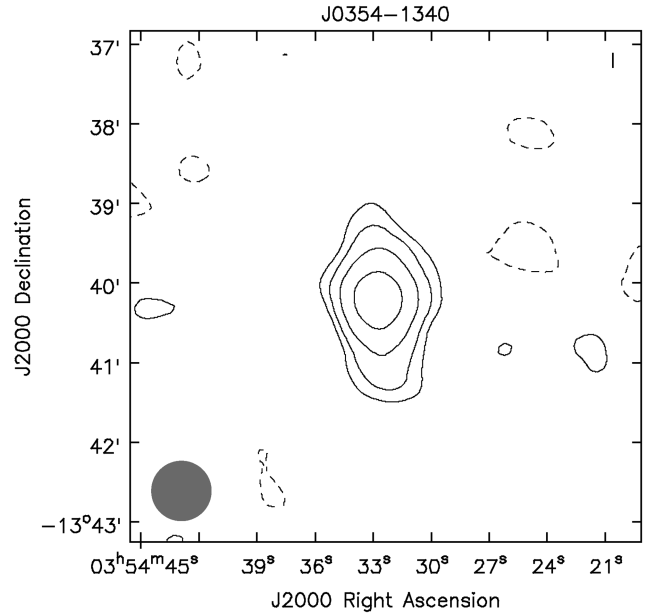


Figure A2. J0354–1340 with the NVSS detection, rms = $300 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 3]$, beam size 60.41×60.41 kpc.

density and luminosity are $S_{\text{int}} \sim 1.35$ mJy and $\log L_{\text{int}} \sim 39.05$, respectively, which are weaker than the central core. It shows diffuse emission. We measured a spectral index of $\alpha_{\text{in-band}} \sim 1.7$ that is steeper than the core.

A northern component (centred at RA 03:54:32.95 and Dec. $-13:39:11.57$) has a separation of 55.7 arcsec from the central target, corresponding to a projected linear size of 83.9 kpc. The flux density and luminosity are $S_{\text{int}} \sim 0.45$ mJy and $\log L_{\text{int}} \sim 38.43$, respectively, which are fainter and more diffuse than the southern one. We did not measure its spectral index because it is not detected in the 5 and 6 GHz images that have a higher RMS than the original one.

This source was also detected at 1.4 GHz associating with a counterpart of NVSS J035432–134012. It has a flux density of $S_{1.4} \sim 14.9$ mJy and a luminosity of $\log L_{1.4} \sim 39.46$. The NVSS map exhibits an elongated structure shown in Fig. A2. The southern and northern components are unresolved, thus the flux measurement includes the radio emission from both the core and the two components. The spectral index between 1.4 and 5.5 GHz was estimated using the NVSS flux and the VLA flux of a core plus two components giving $\alpha_{1.4-5.5} \sim 0.56$. We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C1, right-hand panel) and found that no optical counterpart is associated with the southern and northern components.

A4 J0436–1022 (Mrk 618)

Mrk 618 looks like a spiral galaxy on the optical image from the Pan-STARRS (Fig. C2, left-hand panel). The VLA map shows an extended morphology at kpc-scale (Fig. B9, right-hand panel) with a flux density of $S_{\text{int}} \sim 4.62$ mJy. The southern extended structure is twice as bright as the northern one. We overlaid the radio contours on the optical image and found that the morphology of the extended structure in the radio map is coincident with that of the host galaxy in the optical image. The spectral slope is steep with $\alpha_{\text{in-band}} \sim 0.60$. It has a luminosity of $\log L_{\text{int}} \sim 38.86$ and a black hole mass of $\log (M_{\text{BH}}/M_{\odot}) \sim 6.90$, which are at an intermediate level.

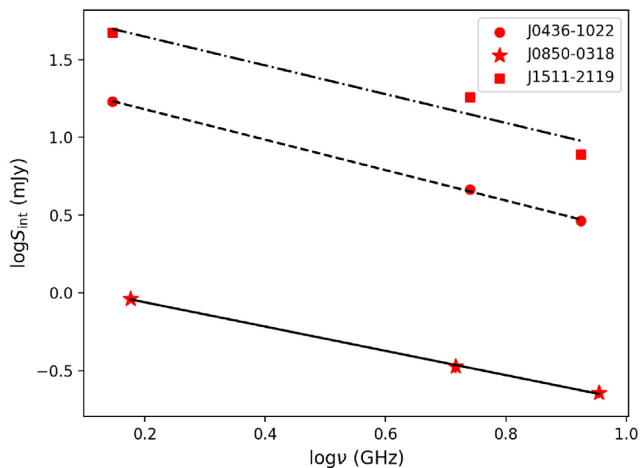


Figure A3. The radio spectra with a power-law fit of J0436–1022 (circles and dashed line), J0850–0318 (stars and solid line), and J1511–2119 (squares and dashdot line).

This object was detected with the NVSS at 1.4 GHz giving a flux density of $S_{1.4} \sim 17.0$ mJy, a luminosity of $\log L_{1.4} \sim 38.83$, and a steep spectral index of $\alpha_{1.4-5.5} \sim 0.95$. It was also observed with the VLA A-configuration at 8.4 GHz having a flux density of $S_{8.4} \sim 2.9$ mJy (Thean et al. 2000). The source displays a compact core in the 8.4 GHz map, where the northern and southern extended structure in the 5.5 GHz map is probably out-resolved at a high-resolution observation. We modelled the radio spectrum between 1.4 and 8.4 GHz with a power law yielding a steep spectral index of $\alpha_{1.4-8.4} \sim 0.98$, as shown in Fig. A3.

A5 J0447–0508

This source has a compact core at kpc-scale (Fig. B10, right-hand panel) with a flux density of $S_{\text{int}} \sim 4.04$ mJy. The luminosity is slightly higher than the average and the black hole mass is comparable to the average, having $\log L_{\text{int}} \sim 39.56$ and $\log (M_{\text{BH}}/M_{\odot}) \sim 6.85$, respectively. The spectral slope is steep with $\alpha_{\text{in-band}} \sim 0.83$. It was not detected at 1.4 GHz.

We note a very bright radio source to the north-east of the central target. It is a typical FR II radio galaxy according to the morphology. The central core is probably hidden. The radio emission is dominated by the radio lobes on both sides. Two hotspots are likely to be present at the end of the radio lobes. It has a flux density of $S_{\text{int}} \sim 22.70$ mJy and a steep slope of $\alpha_{\text{in-band}} \sim 0.66$. This radio galaxy has an associated counterpart of NVSS J044722–050750 at 1.4 GHz, with a flux density of $S_{1.4} \sim 68.4$ mJy and a steep spectral index of $\alpha_{1.4-5.5} \sim 0.81$. We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C2, middle panel). It could be associated with an optical counterpart of 2MASS J04472253–0507494, but no redshift is available.

A6 J0850–0318

This target shows a slightly extended structure at kpc-scale (Fig. B15, left-hand panel) with a flux density of $S_{\text{int}} \sim 0.43$ mJy. Its radio power is at a mean level with $\log L_{\text{int}} \sim 39.30$. But the black hole mass is large in this sample with $\log (M_{\text{BH}}/M_{\odot}) \sim 7.17$. It has a steep spectral index of $\alpha_{\text{in-band}} \sim 2.15$.

We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C2, right-hand panel), and note that a nearby source to

the west of the central target is associated with an optical counterpart of 2MASS J08502683–0318107. But no redshift is available. It has a flux density of $S_{\text{int}} \sim 0.37$ mJy and a very steep slope of $\alpha_{\text{in-band}} \sim 2.38$. To highlight the sensitivity of weak diffuse emission possibly connecting these two sources, we added a Gaussian taper with a radius of $30k\lambda$ (the longest baseline is $170k\lambda$). Even in the tapered map, the two sources do not appear to be connected. Besides, they were not detected with the 1.4 GHz surveys.

This target was also observed with the VLA A-configuration in L, C, and X bands (Proposal ID: VLA/19B-285) for the purpose of studying the steep in-band spectral index. The exposure time is 10 min on source in each band, yielding an average RMS of $30 \mu\text{Jy beam}^{-1}$. It exhibits a compact core in the A-array images at all frequencies. The flux densities are $S_{1.5} \sim 0.92$ mJy, $S_{5.2} \sim 0.34$ mJy, and $S_{9.0} \sim 0.23$ mJy. We modelled the radio spectrum from 1.5 to 9.0 GHz with a power law giving a simultaneous steep slope of $\alpha_{1.5-9.0} \sim 0.78$, as shown in Fig. A3. The nearby source was not detected with the A-array due to either the emission being out-resolved or the high RMS compared with the C-array observation.

A7 J0952–0136 (Mrk 1239)

Mrk 1239 is a well-studied object in literature. It exhibits a very compact core with the C-configuration at 5.5 GHz (Fig. B15, right-hand panel) and has a flux density of $S_{\text{int}} \sim 21.28$ mJy. The radio luminosity is at an intermediate level with $\log L_{\text{int}} \sim 39.01$. The black hole mass is relatively small with $\log (M_{\text{BH}}/M_{\odot}) \sim 6.33$, which is consistent with other measurements in the range of $\log (M_{\text{BH}}/M_{\odot}) \sim 5.89-7.02$ (Ryan et al. 2007; Berton et al. 2015). The radio spectrum is steep with $\alpha_{\text{in-band}} \sim 0.91$. This object was also observed with the A-configuration at 5 GHz giving similar measurements and showing extended emission (Berton et al. 2018).

Other frequency observations are available as well for this target. At 1.4 GHz, it has a flux density of $S_{1.4} \sim 59.8$ mJy with the FIRST detection, yielding a luminosity of $\log L_{1.4} \sim 38.86$ and a steep slope of $\alpha_{1.4-5.5} \sim 0.76$. In addition, a VLBA observation at 1.7 GHz gave a flux density of $S_{1.7} \sim 41.8$ mJy and provided an image exhibiting a resolved and linear extended structure (Doi et al. 2013, 2015). Another observation with the VLA A-configuration at 8.5 GHz also displayed a slightly elongate structure and gave a flux density of $S_{8.5} \sim 14.6$ mJy (Orienti & Prieto 2010; Doi et al. 2013). Furthermore, a radio spectrum was derived by Doi et al. (2013) with a VLA B-configuration quasi-simultaneous observation, finding spectral indexes of $\alpha_{1.4-5} \sim 0.52$ at 1.4–5 GHz, $\alpha_{5-8.5} \sim 0.94$ at 5–8.5 GHz, and $\alpha_{8.5-15} \sim 1.64$ at 8.5–15 GHz, which become steeper towards higher frequencies. Besides, this source was not detected at 22 GHz, which could be due to the steep spectral slope leading to the flux density at 22 GHz below the limit of the instrument (Doi et al. 2016).

A8 J1014–0418 (PG 1011–040)

This target is in the PG quasar sample (Boroson & Green 1992). It is a RQ source having a slightly extended structure at kpc-scale (Fig. B16, left-hand panel) and a flux density of $S_{\text{int}} \sim 0.38$ mJy. The radio power is relatively weak with a luminosity of $\log L_{\text{int}} \sim 38.23$. The black hole mass is $\log (M_{\text{BH}}/M_{\odot}) \sim 6.86$ which is at a mean level and in agreement with previous study (Davis & Laor 2011). The spectral slope is steep with $\alpha_{\text{in-band}} \sim 0.79$. It was not detected at 1.4 GHz. Additionally, this object was also observed with the A-array at 5 GHz giving a similar flux density and with the B-array at 8.5 GHz giving $S_{8.5} \sim 0.27$ mJy, which result in a steep slope of

$\alpha_{5-8.5} \sim 0.57$ between 5 and 8.5 GHz (Kellermann et al. 1989; Laor, Baldi & Behar 2019).

A9 J1511–2119

This source has a very compact core at kpc-scale (Fig. B20, left-hand panel) with a flux density of $S_{\text{int}} \sim 18.19$ mJy. It is a radio bright object with $\log L_{\text{int}} \sim 39.66$ above the average, but has a small black hole with $\log (M_{\text{BH}}/M_{\odot}) \sim 6.63$ below the average. It was detected at 1.4 GHz with an associated counterpart of NVSS J151159–211900. The flux density and luminosity are $S_{1.4} \sim 46.9$ mJy and $\log L_{1.4} \sim 39.48$, respectively. The spectral slope is steep with $\alpha_{\text{in-band}} \sim 0.81$ and $\alpha_{1.4-5.5} \sim 0.69$. An observation with the VLA A-configuration at 8.4 GHz gave a flux density of $S_{8.4} \sim 7.8$ mJy and a slightly extended morphology (Thean et al. 2000). We modelled the radio spectrum from 1.4 to 8.4 GHz with a power law obtaining a steep spectral index of $\alpha_{1.4-8.4} \sim 0.93$, as shown in Fig. A3.

A10 J1615–0936

This target is detected at radio frequencies for the first time, no previous detection with the 1.4 GHz surveys. The central core is slightly extended at kpc-scale (Fig. B21, left-hand panel) with a flux density of $S_{\text{int}} \sim 0.68$ mJy. The luminosity and black hole mass are below and above the averages having $\log L_{\text{int}} \sim 38.59$ and $\log (M_{\text{BH}}/M_{\odot}) \sim 6.94$, respectively. It has a steep radio spectrum with $\alpha_{\text{in-band}} \sim 1.15$.

A nearby component (centred at RA 16:15:18.66 and Dec. –09:36:12.75) to the west of the central core, has a separation of 5.9 arcsec corresponding to a projected linear size of 7.7 kpc at $z = 0.065$. We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C3, left-hand panel). No optical counterpart is associated with it. The flux density and luminosity are $S_{\text{int}} \sim 0.15$ mJy and $\log L_{\text{int}} \sim 37.91$, respectively, which are weaker than the central core. It shows slightly diffuse emission. The spectral slope is steeper than the core with $\alpha_{\text{in-band}} \sim 1.6$.

A11 J1937–0613

This target exhibits diffuse emission around the central core at kpc-scale (Fig. B23, right-hand panel) with a flux density of $S_{\text{int}} \sim 12.12$ mJy. The luminosity and black hole mass are $\log L_{\text{int}} \sim 38.20$ and $\log (M_{\text{BH}}/M_{\odot}) \sim 6.03$, respectively, which are lower than the average values. It was also detected at 1.4 GHz associating with a counterpart of NVSS J193733–061304. The flux density and luminosity are $S_{1.4} \sim 42.2$ mJy and $\log L_{1.4} \sim 38.15$, respectively. The spectrum is steep with $\alpha_{\text{in-band}} \sim 0.64$ and $\alpha_{1.4-5.5} \sim 0.91$. It looks like a spiral galaxy on the optical image from the Pan-STARRS (Fig. C3, middle panel). We overlaid the radio contours on the optical image and found that the morphology of the diffuse emission is similar to that of the host galaxy.

A12 J2021–2235

This target is one of the RL NLS1s in the sample with a flux density of $S_{\text{int}} \sim 9.47$ mJy. It is a radio bright source having $\log L_{\text{int}} \sim 40.70$, which is consistent with previous study (Komossa et al. 2006). The black hole mass is relatively small with $\log (M_{\text{BH}}/M_{\odot}) \sim 6.65$, which lies in the range of $\log (M_{\text{BH}}/M_{\odot}) \sim 6.48$ – 7.57 according to different methods (Komossa et al. 2006; Foschini et al. 2015). It has an associated counterpart of NVSS J202104–223520 at 1.4 GHz with a flux density of $S_{1.4} \sim 24.6$ mJy and a luminosity of $\log L_{1.4} \sim 40.50$. The spectral slope is steep with $\alpha_{\text{in-band}} \sim 0.95$ and $\alpha_{1.4-5.5} \sim 0.70$. This source exhibits a compact core at kpc-scale (Fig. B24, left-hand panel). An optical morphological study reveals that it is interacting with a nearby galaxy to the north-east of the target (Berton et al. 2019). We overlaid the radio contours on the optical image from the Pan-STARRS (Fig. C3, right-hand panel). The nearby galaxy cannot be resolved in the current VLA map.

APPENDIX B: RADIO MAPS

We present the radio maps of the southern NLS1s with the VLA C-configuration observations at 5.5 GHz in Figs B1–B24.

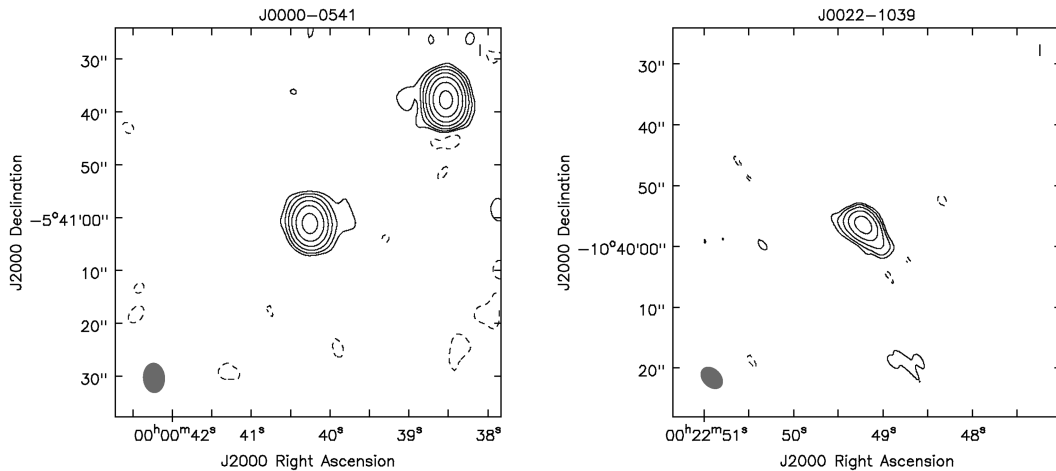


Figure B1. Left-hand panel: J0000–0541, $\text{rms} = 8 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 6]$, beam size 11.58×8.37 kpc. Right-hand panel: J0022–1039, $\text{rms} = 6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 4]$, beam size 44.67×31.72 kpc.

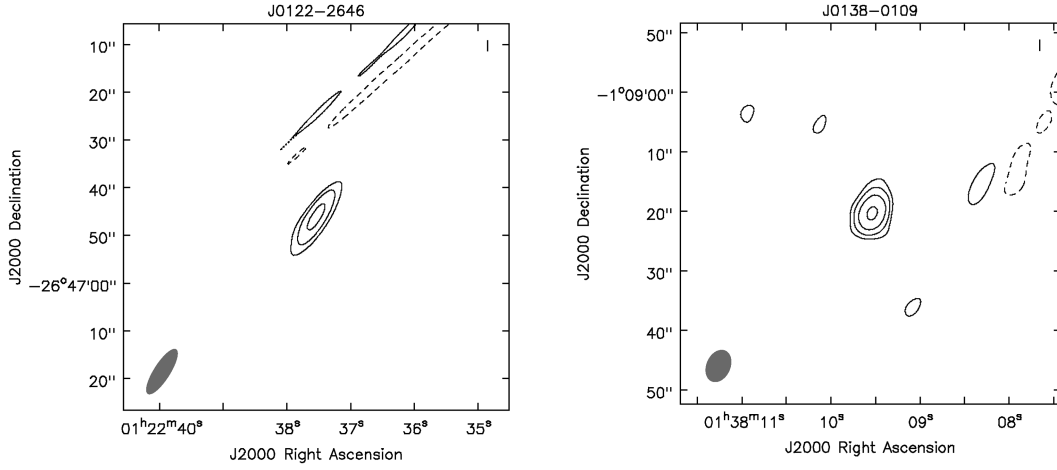


Figure B2. Left-hand panel: J0122-2646, rms = $55 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 120.29×35.55 kpc. Right-hand panel: J0138-0109, rms = $8 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 37.10×26.15 kpc.

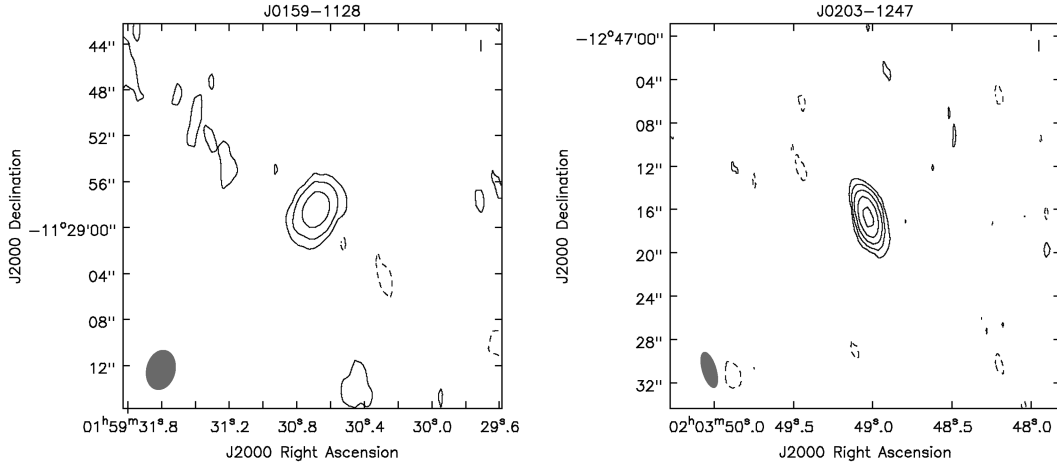


Figure B3. Left-hand panel: J0159-1128, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 12.92×9.22 kpc. Right-hand panel: J0203-1247, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 4]$, beam size 3.82×1.42 kpc.

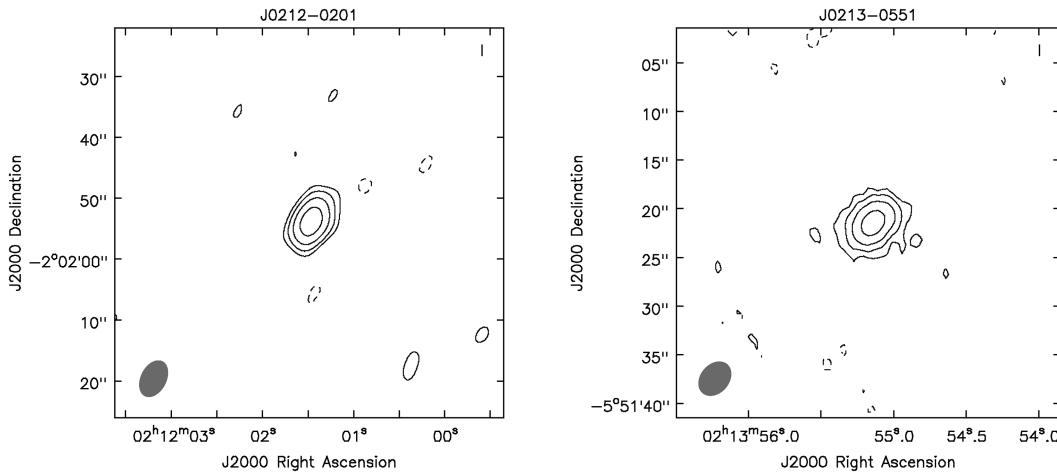


Figure B4. Left-hand panel: J0212-0201, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 73.33×47.79 kpc. Right-hand panel: J0213-0551, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 12.24×9.13 kpc.

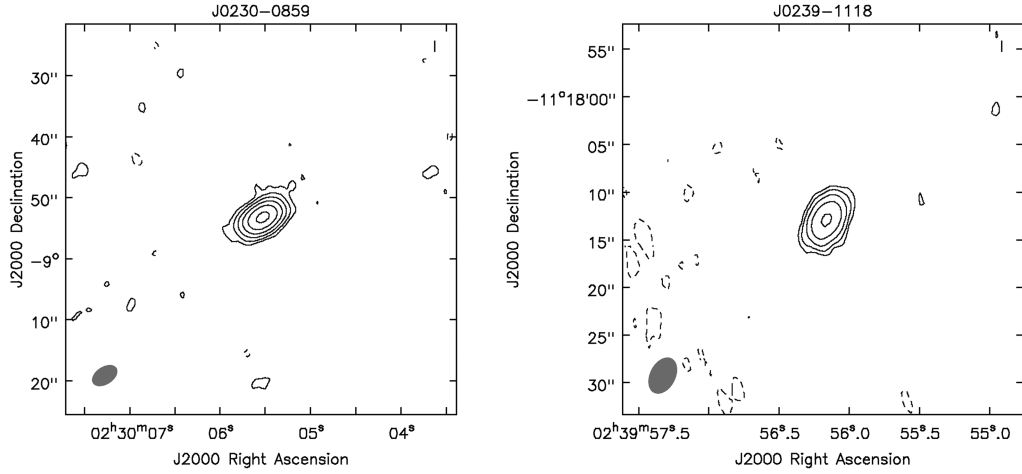


Figure B5. *Left-hand panel:* J0230–0859, $\text{rms} = 9 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 1.50×0.92 kpc. *Right-hand panel:* J0239–1118, $\text{rms} = 6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 4]$, beam size 19.13×12.45 kpc.

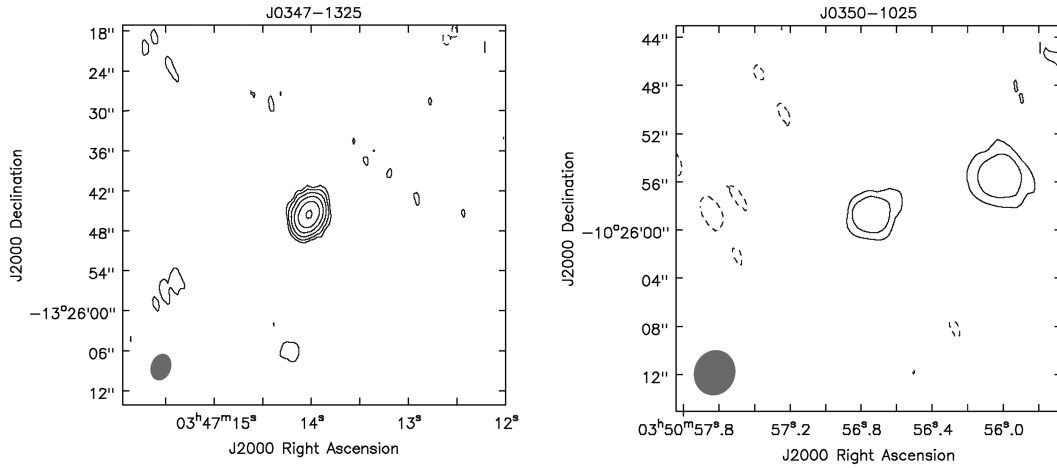


Figure B6. *Left-hand panel:* J0347–1325, $\text{rms} = 6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 17.82×13.10 kpc. *Right-hand panel:* J0350–1025, $\text{rms} = 5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 1]$, beam size 10.83×9.72 kpc.

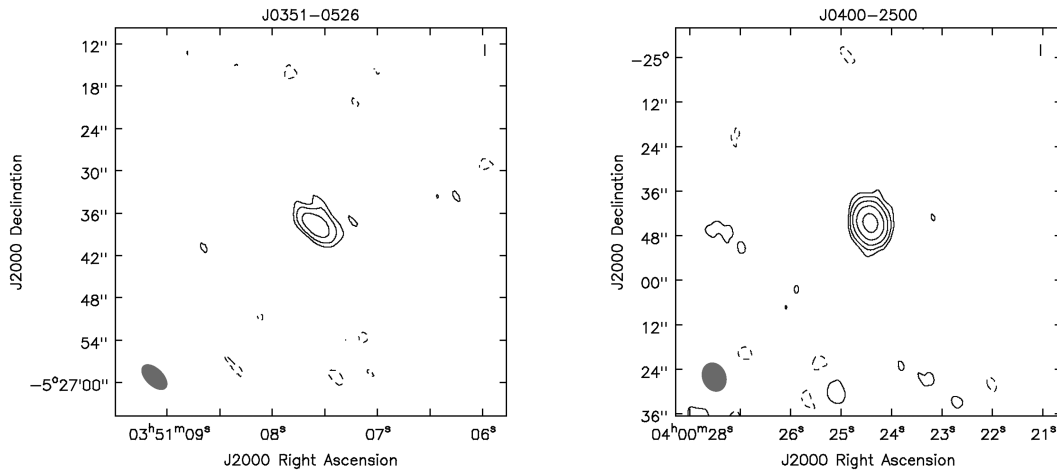


Figure B7. *Left-hand panel:* J0351–0526, $\text{rms} = 5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 6.51×3.52 kpc. *Right-hand panel:* J0400–2500, $\text{rms} = 19 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 4]$, beam size 17.04×13.46 kpc.

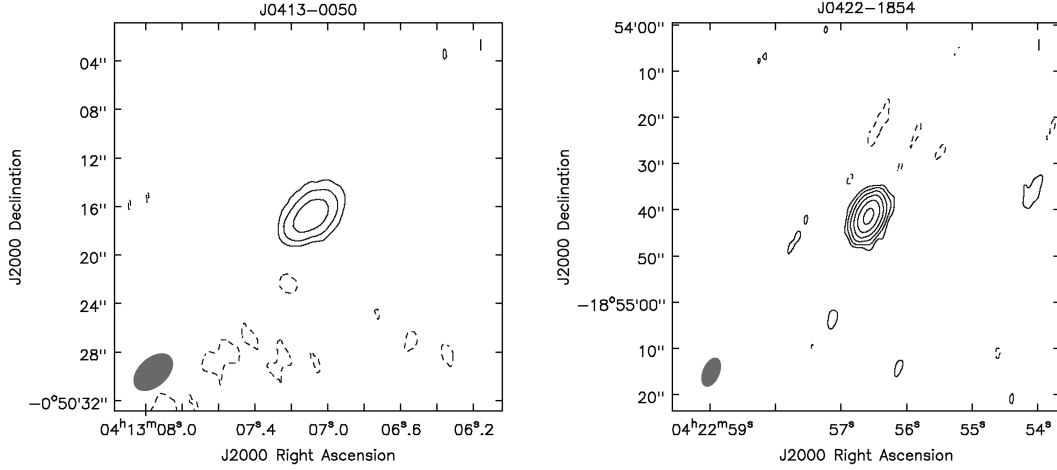


Figure B8. *Left-hand panel:* J0413–0050, rms = $7 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 3.22×1.98 kpc. *Right-hand panel:* J0422–1854, rms = $9 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 9.00×4.84 kpc.

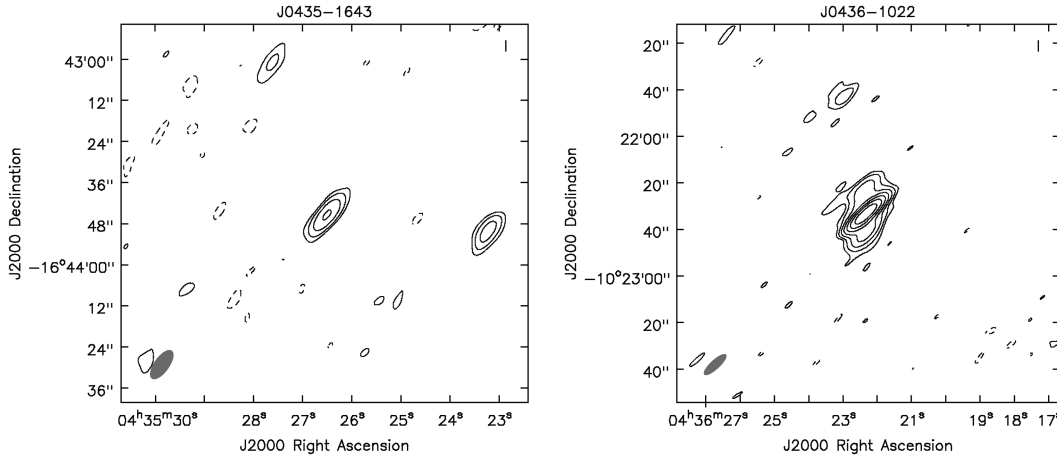


Figure B9. *Left-hand panel:* J0435–1643, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 21.33×9.02 kpc. *Right-hand panel:* J0436–1022, rms = $11 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 6]$, beam size 9.22×2.65 kpc.

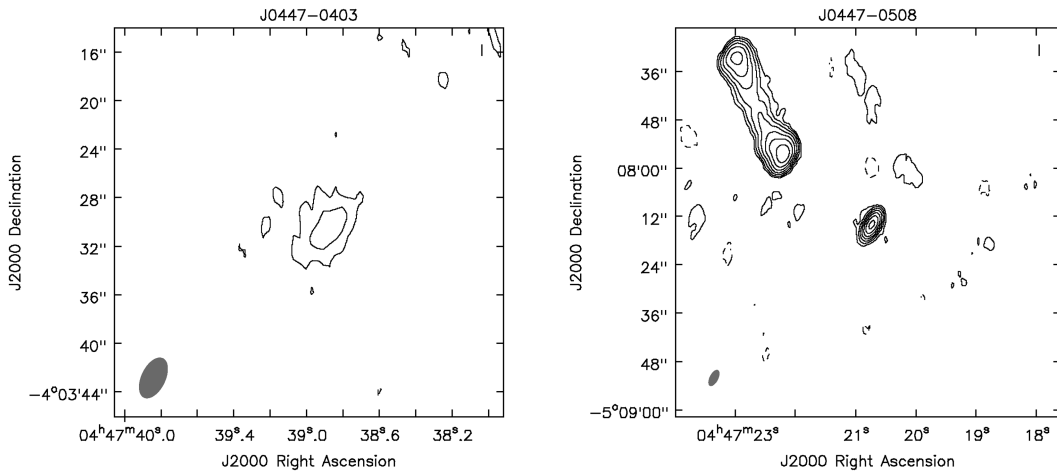


Figure B10. *Left-hand panel:* J0447–0403, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 1]$, beam size 3.39×1.84 kpc. *Right-hand panel:* J0447–0508, rms = $15 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 6]$, beam size 7.44×3.39 kpc.

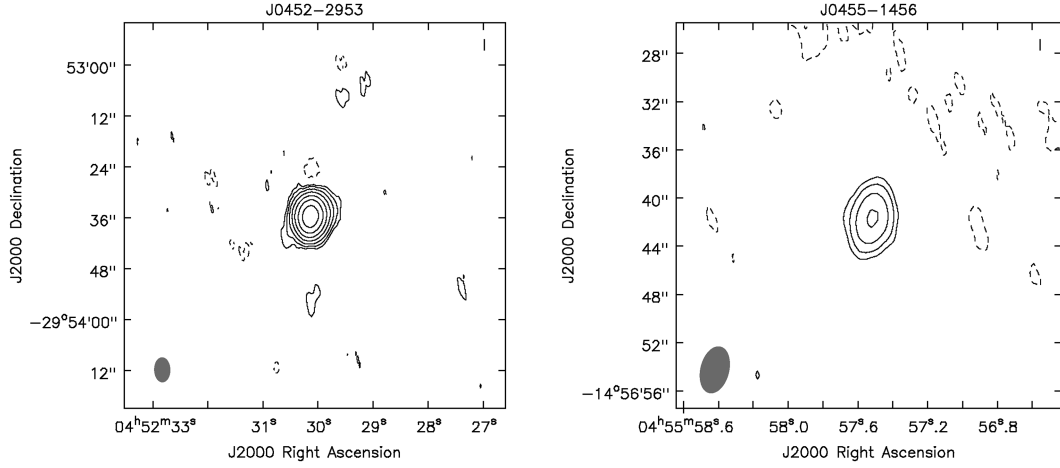


Figure B11. *Left-hand panel:* J0452–2953, rms = $8 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 6]$, beam size 40.37×25.70 kpc. *Right-hand panel:* J0455–1456, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 12.13×7.22 kpc.

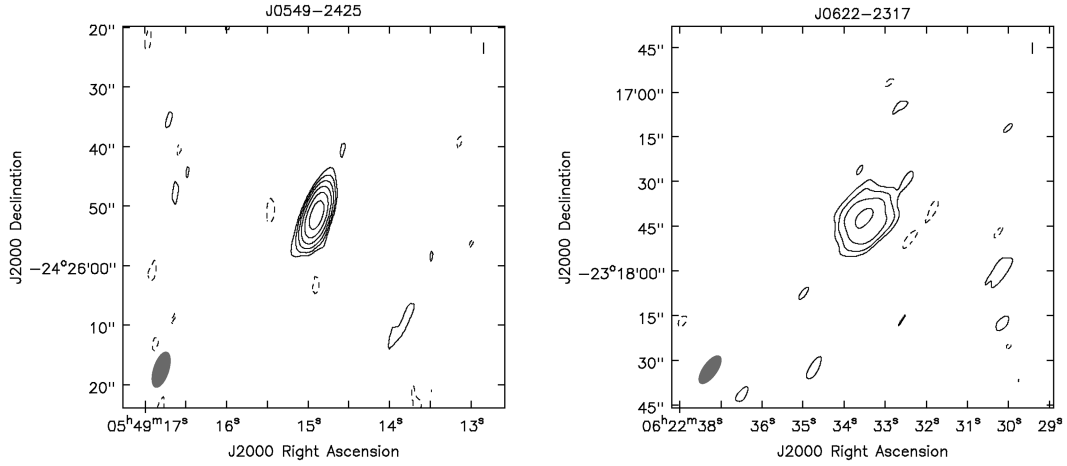


Figure B12. *Left-hand panel:* J0549–2425, rms = $10 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 5.95×2.43 kpc. *Right-hand panel:* J0622–2317, rms = $10 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 8.89×3.67 kpc.

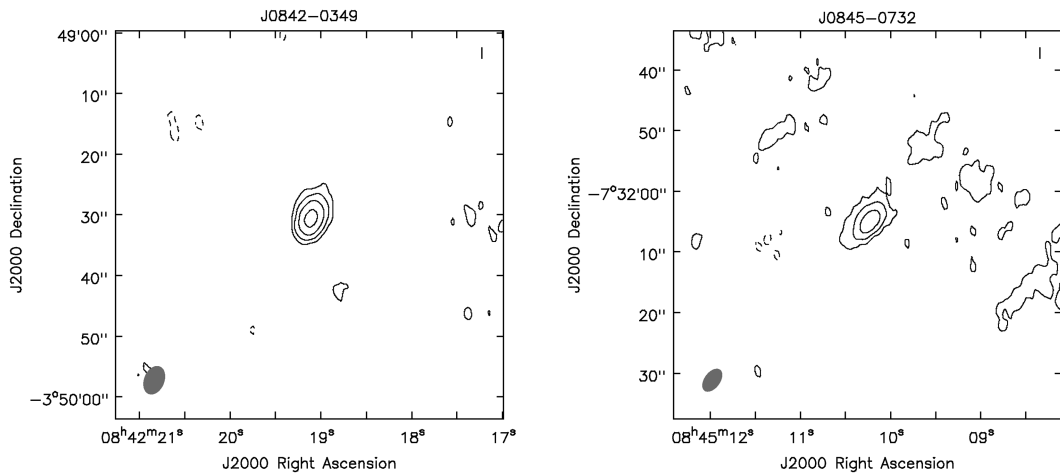


Figure B13. *Left-hand panel:* J0842–0349, rms = $9 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 43.52×29.16 kpc. *Right-hand panel:* J0845–0732, rms = $7 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 9.62×5.58 kpc.

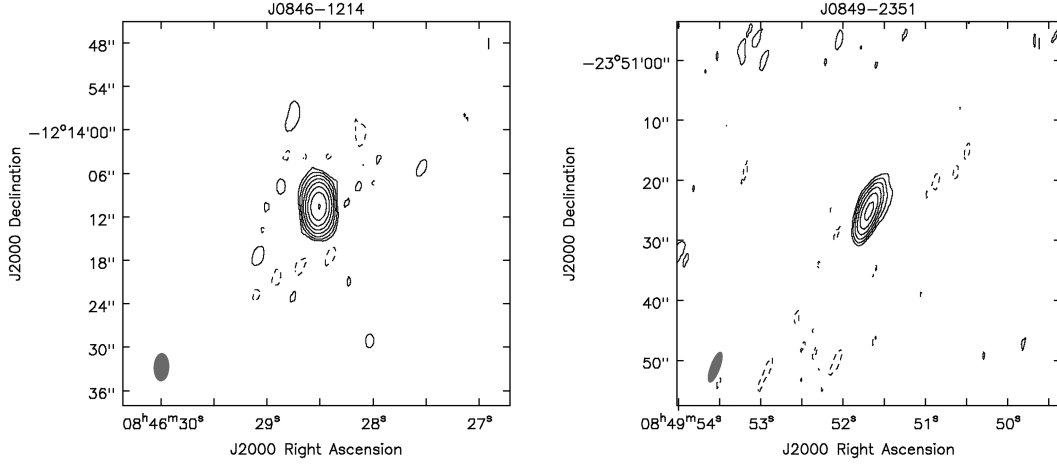


Figure B14. *Left-hand panel:* J0846–1214, rms = $13 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 7]$, beam size 9.14×4.93 kpc. *Right-hand panel:* J0849–2351, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 15.29×4.36 kpc.

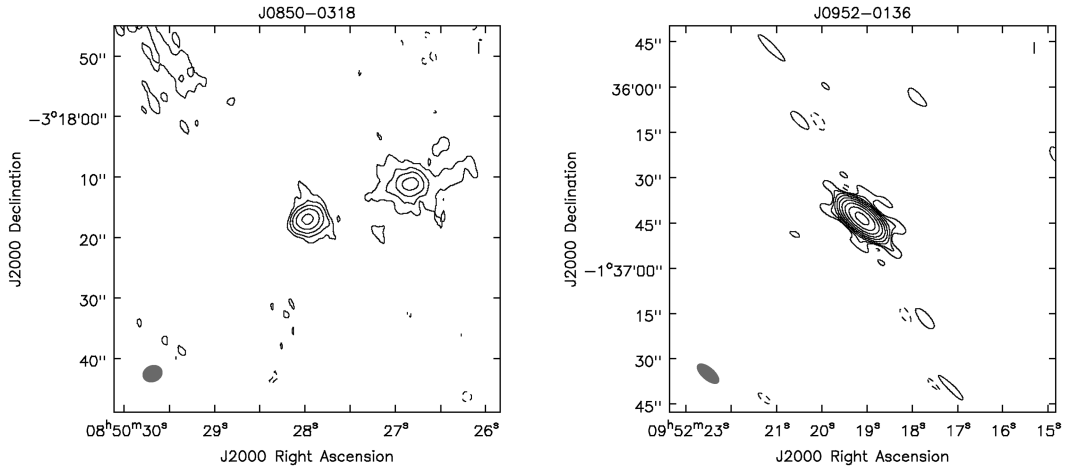


Figure B15. *Left-hand panel:* J0850–0318, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 4]$, beam size 12.25×9.91 kpc. *Right-hand panel:* J0952–0136, rms = $20 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 8]$, beam size 3.63×1.63 kpc.

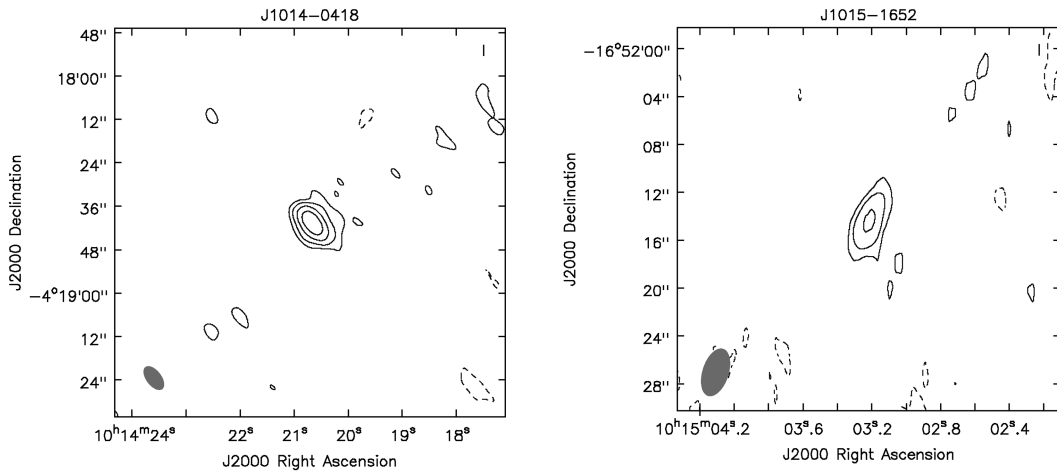


Figure B16. *Left-hand panel:* J1014–0418, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 9.34×4.91 kpc. *Right-hand panel:* J1015–1652, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 46.56×24.32 kpc.

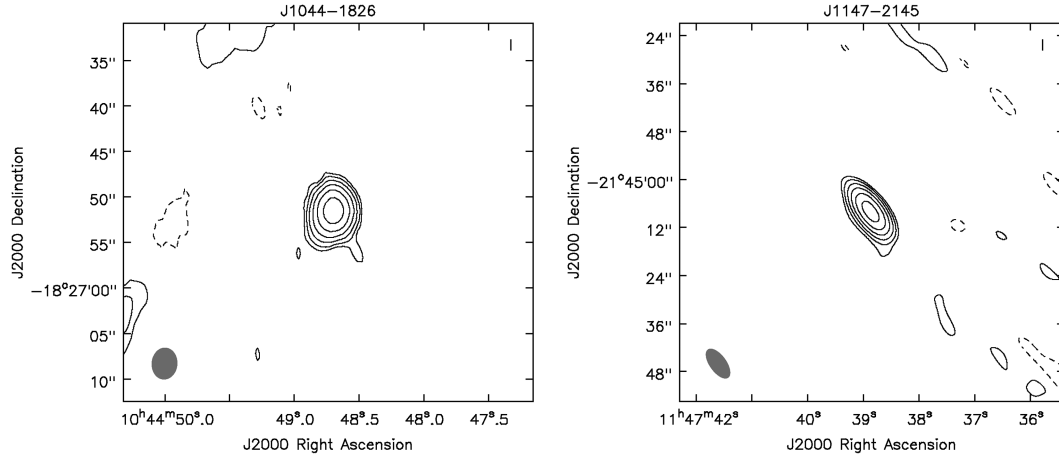


Figure B17. *Left-hand panel:* J1044–1826, rms = $10 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 8.66×6.96 kpc. *Right-hand panel:* J1147–2145, rms = $15 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 5]$, beam size 44.47×20.74 kpc.

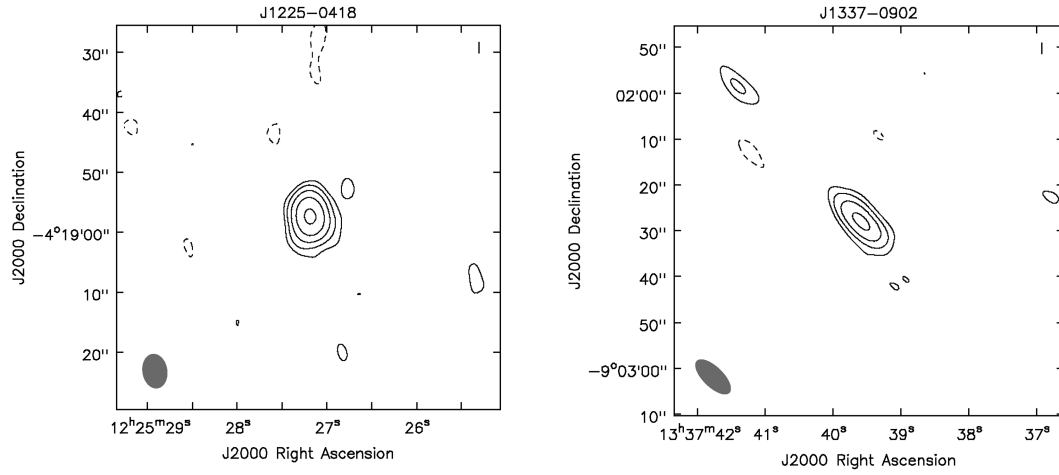


Figure B18. *Left-hand panel:* J1225–0418, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 4]$, beam size 17.65×12.54 kpc. *Right-hand panel:* J1337–0902, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 17.23×7.44 kpc.

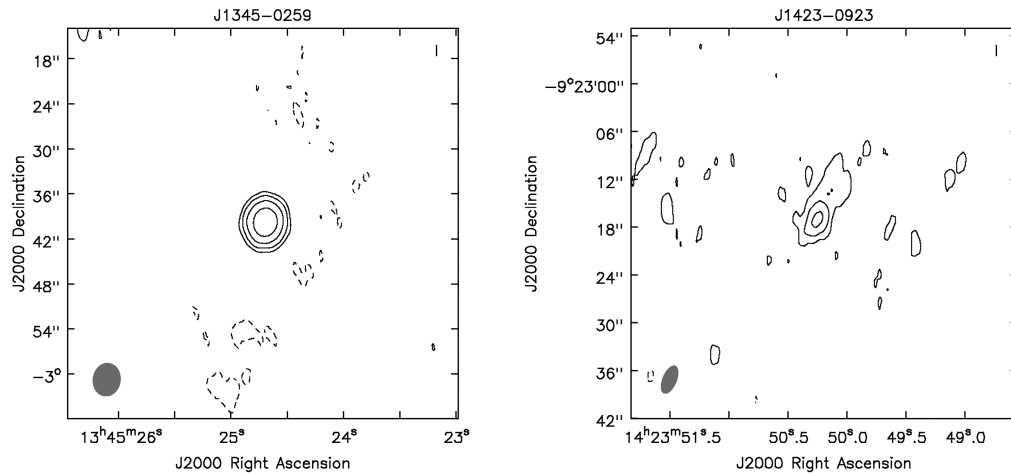


Figure B19. *Left-hand panel:* J1345–0259, rms = $9 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 3]$, beam size 8.20×6.78 kpc. *Right-hand panel:* J1423–0923, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 5.42×2.50 kpc.

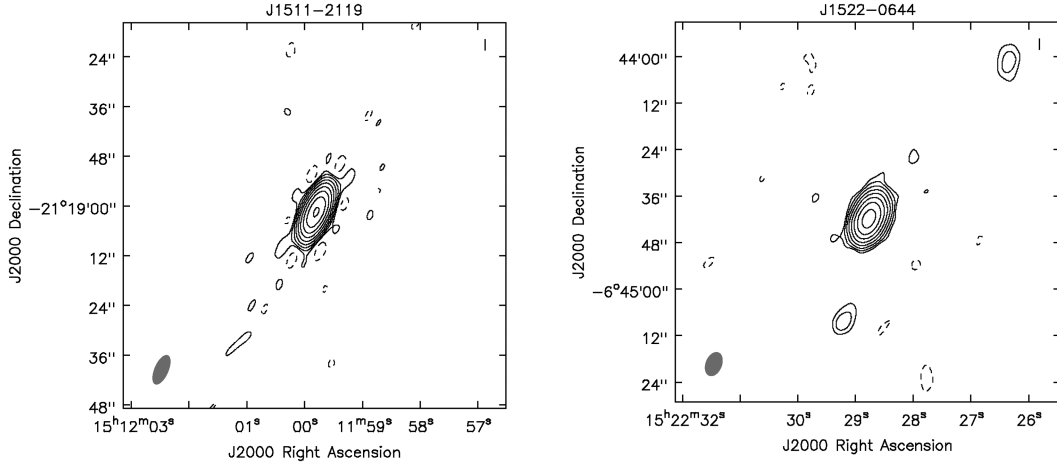


Figure B20. *Left-hand panel:* J1511–2119, rms = $21 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 8]$, beam size 7.07×3.02 kpc. *Right-hand panel:* J1522–0644, rms = $8 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 7]$, beam size 11.46×7.32 kpc.

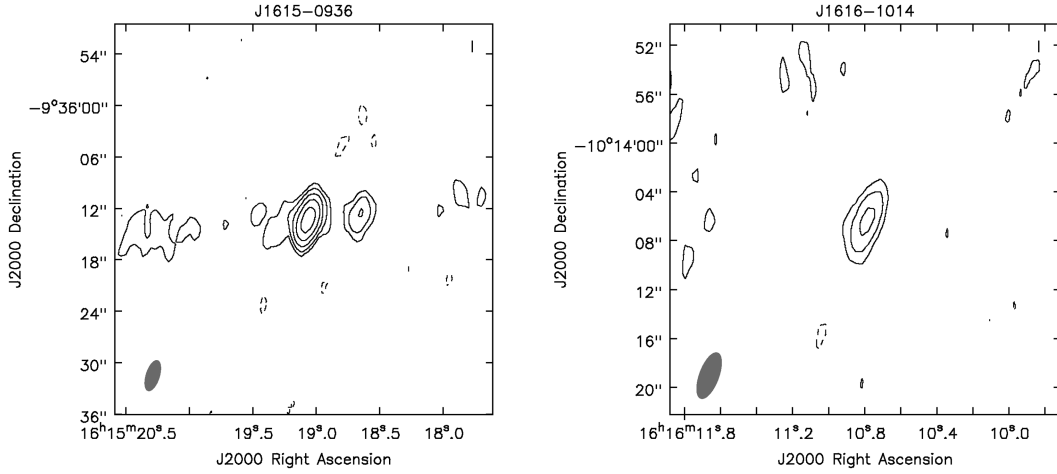


Figure B21. *Left-hand panel:* J1615–0936, rms = $7 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 4]$, beam size 5.17×2.17 kpc. *Right-hand panel:* J1616–1014, rms = $5 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 2]$, beam size 6.73×2.66 kpc.

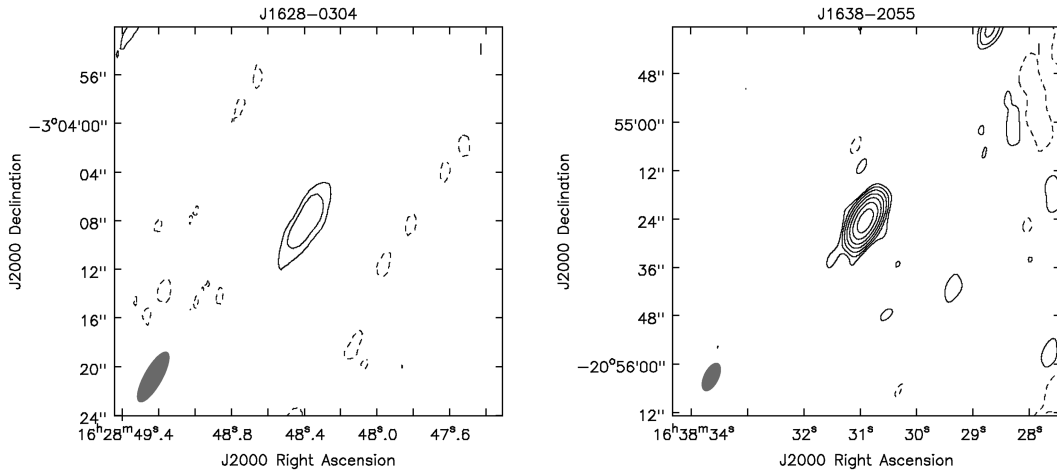


Figure B22. *Left-hand panel:* J1628–0304, rms = $6 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 1]$, beam size 9.58×2.99 kpc. *Right-hand panel:* J1638–2055, rms = $11 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n, n \in [0, 6]$, beam size 4.22×2.01 kpc.

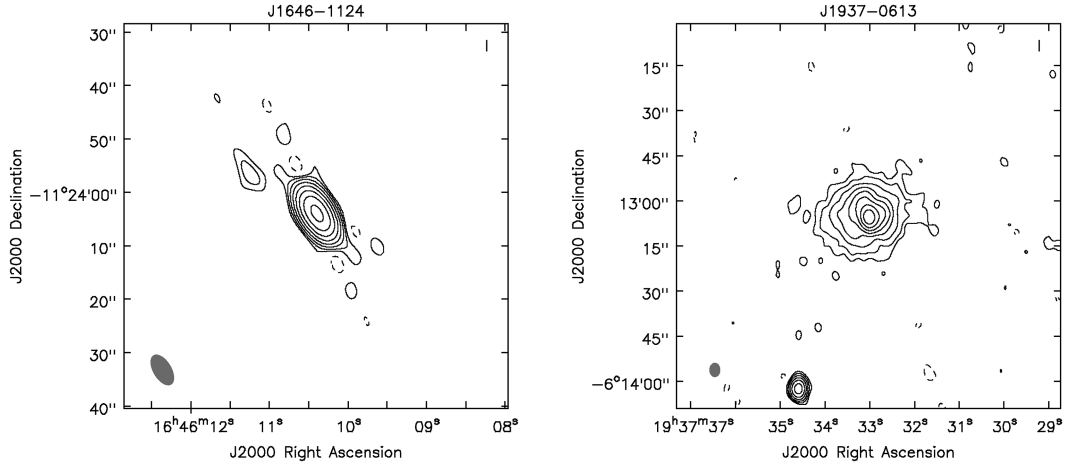


Figure B23. *Left-hand panel:* J1646–1124, rms = $22 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 7]$, beam size 10.09×5.26 kpc. *Right-hand panel:* J1937–0613, rms = $9 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 7]$, beam size 0.95×0.72 kpc.

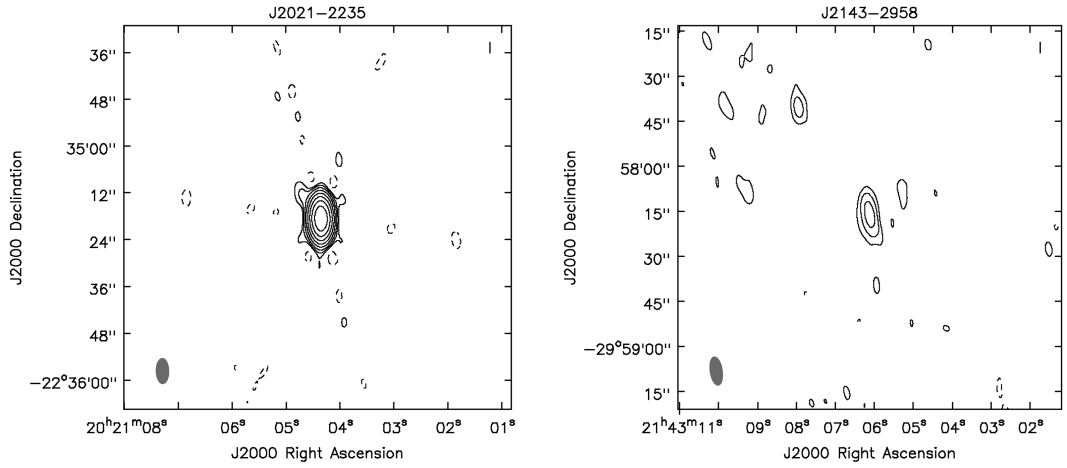


Figure B24. *Left-hand panel:* J2021–2235, rms = $12 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 7]$, beam size 28.12×14.02 kpc. *Right-hand panel:* J2143–2958, rms = $8 \mu\text{Jy beam}^{-1}$, contour levels at $-3, 3 \times 2^n$, $n \in [0, 2]$, beam size 25.77×10.94 kpc.

APPENDIX C: OPTICAL IMAGES OVERLAID WITH RADIO CONTOURS

We present the optical images from the Pan-STARRS overlaid with the radio contours for nine NLS1s in Figs C1–C3. The radio contours are the same as Figs A1 and Appendix B.

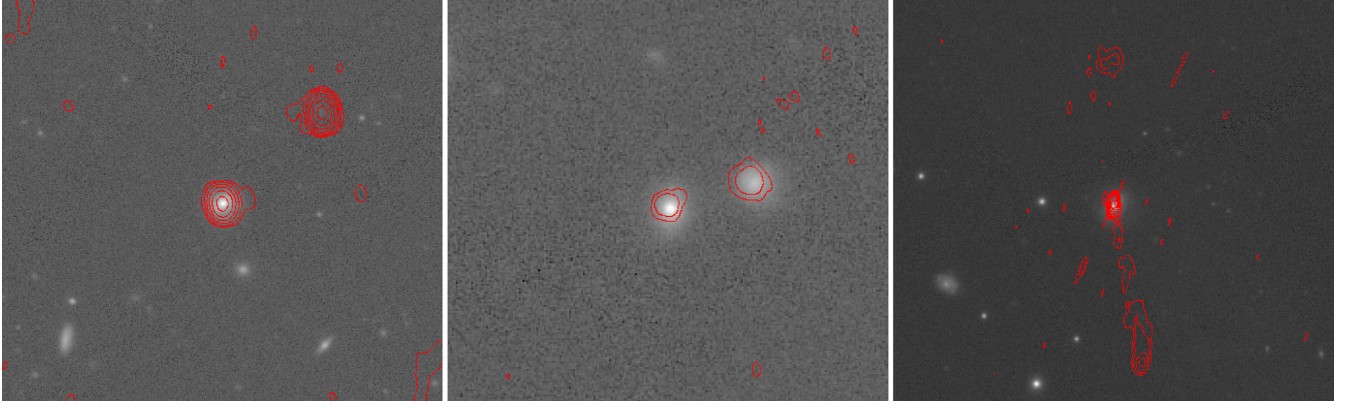


Figure C1. *Left:* J0000–0541, *middle:* J0350–1025, *right:* J0354–1340.

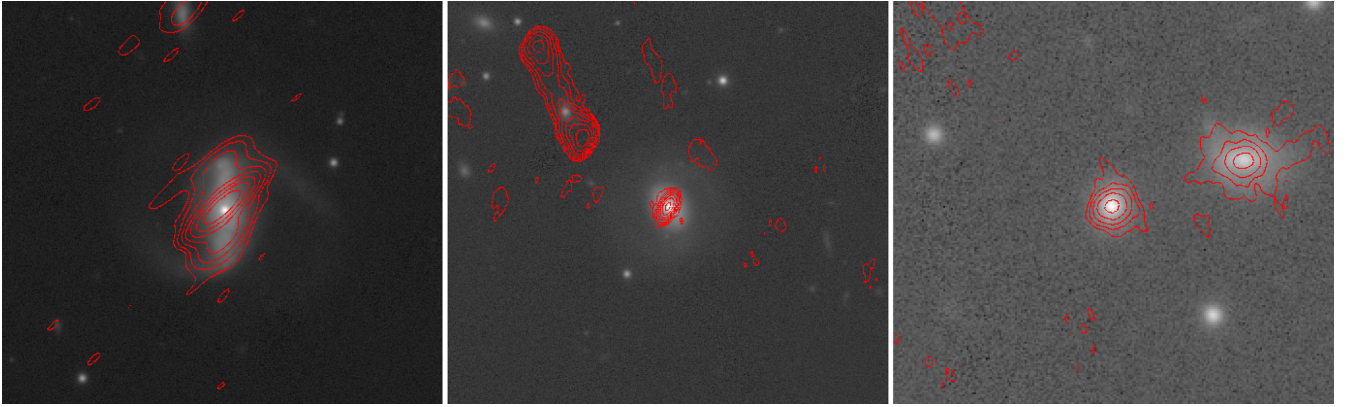


Figure C2. *Left:* J0436–1022, *middle:* J0447–0508, *right:* J0850–0318.

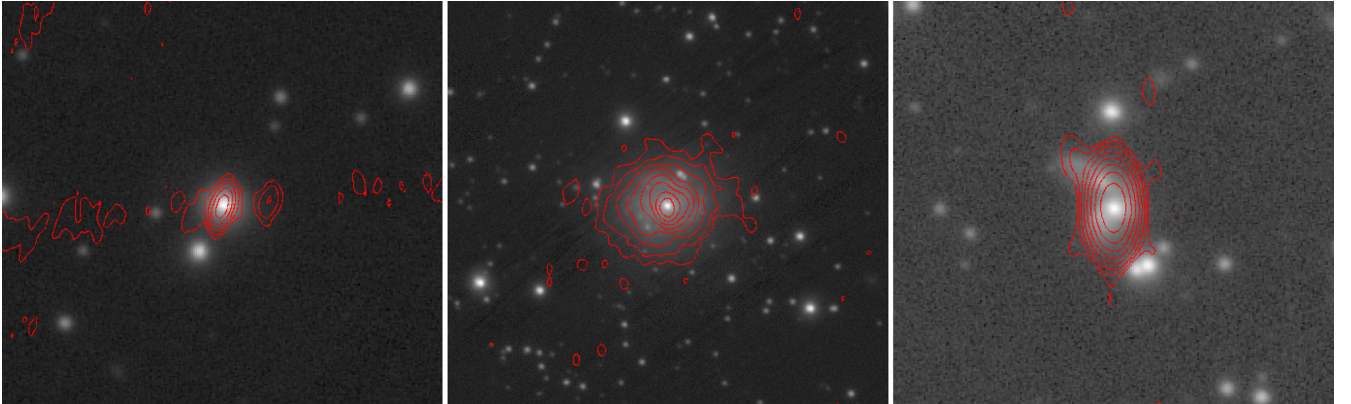


Figure C3. *Left:* J1615–0936, *middle:* J1937–0613, *right:* J2021–2235.

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