

INSPIRE: INvestigating Stellar Populations In RELics – IX. KiDS J0842 + 0059: the first fully confirmed relic beyond the local Universe

C. Tortora¹,¹★ G. Tozzi,² G. Agapito³,³ F. La Barbera,¹ C. Spiniello^{1,4},^{1,4}★ R. Li⁵,⁵ G. Carlà³,³ G. D’Ago⁶,⁶ E. Ghose,³ F. Mannucci³,³ N. R. Napolitano⁷,⁷ E. Pinna,³ M. Arnaboldi,⁸ D. Bevacqua^{9,10},^{9,10} A. Ferré-Mateu^{11,12},^{11,12} A. Gallazzi³,³ J. Hartke^{13,14},^{13,14} L. K. Hunt³,³ M. Maksymowicz-Maciata,^{4,15} C. Pulsoni³,³ P. Saracco⁹,⁹ D. Scognamiglio¹⁶ and M. Spavone¹

Affiliations are listed at the end of the paper

Accepted 2025 May 15. Received 2025 May 15; in original form 2025 April 2

ABSTRACT

Relics are massive, compact and quiescent galaxies that assembled the majority of their stars in the early Universe and lived untouched until today, completely missing any subsequent size growth caused by mergers and interactions. They provide the unique opportunity to put constraints on the first phase of mass assembly in the Universe with the ease of being nearby. While only a few relics have been found in the local Universe, the INSPIRE project has confirmed 38 relics at higher redshifts ($z \sim 0.2\text{--}0.4$), fully characterizing their integrated kinematics and stellar populations. However, given the very small sizes of these objects and the limitations imposed by the atmosphere, structural parameters inferred from ground-based optical imaging are possibly affected by systematic effects that are difficult to quantify. In this paper, we present the first high-resolution image obtained with Adaptive Optics K_s -band observations on SOUL-LUCI@LBT of one of the most extreme INSPIRE relics, KiDS J0842 + 0059 at $z \sim 0.3$. We confirm the discy morphology of this galaxy (axis ratio of 0.24) and its compact nature (circularized effective radius of ~ 1 kpc) by modelling its 2D surface brightness profile with a point-spread function-convolved Sérsic model. We demonstrate that the surface mass density profile of KiDS J0842 + 0059 closely resembles that of the most extreme local relic, NGC 1277, as well as of high-redshift red nuggets. We unambiguously conclude that this object is a remnant of a high-redshift compact and massive galaxy, which assembled all of its mass at $z > 2$, and completely missed the merger phase of the galaxy evolution at high redshift.

Key words: galaxies: clusters: general – galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: general – galaxies: structure.

1 INTRODUCTION

Relic galaxies are the final stage of high redshift passive, massive and compact galaxies, the red nuggets (Damjanov et al. 2009; Trujillo et al. 2009), that evolved undisturbed until the nearby Universe, missing the second phase of the so-called ‘two-phase’ scenario (Oser et al. 2010). According to this formation scenario, red nuggets form at high redshift through an intense and rapid series of dissipative processes. These processes produce a massive, passive, and very compact galaxy after star formation (SF) quenches (Labbé et al. 2023). Once this dissipative phase concludes (around $z \sim 2$, Zolotov et al. 2015), a second phase begins, characterized by gas inflows, dry mergers, and interactions. This accretion phase, which is significantly more extended in time than the first one, drives the dramatic structural evolution and size growth of early-type galaxies (ETGs) from $z \sim 2$ to the present day (Naab, Johansson & Ostriker 2009; Buitrago et al. 2018). However, since mergers are stochastic

events, a small percentage (1–10 per cent) of high-redshift red nuggets could miss completely the size-growth phase, and evolve undisturbed and passively in cosmic time (Hopkins et al. 2009; Quilis & Trujillo 2013). These nearby objects are the ancient relics of the early Universe.

In the local Universe, only few relic galaxies are known and have been studied in detail (e.g. Trujillo et al. 2014; Ferré-Mateu et al. 2015, 2017; Yıldırım et al. 2017; Beasley et al. 2018; Martín-Navarro, van de Ven & Yıldırım 2019; Salvador-Rusiñol et al. 2022; Comerón et al. 2023), pointing to an abundance consistent with theoretical expectations (e.g. Quilis & Trujillo 2013; Wellons et al. 2016; Flores-Freitas et al. 2022; Moura et al. 2024). Since these galaxies are expected to be slightly more common at higher redshifts, efforts have been dedicated to explore redshifts up to $z \sim 0.7$, using HST data (e.g. Damjanov et al. 2015), or wide-field ground-based images (e.g. Tortora et al. 2016, 2018, 2020; Charbonnier et al. 2017; Buitrago et al. 2018; Scognamiglio et al. 2020; Lisiecki et al. 2023). This has allowed the creation of a statistically large sample of relic candidates, i.e. ultra-compact massive galaxies (UCMGs; effective radius $R_e \lesssim 2$ kpc and stellar mass $M_* \gtrsim 6 \times 10^{10} M_\odot$) characterized by very red colours.

* E-mail: chiara.spiniello@physics.ox.ac.uk (CS); crescenzo.tortora@inaf.it (CT)

With the INvestigating Stellar Populations In RELics (INSPIRE) Project¹, we obtained the first systematic confirmation of relic galaxies beyond the local Universe. Thanks to high signal-to-noise, UVB-to-NIR X-Shooter (XSH) spectra, we obtained precise determinations of the stellar age, metallicity, [Mg/Fe], initial mass function (IMF) and velocity dispersion for the full sample of 52 UCMGs (Spiniello et al. 2021a, b, c, 2024; D’Ago et al. 2023; Martín-Navarro et al. 2023; Maksymowicz-Maciata et al. 2024; Scognamiglio et al. 2024). Out of these, 38 assembled more than 75 per cent of their mass by $z = 2$, and hence have been classified as relic galaxies. This has more than tripled the number of currently known relic galaxies, also pushing the redshift boundaries, and confirming the existence of a ‘degree of relicness’ (DoR; Ferré-Mateu et al. 2017) in terms of the star formation history (SFH). Indeed, from a detailed stellar population study of the final INSPIRE sample, in Spiniello et al. (2024) we have quantified the DoR as a dimensionless number, varying from 1 to 0, and defined in terms of (a) the fraction of stellar mass formed by $z = 2$, (b) the time at which a galaxy has assembled 75 per cent of its mass, and (b) the final assembly time. A high DoR (> 0.7) indicates a very early complete mass assembly, with almost no contribution from later star formation episodes. A lower DoR (≤ 0.34) means that there is a non-negligible percentage of populations with younger ages and different metallicities, and hence a much later time of final assembly.

INSPIRE has been the first important step to understand the physical processes which drove red nuggets formation at high redshift and their undisturbed evolution from $z \gtrsim 2$ until the local Universe, providing constraints on the *in situ* stellar populations in the central regions of the most massive and passive galaxies. However, both the imaging used to select the galaxies² and XSH spectroscopy to confirm these systems as relics are seeing dominated. Given the very small sizes of these objects on sky, compared to the ground-based FWHM seeing (i.e. effective radii ~ 0.2 – 0.5 arcsec versus FWHM ~ 0.7 arcsec in the best cases), the only viable way to securely estimate structural properties, constrain their morphology with high precision and confidence, and confirm the pureness of our selected sample is the use of Adaptive Optics (AO) supported imaging or space imaging data (FWHM ~ 0.1 arcsec).

In this paper, the 9th of the INSPIRE series, we present the first AO observations, obtained with the Single conjugated adaptive Optics Upgrade for LBT (SOUL) imager (Pinna et al. 2019) on the Large Binocular Telescope (LBT; Hill 2010), of one of the most extreme relic galaxies in our sample: KiDS J0842+0059. At the time of the proposal submission, KiDS J0842 + 0059 was the most promising among the INSPIRE galaxies that could be followed up with LBT. From KiDS seeing-limited data, we computed an effective radius of $R_e = 1$ kpc averaging the values obtained from g , r , and i images (Scognamiglio et al. 2020). Here, using the new AO observations, we unambiguously confirm the compact nature of KiDS J0842 + 0059, by determining more accurate size and structural parameters. This is the first time that a confirmed relic galaxy at these redshifts is observed with such high data quality and a spatial resolution of ~ 0.1 arcsec. Our study is the first step to confirm the truly compact nature of the INSPIRE sample.

The paper is organized as follows. In Section 2, we present the observations and the data analysis. Structural parameters are presented in Section 3, alongside with the target and its characteristics, which confirm the compact nature of the target. In Section 4, we compute

mass density profile of the target and compare it with these of other ultra-compact massive galaxies. A summary of the results is outlined in Section 5.

2 OBSERVATIONS, DATA REDUCTION, AND PSF MODELLING

The images analysed in this paper have been obtained on January 2021 with the SOUL imager that feeds the instrument LUCI (Heidt et al. 2018) on LBT, in the K_s filter. These observations are part of the proposal INVEstigating Relic galaxies with SOUL (INVERSO, P.I.: C. Tortora). The pixel scale of the imager is of 0.015 arcsec/pix.

2.1 Observations

The single-conjugate (SCAO) systems mounted on LBT use a natural guide star (NGS) to drive the deformable mirror. Under median seeing conditions, Strehl ratios (SR) above 0.1 can be obtained in the K_s band for NGS (AO star in the following) brighter than $R = 15$ mag and separations below 40 arcsec, limiting the number of observable targets (Pinna et al. 2023). Furthermore, bright closer stars are necessary to carefully map the point-spread function (PSF) of the instrument. The further away the target is from the AO star, the worse is the modelling of the PSF and the larger is the distortion of the image along the direction of the star. Hence, the ideal strategy when using AO systems is to use PSF stars at a distance from the AO star comparable to that of the target.

Only four objects from the current INSPIRE sample have a star in their proximity, bright enough to be used as AO star. At the time of the proposal submission, KiDS J0842 + 0059 was the most interesting of the four targets with a star in their proximity. Indeed, this is one of the most extreme relics of the sample, having formed more than 98 per cent of its stellar mass within ~ 3 Gyr after the big bang, and has three nearby bright stars to be used for PSF modelling. However, unfortunately these stars are further away from the AO star than the relic galaxy. In particular, the AO star is located at a distance of 23.2 arcsec from the target. It has a *Gaia*-band magnitude of 15.1 mag (KiDS g - and r -band of 15.5 mag and 15.1 mag, respectively, and 2MASS K band of 13.8 mag). The three PSF stars are instead at the distances of 42.3 ($K = 15$ mag), 55.4 ($K = 15.4$ mag), and 62.9 ($K = 14.6$ mag) arcsec from the AO star.

The LUCI observations consist of two data sets relative to two different sky pointings: one centred on the target galaxy and the AO star, and the other one containing the other stars used to estimate the PSF, as described below. In each field, the main target (i.e. either the galaxy or the stars) was dithered on the field of view, hence observed in different positions on the LUCI detector.³ For each dithered target position, an observing shot (or frame) of $NDIT = 12$ exposures of 10 s each was carried out. We observed 27 frames for the pointing centered on the target and the AO star, and 2 frames for the pointing containing the three PSF stars.

2.2 Data reduction

We reduced and combined the data frames of the galaxy and, separately, of the PSF stars, by means of our custom-made PYTHON routine (PySNAP), which extensively makes use of functions from

¹Based on the ESO Large Program ID: 1104.B-0370, PI: C. Spiniello.

²Multiband images from the Kilo Degree Survey (de Jong et al. 2015, 2017).

³The dithering has been performed with shifts in the range 2–3.5 arcsec from the reference position of the target along two orthogonal axes.

the `ccdproc` package⁴ included in the `astropy` library. Our procedure performs sky subtraction, accounts for the presence of bad pixels, and applies dark and flat field corrections. Being all on-target exposures, we produced a master sky frame using the `combine` function: all dithered on-target frames were combined by calculating the median value pixel-by-pixel, and rejecting the most deviating pixels (those beyond 2σ) with a sigma clipping technique. A flat field image was created with the same combining method as for the sky frame. We then corrected for dark using a dedicated dark exposure and finally divided the combined image for the median values of its pixels. Finally, we used the resulting flat field frame as input of the `ccdmask` function to generate the bad pixel mask.

With these ingredients, for each 12-exposure shot we generated the mean frame, in order to increase the signal-to-noise ratio (S/N). All corrections were then applied to each mean frame, namely the master-sky subtraction and the corrections for dark, flat field, and bad pixels. At this point, the different dithered, corrected mean frames were aligned first using the World Coordinate System (WCS) offsets stored in the header file, then applying an additional shift correction by fitting a 2D elliptical Moffat profile to the target (i.e. either the galaxy or the stars) surface brightness (SB). Although better suitable to reproduce the SB of point-like sources (e.g. stars) and not of extended sources, such as our galaxy, a 2D Moffat profile is, however, successful in identifying the galaxy emission centroid, necessary to align the distinct mean frames. Out of the 27 frame sequences where the galaxy appears, we discarded three sequences of clearly much worse quality, likely obtained in either bad atmospheric conditions or during instrumental issues (e.g. open loop of the AO system). The poor quality of these frames is confirmed by the larger FWHM (i.e. $\gtrsim 10$ pixels) of the 2D Moffat profile fitted to the AO star, present in all these frames, compared to the good galaxy frames where the AO star is also covered (FWHM ~ 5 pixels).⁵ For the galaxy, three frames were rejected, for the pointing containing the three PSF stars all frames were used. Finally, as last step, we combined all non rejected frames for each pointing. The final combined images for each pointing were created as an average frame of the corrected, aligned frames, using the `combine` function. In details, for the pointing including the target, we observed a total of 27 frames, and we combined 24 of these with optimal FWHM, totalling an exposure time of 2880s. For the pointing including the PSF stars, we combined 2 frames of 12×10 s each, hence reaching a total exposure time of 240s.

Along with the final combined image, we produced a weight map reporting pixel-by-pixel the value of $1/\sigma^2$, where σ is the uncertainty on photon counts in a given pixel. Since in the near-IR the dominant uncertainty is due to the Poissonian noise, the weight map was created in the following way. We first associated a single noise value σ_i to each corrected frame i , calculated as median absolute distance (MAD) value among those pixels not contaminated by source emission. Then, the value of σ in a given pixel (x, y) of the final combined image was computed as:

$$\sigma(x, y) = \sqrt{\frac{\sum_{i=1}^N \sigma_i^2(x, y)}{N(x, y)}}, \quad (1)$$

⁴All functions mentioned in the text belong to this package.

⁵Unfortunately, the AO star was not visible in all galaxy frames. As a result, we could neither use it to align the different frame sequences nor to thoroughly assess the quality of each sequence.

where N is the total number of combined frames in the pixel (x, y) , considering both the overlapping of different dithered frames and masked bad pixels.

The KiDS@VST and SOUL@LBT field of views, oriented with East and North along the horizontal and vertical directions, and including target, AO star, and PSF stars, are shown in Fig. 1. The figure highlights the much higher resolution we achieved with AO-supported images.

2.3 PSF modelling

Since the PSF FWHM varies with distance from the target, we conducted a detailed PSF modelling to accurately determine the PSF FWHM at the target's location and assess whether the planned objectives in terms of FWHM have been met. This is crucial to assess the possible uncertainties and systematics that could affect the determination of the structural parameters of the galaxy, which is the main goal of this paper.

Our approach consisted of two steps. First, we fitted the four stars (AO and PSF stars) using *TIPTOP* (Neichel et al. 2021), and then we extrapolated the PSF to the galaxy position. Technical details are provided in Appendix A.

The circularized FWHM of the three PSFs is of ~ 0.20 , ~ 0.24 , and ~ 0.25 arcsec, calculated using, independently, *TIPTOP* or (equivalently) fitting a Moffat profile to the PSF star 1, 2, and 3, respectively. The extrapolated FWHM value to the target position is of ~ 0.14 arcsec, with a major axes value of 0.168 arcsec and minor-axes of 0.116 arcsec. This is in line with the target value planned during the proposal preparation using the SOUL SR calculator.⁶

3 KIDS J0842 + 0059: A FULLY CONFIRMED RELIC BEYOND THE LOCAL UNIVERSE

In this section, we first provide a detailed discussion of the characteristics of KiDS J0842 + 0059, derived from past seeing-limited observations from KiDS and INSPIRE. We then complement this analysis with the structural parameter analysis based on this new observations from LUCI-SOUL@LBT, which confirm its ultra-compact nature.

3.1 Physical and stellar population properties

KiDS J0842 + 0059 is one of the most extreme relic galaxies in the INSPIRE sample. This galaxy, at $z = 0.2959$, has formed ~ 75 per cent of its stellar mass within the first Gyr after the big bang and fully completed its assembly 3.5 Gyr after it. Within the INSPIRE project, galaxies have been operationally classified in terms of their DoR into three broad families: non-relics (DoR < 0.34, 14 objects), relics ($0.34 \leq \text{DoR} \leq 0.7$, 29 objects), and extreme relics (DoR > 0.7, 9 objects). The DoR ranges between 0.06 and 0.83 with a median of 0.4. KiDS J0842 + 0059 has a DoR of 0.73, ranking as the seventh most extreme relic galaxy in the INSPIRE sample.

KiDS J0842 + 0059 has a stellar mass⁷ of $M_* \sim 10^{11} M_\odot$, inferred from spectral energy distribution (SED) fitting in the KiDS *ugri* bands using LE PHARE (see more details in Tortora et al. 2018; Scognamiglio et al. 2020). In Spiniello et al. (2024), we infer an integrated stellar velocity dispersion of $\sigma_e \sim 324 \pm 32 \text{ km s}^{-1}$, from a XSH spectra extracted over an aperture encapsulating 50 per cent

⁶<http://adopt.arcetri.inaf.it/strehl.html>

⁷The error on the stellar mass is of the order of 0.1 dex.

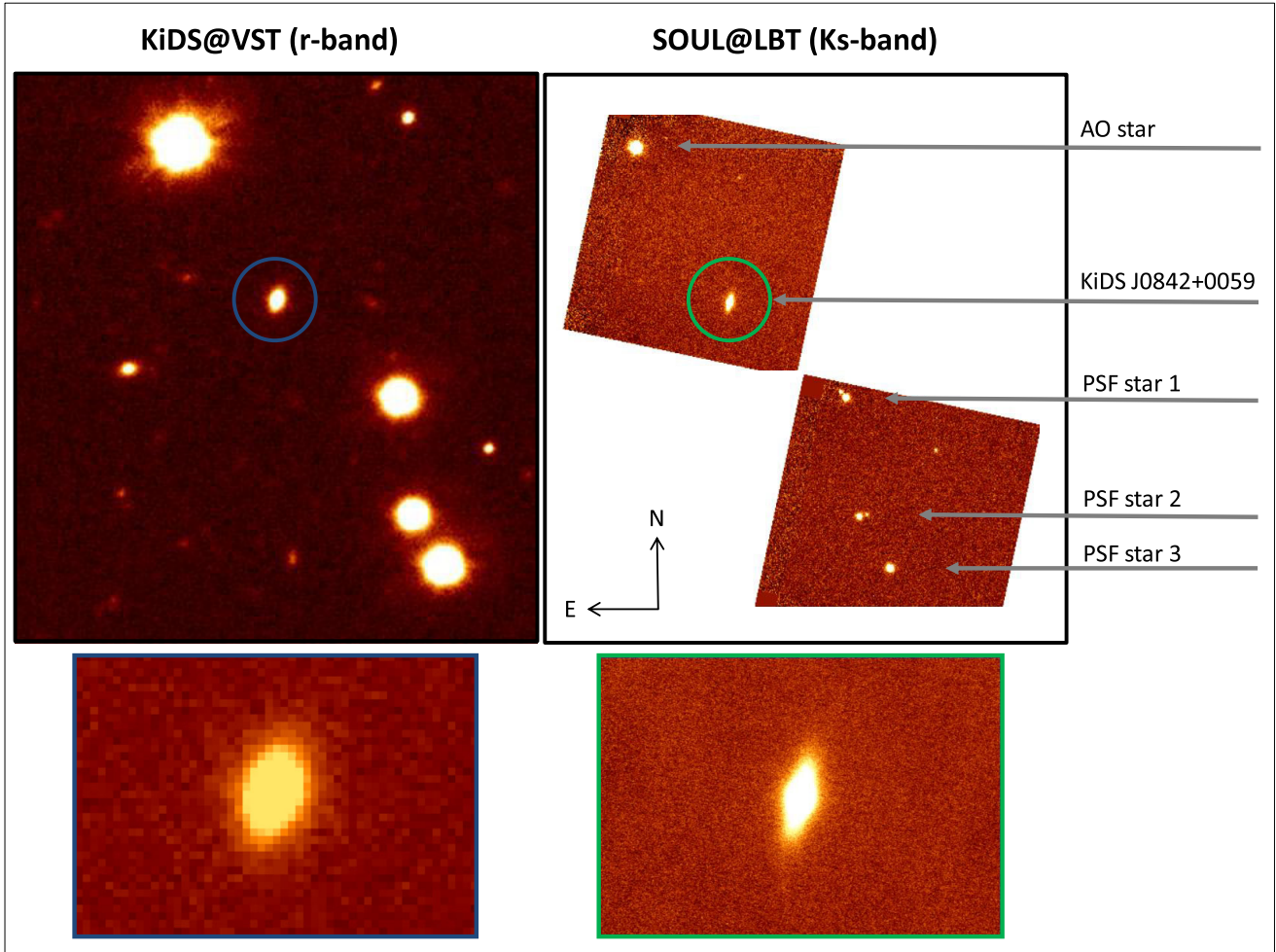


Figure 1. *Top.* Field of view of $\sim 1.1 \times 1.2 \text{ arcmin}^2$ including the target (within central circles), the AO star, and the three stars used to determine the PSF of the SOUL data. We show the KiDS@VST r -band image on the left panels and a SOUL K_s -band mosaic built merging the pointing on the AO star and the target and that containing the PSF stars on the right. Each SOUL pointing has a size of $\sim 30 \times 30 \text{ arcsec}^2$. Both images are oriented with East and North along the horizontal and vertical directions. *Bottom.* Cutouts of $\sim 8 \times 11 \text{ arcsec}^2$ from KiDS (left) and SOUL (right) centered on KiDS J0842 + 0059.

of the light (R50).⁸ Like for the other extreme relics, the σ_e is larger than the average values computed for normal-sized galaxies and non-relics with similar stellar masses (see fig. 9 in Spiniello et al. 2024).

The stellar population properties of KiDS J0842 + 0059 fully resemble those of local relics: it has a very old mass-weighted age of $\sim 10 \text{ Gyr}$, compared to the age of the Universe at $z = 0.3$ (Spiniello et al. 2024), supersolar metallicity ($[M/H] = 0.36_{-0.04}^{+0.03}$) and $[Mg/Fe] (0.26_{-0.06}^{+0.06})$, and a dwarf-rich IMF with a low-mass end slope of $\Gamma_b = 2.80_{-0.32}^{+0.26}$ (Maksymowicz-Maciata et al. 2024). When instead assuming a universal, Kroupa-like IMF (as in Spiniello et al. 2024), the galaxy resulted equally old ($\sim 10 \text{ Gyr}$), even more Mg-enhanced ($[Mg/Fe] \sim 0.4$) and still metal rich, although less super solar than when changing the IMF slope ($[M/H] \sim 0.25$). These stellar population estimates allow us to update the stellar mass of KiDS J0842 + 0059 obtained from photometry. Specifically, adopting the best-fitting model from Spiniello et al. (2024), we find a mass-to-

light ratio of $M/L_K = 0.94$, which, combined with a total K -band magnitude of 17.6, yields a stellar mass of $M_* \sim 2.5 \times 10^{11} M_\odot$. Using instead the stellar population results from Maksymowicz-Maciata et al. (2024) with a variable IMF, leads to a doubled mass estimate of $M_* \sim 5 \times 10^{11} M_\odot$. However, here, we will stick with our previous mass estimate from photometry, for uniformity with the rest of galaxies analysed in the INSPIRE sample. In a future publication, we will revise all the masses for the INSPIRE sample. Finally, in terms of mass density profile, KiDS J0842 + 0059 matches very well the local relic prototype NGC 1277 (Ferré-Mateu et al. 2017) and high- z red nuggets (Szomoru, Franx & van Dokkum 2012).

Modelling ground-based KiDS optical images, we estimated an effective radius of $R_e = 1 \text{ kpc}$ (Scognamiglio et al. 2020), and therefore classified KiDS J0842 + 0059 as ultra-compact. However, as shown in appendix B of Tortora et al. (2018), while KiDS effective radii are statistically precise and accurate on a population-level, individual cases – particularly those with the most compact sizes – exhibit larger uncertainties (up to ~ 50 per cent), and results can be affected by potential biases. Therefore, the main scope of this paper is to get an accurate measurement of the structural parameters of KiDS exploiting the accuracy of the unprecedented AO observations described in Section 2. A much higher precision is indeed achievable

⁸This is likely a lower-limit for σ_e , since the light contained in the R50 aperture is a mixture from inside and outside the effective radius, see appendix A in Spiniello et al. (2021c) for details.

thanks to a PSF which is around 6 times narrower than that of ground-based, seeing-limited data, and hence comparable to the effective radius of the galaxy. The following subsection presents the main results of the paper: a precise estimate of the structural parameters (Sérsic index, axis ratio, and effective radius) from the AO-supported NIR images.

3.2 Structural parameter estimates from AO observations

We first performed a preliminary analysis, to characterize in a non-parametric way the shape of the galaxy, by running the code AUTOPROF (Stone et al. 2021) on the AO-supported image. AUTOPROF fits elliptical isophotes to the image and then performs a non-parametric extraction of the azimuthally averaged light profile of the galaxy. The image of the galaxy is shown in an inset on the left panel of Fig. 2, together with the fitted isophotes; for visualization purposes, the isophotes are coloured in red, with one in cyan every fourth isophote. These are rounder in the center ($q = 0.8$) and more elongated in the outskirts ($q = 0.45$). The artificial twist in the position angle of the ellipses is also clear, with the central ($\lesssim 0.2''$) isophotes, where the effect of the PSF is stronger, rounder, and stretched towards the AO star (along the north-east direction). The AO higher-resolution image clearly shows the discy shape of the galaxy isophotes, which were blurred by the worse PSF in the KiDS image.

We used 2DPHOT (La Barbera et al. 2008) to fit the target image with a PSF-convolved 2D Sérsic model. The PSF used to convolve the Sérsic profile has been modelled as described in Section 2.3 and discussed in greater detail in Appendix A. We have also explored a more complex model and its impact on the fit using the combination of a Sérsic and an exponential profile in Appendix B. We have estimated the following parameters with their uncertainties: Sérsic index (n), axis ratio (q), position angle (θ_q), major axis, and circularized effective radius ($R_{e,\text{maj}}$ and $R_{e,\text{circ}} = \sqrt{q} R_{e,\text{maj}}$). The results of the fitting are compared with those obtained from KiDS seeing limited data and summarized in Table 1.

We have run AUTOPROF not only on the galaxy image, but also on the image of the best-fitting models convolved with the PSF.⁹ The radial SB profiles for image (black symbols) and best-fitting models (blue lines) are shown in the left panel of Fig. 2.

By fitting a single-Sérsic PSF-convolved profile we have found $n = 5.7 \pm 0.5$, $q = 0.24 \pm 0.1$, and $R_{e,\text{circ}} = 0.214 \pm 0.001$ arcsec, which corresponds to $R_e = 1$ kpc. The reduced χ^2 is 1.4. All parameters are determined with high accuracy, except for the axis ratio, which exhibits the largest relative error of approximately ~ 40 per cent.

The residual image is shown in the right panel of Fig. 2. We find an asymmetrical residual concentrated in the very central galactic region ($\lesssim 0.3$ arcsec), which can be probably connected to the inability to model the PSF with the needed accuracy. For example, imperfect PSF modelling, such as errors in the PSF slope, can influence the slope of the best-fitting profile and consequently the value of n . Nevertheless, the residuals calculated at different polar angles as a function of the distance from the center give a clear indication of a quite good fit.

By calculating the growth curve from the data shown in Fig. 2 and determining $R_{e,\text{maj}}$ as the radius enclosing half of the total light,

we find $R_{e,\text{maj}} = 0.45$ arcsec.¹⁰ It is closely matching the estimate obtained from fitting the 2D Sérsic profile to the image.

In Appendix B we also explore a double-component fit. Although the central component seems to address our challenge in accurately modelling the PSF at the galaxy position and therefore a slightly lower χ^2 , we express reservations about the two-component model's ability to represent a physically meaningful solution for the relic. This consideration is justified by the fact that the central component specifically fits the central SB profile, where the PSF plays a crucial role. Additionally, such best-fitting model's axis ratio and orientation mirror the characteristics of the more central isophotes.

In conclusion, the circularized effective radius is robustly determined with a value of $R_e = 0.214 \pm 0.001$, which corresponds to $R_e \sim 1$ kpc, confirming the KiDS seeing-dominated results reported in Scognamiglio et al. (2020). We have also confirmed that the galaxy presents discy isophotes. Therefore, we can now securely conclude that J0842 + 0059 is an ultra-compact (massive) galaxy, with size and shape resembling the local relic prototype NGC 1277 and high- z red nuggets.

4 MASS DENSITY PROFILES OF RELICS AND RED NUGGETS

The new high-resolution AO imaging has allowed us to confirm that KiDS J0842 + 0059 is an ultra-compact massive galaxy, according to the INSPIRE threshold on the circularized radius ($R_e \lesssim 2$ kpc, Tortora et al. 2018; Scognamiglio et al. 2020; Spiniello et al. 2021c).

We note that we cannot determine the intrinsic ellipticity of the galaxy. As a result, the circularized radius may underestimate the true effective radius. In the most extreme case, where the galaxy is a circular inclined disc seen in projection, the actual effective radius could be as large as the major-axis value. However, given the high metallicity, α -enhancement, and velocity dispersion of KiDS J0842 + 0059, it is highly unlikely that the galaxy is disc-dominated, reducing the possible impact of such projection effect. Nevertheless, since determining the intrinsic ellipticity and effective radius of galaxies is typically challenging, we limit our analysis to projected quantities and adopt homogeneous definitions across different galaxy samples.

The scope of this section is to compare the mass density profile of KiDS J0842 + 0059 to other galaxies, at different cosmic time, which are believed to have similar formation path. We therefore only focus on the circularized radius, since for local relics and high- z red nuggets it is the one most commonly used in the literature.

We first converted the measured SB profile of KiDS J0842 + 0059 into a projected mass density profile $\Sigma(R) = M/L \times I(R)$, where M/L is a constant stellar mass-to-light ratio¹¹ and $I(R)$ is the intensity of the best-fitting Sérsic profile. We determined the M/L by imposing that the total mass obtained by integrating $\Sigma(R)$ to large radii is equal to the value of $10^{11} M_\odot$ we have determined from KiDS photometric data in Scognamiglio et al. (2020). We then applied the same procedure to other samples. In particular, we determined the $\Sigma(R)$ for the entire INSPIRE sample, including ultra-compact

¹⁰Since no PSF deconvolution has been applied, this value represents an upper limit.

¹¹For simplicity, we assume no colour gradients, implying a constant M/L as a function of radius. Even if this assumption is not entirely accurate, colour gradients in massive galaxies are generally weak (Tortora et al. 2011), and their impact is minimal in the K band, which closely traces the stellar mass.

⁹The latter image is provided in output by 2DPHOT.

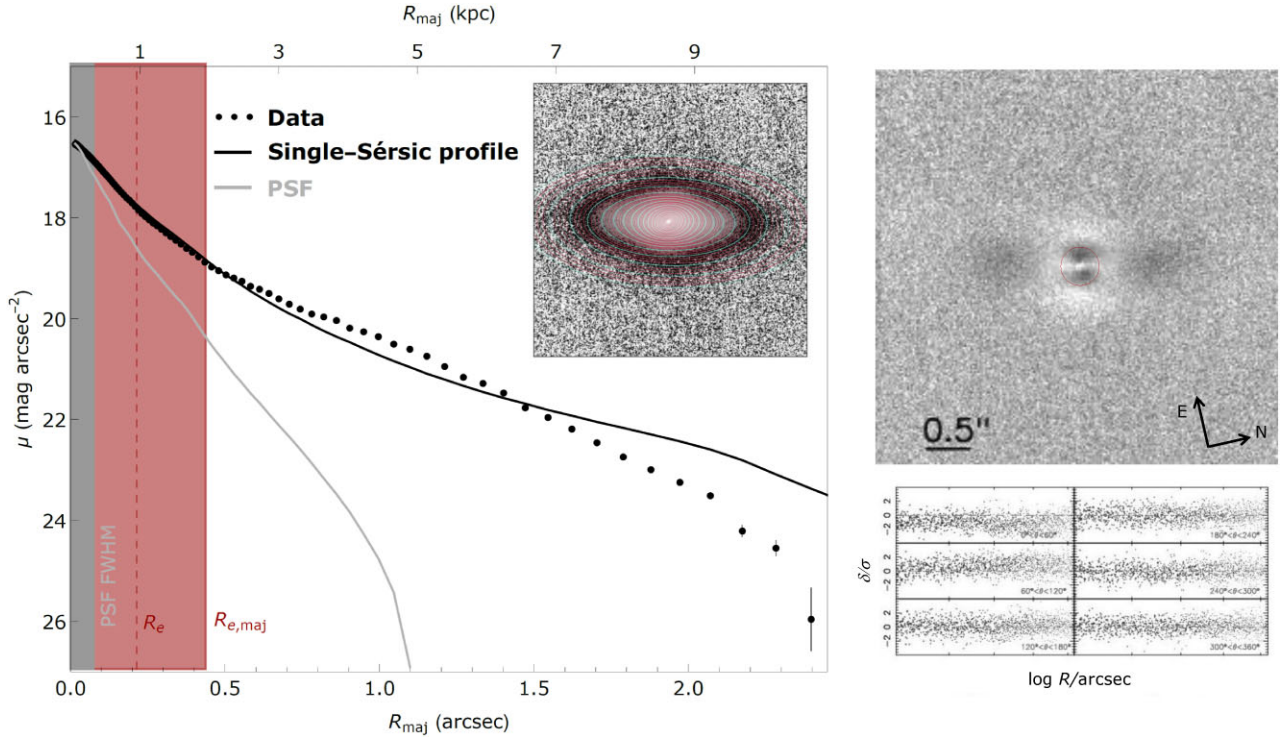


Figure 2. *Left.* The azimuthally averaged (1D) SB profile, derived from the isophotes shown in the inset panel, plotted as a function of the major axis R_e (black symbols with uncertainties). The solid black curve represents the best-fitting PSF-convolved model, obtained by fitting a 2D single-Sersic profile to the image using 2DPHOT. The parameters of this fit are listed in Table 1. The light-red (larger) region marks the major-axis R_e , while the red vertical dashed line indicates the circularized R_e for the best-fitting single-Sersic profile. The grey curve corresponds to the PSF used in the fitting, and the grey (smaller) shaded area represents the PSF’s FWHM. The galaxy image in the inset was created with AUTOPROF, with isophotal ellipses overlaid. *Right.* Residual image generated by 2DPHOT for the single-Sersic fit. Darker (lighter) grey is for larger (smaller) residuals. The red circle marks the circularized R_e . The image is aligned according to the SOUL image, with the AO star located in the upper-right direction. The residual image is computed as the difference between the observed image and the best-fitting model, divided by the error in the pixel. In the six lower panels, the residuals of the galaxy flux per pixel, after model subtraction, are plotted as a function of distance from the galaxy center, separated into different bins of polar angle.

Table 1. KiDS J0842 + 0059 structural parameters computed using KiDS and AO data for a single-Sersic profile. Angles are counted from west to north.

Data	PSF	n	q	θ_q (deg)	$R_{e,maj}$ (arcsec)	$R_{e,circ}$ (arcsec)	χ^2
KiDS	2-Moffat	3.27	0.29	76.8	0.420	0.226	1
AO	AO-Extrapolated	5.72 ± 0.5	0.24 ± 0.10	74 ± 2	0.437 ± 0.1	0.214 ± 0.001	1.4

galaxies of any DoR, and a randomly selected sample of 500 normal-size galaxies at $z < 0.5$ with $R_e > 2.5$ kpc and stellar masses similar to those of INSPIRE galaxies.

For these normal-sized galaxies we rely on KiDS r -band structural parameters and stellar masses obtained using KiDS photometry (Roy et al. 2018; Tortora et al. 2018, 2020; Scognamiglio et al. 2020). Finally, we obtained the $\Sigma(R)$ profiles for the red nuggets with $R_e < 2$ kpc and $M_* \gtrsim 6 \times 10^{10} M_\odot$ in the sample of $z \sim 2$ quiescent galaxies in Szomoru et al. (2012).

We plot the distribution of the mass density profiles of these samples in the left panel of Fig. 3 as shaded regions,¹² and compared to KiDS J0842+0059. KiDS J0842 + 0059 (black solid line) exhibits average structural properties and surface mass density profile with respect to the INSPIRE sample. What is clearly visible from this

figure is that the sample of INSPIRE galaxies (red shaded region) have mass density profiles which resemble perfectly those of high- z red nuggets (cyan shaded region).

Selecting only relic galaxies (DoR > 0.34, according to Spiniello et al. 2024) from INSPIRE does not alter these conclusions. Interestingly, when we focus on INSPIRE galaxies with lower DoR values (e.g. DoR < 0.7) or lower velocity dispersions (< 200 km s⁻¹)¹³, the scatter in the mass density profiles increases. We show this trend in the middle panel of Fig. 3, where the INSPIRE sample is divided in extreme relics (DoR > 0.7, red), relics (0.34 < DoR ≤ 0.7, green), and non-relics (DoR ≤ 0.34, blue). The average mass density profiles do not depend on the DoR, instead their distribution is broader among galaxies with low DoR compared to the most extreme relics in the sample. This probably reflects the more diverse nature of these

¹²We evaluated all the profiles at given radial values and joined all the data points. Then, we only retained points within the 16th and 84th percentile of the sample distribution in each radial bin, to exclude outlier profiles.

¹³The most extreme relics, including KiDS J0842 + 0059, exhibit higher velocity dispersions compared to UCMGs with lower DoR (Spiniello et al. 2024).

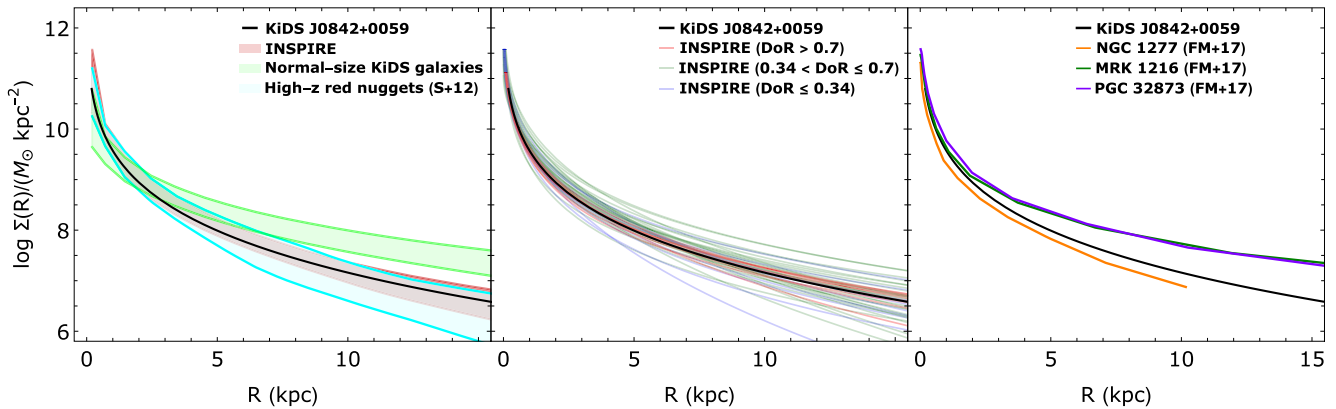


Figure 3. Projected mass density profile of KiDS J0842 + 0059 (black line). In the left panel, this is compared to the range around the median including 68 per cent of profiles for INSPIRE galaxies (Spiniello et al. 2024), normal-size KiDS galaxies (Roy et al. 2018; Tortora et al. 2020) and high- z red-nuggets (Szomoru et al. 2012). In the middle panel, the mass density profile of KiDS J0842 + 0059 is compared to the profiles of INSPIRE extreme relics (DoR > 0.7), relics ($0.34 < \text{DoR} \leq 0.7$), and non-relics (DoR ≤ 0.34). Finally, in the right panel, the mass density profile of KiDS J0842 + 0059 is compared to the profiles of the three local relics: NGC 1277, MRK 1216, and PGC 32 873 (Ferré-Mateu et al. 2017).

galaxies, which exhibit a broad range of star formation histories. Instead, extreme relic galaxies form a more homogeneous class of objects in terms of their stellar population properties, a consistency that also extends to their mass density profiles. In contrast, galaxies with lower DoR display more diverse properties, which likely contribute to the larger scatter observed in their stellar mass profiles. In future work, we will further investigate the stellar mass profiles of these galaxies to better understand the origin of this diversity.

Interestingly, only in the central regions ($R < 5$ kpc) we observe a mild tendency towards shallower profiles in galaxies with lower DoR. The profiles for normal-sized galaxies at similar redshift (green shaded region in the left panel) are similar to those of ultra-compact galaxies (and in particular to those with lower DoR) at all redshifts only in the innermost region, where the compact progenitor dominates the light budget (e.g. Barbosa et al. 2021). Instead, at larger radii ($> 3 - 5$ kpc), compact galaxies at both $z < 0.5$ and $z \sim 2$ have smaller mass densities.

Finally, the right panel of Fig. 3 compares the mass density profile of KiDS J0842 + 0059 with that of the three local relics. According to the DoR definition provided by Spiniello et al. (2024) NGC 1277 has a DoR = 0.95, therefore classified as the most extreme relic found so far. KiDS J0842 + 0059 (DoR = 0.73) exhibits a mass density profile that falls between that of NGC 1277 and those of the other two local relics, MRK 1216 and PGC 32873. The latter two, with their more (relatively) extended star formation histories, as found in Ferré-Mateu et al. (2017), have higher mass densities at all radii, suggesting they accreted more mass over cosmic time. They would have a slightly lower DoR according to the INSPIRE definition, but would still be classified as relics. The most extreme relics, such as NGC 1277 and KiDS J0842 + 0059, have not accreted significant *ex situ* mass throughout their entire cosmic histories. The behaviour of Mrk 1216 and PGC 32873, which both exhibit lower DoR than NGC 1277, is consistent with INSPIRE galaxies with lower-DoR.

We note that although all profiles would be shallower when surface mass density is plotted as a function of the major-axis radius instead of the circularized one (as done here), their relative differences would change only negligibly and therefore the conclusions of this comparison remain unaffected.

Summarizing the results, in terms of mass density profile, the extreme relic J0842 + 0059 and NGC 1277 are the perfect descendant of high- z red nuggets. Relics with more prolonged SFH (such

as MRK 1216 and PGC 32873) and normal-sized galaxies have shallower mass density profile, which is an indication of an accreted *ex situ* mass component that is instead missing in low- z extreme relics.

5 CONCLUSIONS

With the INSPIRE Project we have assembled a catalogue of 52 confirmed UCMGs at $z = 0.2-0.4$. Thanks to X-Shooter@VLT spectra, we have obtained, for each system, a precise estimate of the integrated velocity dispersion (D’Ago et al. 2023), stellar age, metal-content, and star formation history (Spiniello et al. 2021a, b, c, 2024), as well as individual elemental abundance and an estimate of the low-mass end of the IMF slope (Martín-Navarro et al. 2023; Maksymowicz-Maciata et al. 2024). Out of the 52 UCMGs, 38 have been classified as relic galaxies, having formed 75 per cent or more of their stellar mass at early cosmic times ($z > 2$). This sample increases by more than 3 times the number of spectroscopically confirmed relic galaxies in the nearby Universe. However, the photometrical selection of these galaxies was based on ground-based seeing dominated observations, which make their sizes and their shapes and morphologies possibly affected by large uncertainties. Observations with high spatial resolution are necessary to firmly confirm the ultra-compact nature of these systems, case-by-case, allowing to compare their shape and morphological properties with similar galaxies in the local Universe as well as at very high redshift, and hence draw a time evolution of the most compact massive galaxies in the Universe.

Therefore, using Adaptive Optics (AO) K_s -band observations obtained with SOUL@LBT, we targeted one of the most extreme relics in the INSPIRE sample, KiDS J0842 + 0059.

In particular, for KiDS J0842 + 0059 ground-based r -band observations indicated a size of $R_e \sim 1$ kpc and an extremely flattened light profile with an axis ratio of $q \sim 0.3$. These measurements, however, could have been highly affected by the fact that the FWHM of these images is much bigger than the apparent size of the object (Tortora et al. 2018). In this paper, using LUCI@LBT AO supported observations with PSF FWHM of 0.14 arcsec on the target, we unambiguously confirm the ultra-compact nature of this galaxy: KiDS J0842 + 0059 is a very compact galaxy with a circularized effective radius of 0.214 ± 0.001 arcsec, corresponding to $R_e \sim$

1 kpc, and presents discy isophotes with $q = 0.24 \pm 0.1$. Although based on a single system, the confirmation of the compactness of this extreme relic supports the goodness of the INSPIRE selection procedure and estimate of effective radii with seeing-limited KiDS data.

The structural properties and discy isophotal shapes of KiDS J0842 + 0059 fully resemble those accurately measured for the three local relic galaxies in Ferré-Mateu et al. (2017), for which HST data is available. They are also very similar to those of high-redshift red nuggets (e.g. Damjanov et al. 2009; Szomoru et al. 2012; Baker et al. 2025; Glazebrook et al. 2024). We have also shown that KiDS J0842 + 0059 has a mass density profile which resembles that of NGC 1277 and of high- z red-nuggets. On the contrary, UCMGs with a relatively more prolonged SFH (lower DoR) present a larger scatter in their distribution, compared to the most extreme relics. Both UCMGs with lower DoR at $R \lesssim 5$ kpc and normal-sized galaxies have shallower mass profiles, which might be explained by the formation of *ex situ* stars (accreted or later formed), a kind of process which is instead completely missing in the most extreme relics in the Universe, as KiDS J0842 + 0059 and NGC 1277.

Our results are thus perfectly consistent with the ‘two-phase’ scenario, where relic galaxies have stopped forming stars and have missed the merging phase (e.g. Oser et al. 2010).

Combined with spectroscopic evidence showing that KiDS J0842 + 0059 had assembled its entire stellar mass by $z = 2$, and is characterized by metal-rich, α -enhanced stellar populations (Spiniello et al. 2021c) and a dwarf-rich IMF (Maksymowicz-Maciata et al. 2024), we have now reached an unambiguous confirmation that KiDS J0842 + 0059 is a remnant of a high- z red nugget.

This is the first time that high-resolution imaging is obtained for a relic galaxy outside the local Universe with such a high data quality and spatial resolution. KiDS J0842 + 0059 is the relic galaxy with the most complete set of observations (seeing-dominated imaging in optical and NIR, high-quality spectroscopy and high-resolution NIR photometry), and thus, up to now, the most secure relic and compact galaxy beyond the local Universe. However, obtaining structural parameter measurements with AO- or space-based data for the remaining galaxies in the INSPIRE sample remains crucial to definitively confirm their compactness, ensuring the uniqueness of this relic galaxy sample. The Euclid Wide Survey (Euclid Collaboration 2022; Euclid Collaboration: Mellier et al. 2025) will help us taking a first step in this direction, providing high spatial resolution data on the 17 INSPIRE galaxies in the Southern hemisphere, as well as observing more than 30 000 new UCMGs, extending the redshift boundaries up to $z \sim 2$. This will finally allow us to bridge the gap between the two phases of the massive galaxies formation scenario (Oser et al. 2010). Furthermore, recently, a proposal to follow-up the most extreme relic, the largest and smaller UCMGs and the object with the largest velocity dispersion in INSPIRE has been approved. We have been awarded MUSE-IFS data in the Narrow Field Mode (NFM) AO configuration ($0.025''/\text{pixel}$ with a $7.5'' \times 7.5''$ field of view) to get unprecedentedly precise structure parameters and to resolve the internal structure and kinematics of relics for the first time ever beyond the local Universe.

ACKNOWLEDGEMENTS

We thank the anonymous referee for their suggestions, which helped improve the paper. CT acknowledges the INAF grant 2022 LEMON.

GT acknowledges financial support from the European Research Council (ERC) Advanced Grant under the European Union’s Horizon Europe research and innovation programme (grant agreement AdG GALPHYS, No. 101055023). GD acknowledges support by UKRI-STFC grants: ST/T003081/1 and ST/X001857/1. AFM has received support from RYC2021-031099-I and PID2021-123313NA-I00 of MICIN/AEI/10.13039/501100011033/FEDER,UE, NextGenerationEU/PRT. DS is supported by JPL, which is operated under a contract by Caltech for NASA. FLB acknowledges support from INAF minigrant 1.05.23.04.01.

DATA AVAILABILITY

The INSPIRE catalogue is publicly available through the ESO Phase 3 Archive Science Portal under the collection INSPIRE (https://archive.eso.org/scienceportal/home?data_collection=INSPIRE, <https://doi.eso.org/10.18727/archive/36>). The LUCI-SOUL@LBT data used in this paper will be made available to the interested readers upon request to the authors.

REFERENCES

- Agapito G., Arcidiacono C., Quirós-Pacheco F., Esposito S., 2014, *Exp. Astron.*, 37, 503
- Baker W. M. et al., 2025, *NatAs*, 9, 141
- Barbosa C. E., Spiniello C., Arnaboldi M., Coccato L., Hilker M., Richtler T., 2021, *A&A*, 649, A93
- Beasley M. A., Trujillo I., Leaman R., Montes M., 2018, *Nature*, 555, 483
- Beltramo-Martin O., Correia C. M., Neichel B., Fusco T., 2018, *MNRAS*, 481, 2349
- Britton M. C., 2006, *PASP*, 118, 885
- Buitrago F. et al., 2018, *A&A*, 619, A137
- Cantalloube F. et al., 2020, *A&A*, 638, A98
- Charbonnier A. et al., 2017, *MNRAS*, 469, 4523
- Comerón S. et al., 2023, *A&A*, 675, A143
- D’Ago G. et al., 2023, *A&A*, 672, A17
- Damjanov I. et al., 2009, *ApJ*, 695, 101
- Damjanov I., Geller M. J., Zahid H. J., Hwang H. S., 2015, *ApJ*, 806, 158
- de Jong J. T. A. et al., 2015, *A&A*, 582, A62
- de Jong J. T. A. et al., 2017, *A&A*, 604, A134
- Euclid Collaboration: Scaramella et al., 2022, *A&A*, 662, A112
- Euclid Collaboration: Mellier et al., 2025, *A&A*, 697, A1
- Ferré-Mateu A., Mezcua M., Trujillo I., Balcells M., van den Bosch R. C. E., 2015, *ApJ*, 808, 79
- Ferré-Mateu A., Trujillo I., Martín-Navarro I., Vazdekis A., Mezcua M., Balcells M., Domínguez L., 2017, *MNRAS*, 467, 1929
- Flores-Freitas R., Chies-Santos A. L., Furlanetto C., De Rossi M. E., Ferreira L., Zenocatti L. J., Alamo-Martínez K. A., 2022, *MNRAS*, 512, 245
- Glazebrook K. et al., 2024, *Nature*, 628, 277
- Haguenauer P., Agapito G., Neichel B., 2022, in Schreiber L., Schmidt D., Vernet E. eds, Proc. SPIE Conf. Ser. Vol. 12185, Adaptive Optics Systems VIII. SPIE, Bellingham, p. 1218565
- Heidt J. et al., 2018, in Evans C. J., Simard L., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 10702, Ground-based and Airborne Instrumentation for Astronomy VII. SPIE, Bellingham, p.107020B
- Hill J. M., 2010, *Appl. Opt.*, 49, 115
- Hopkins P. F., Bundy K., Murray N., Quataert E., Lauer T. R., Ma C.-P., 2009, *MNRAS*, 398, 898
- La Barbera F., de Carvalho R. R., Kohl-Moreira J. L., Gal R. R., Soares-Santos M., Capaccioli M., Santos R., Sant’anna N., 2008, *PASP*, 120, 681
- Labbé I. et al., 2023, *Nature*, 616, 266
- Lisiecki K., Małek K., Siudek M., Pollo A., Krywult J., Karska A., Junais, 2023, *A&A*, 669, A95
- Maksymowicz-Maciata M. et al., 2024, *MNRAS*, 531, 2864

- Martín-Navarro I., van de Ven G., Yıldırım A., 2019, *MNRAS*, 487, 4939
- Martín-Navarro I. et al., 2023, *MNRAS*, 521, 1408
- Moura M. T., Chies-Santos A. L., Furlanetto C., Zhu L., Canossa-Gosteinski M. A., 2024, *MNRAS*, 528, 353
- Naab T., Johansson P. H., Ostriker J. P., 2009, *ApJ*, 699, L178
- Neichel B. et al., 2021, in Schreiber L., Schmidt D., Vernet E., eds, Proc. SPIE Conf. Ser. Vol. 114482, Adaptive Optics Systems VII. SPIE, Bellingham, p. 114482T
- Oser L., Ostriker J. P., Naab T., Johansson P. H., Burkert A., 2010, *ApJ*, 725, 2312
- Pinna E. et al., 2019, in Adaptive Optics for Extremely Large Telescopes 6—Conference Proceedings. <http://ao4elt6.copl.ulaval.ca/proceedings/401-yw8a-251.pdf>
- Pinna E. et al., 2023, *aoel.conf*, 80
- Quilis V., Trujillo I., 2013, *ApJ*, 773, L8
- Roy N. et al., 2018, *MNRAS*, 480, 1057
- Salvador-Rusiñol N., Ferré-Mateu A., Vazdekis A., Beasley M. A., 2022, *MNRAS*, 515, 4514
- Scognamiglio D. et al., 2020, *ApJ*, 893, 4
- Scognamiglio D. et al., 2024, *MNRAS*, 534, 1597
- Spiniello C., Tortora C., D’Ago G., Napolitano N. R., *Inspire Team*, 2021a, *The Messenger*, 184, 26
- Spiniello C. et al., 2021b, *A&A*, 646, A28
- Spiniello C. et al., 2021c, *A&A*, 654, A136
- Spiniello C. et al., 2024, *MNRAS*, 527, 8793
- Stone C. J., Arora N., Courteau S., Cuillandre J.-C., 2021, *MNRAS*, 508, 1870
- Szomoru D., Franx M., van Dokkum P. G., 2012, *ApJ*, 749, 121
- Tortora C., Napolitano N. R., Romanowsky A. J., Jetzer P., Cardone V. F., Capaccioli M., 2011, *MNRAS*, 418, 1557
- Tortora C. et al., 2016, *MNRAS*, 457, 2845
- Tortora C. et al., 2018, *MNRAS*, 481, 4728
- Tortora C. et al., 2020, *A&A*, 638, L11
- Trujillo I., Cenarro A. J., de Lorenzo-Cáceres A., Vazdekis A., de la Rosa I. G., Cava A., 2009, *ApJ*, 692, L118
- Trujillo I., Ferré-Mateu A., Balcells M., Vazdekis A., Sánchez-Blázquez P., 2014, *ApJ*, 780, L20
- Turchi A., Masciadri E., Veillet C., 2022, in Marshall H. K., Spyromilio J., Usuda T., eds, Proc. SPIE Conf. Ser. Vol. 12182, Ground-based and Airborne Telescopes IX. SPIE, Bellingham, p. 121824O
- Wellons S. et al., 2016, *MNRAS*, 456, 1030
- Yıldırım A., van den Bosch R. C. E., van de Ven G., Martín-Navarro I., Walsh J. L., Husemann B., Gültekin K., Gebhardt K., 2017, *MNRAS*, 468, 4216
- Zolotov A. et al., 2015, *MNRAS*, 450, 2327

APPENDIX A: PSF DETERMINATION

We selected *TIPTOP* (Neichel et al. 2021) as simulation software. Our approach consists of two steps. First, we fitted the 4 stars (AO and PSF stars) with *TIPTOP*, and then we extrapolated the PSF to the galaxy position. The PSF modelling includes mainly four atmospheric parameters: (a) seeing, (b) outer scale, (c) vertical turbulence profile, and (d) wind profile, and the telescope and AO loop parameters. As merit function for the fit, we used the uncertainty on the FWHM of the 4 stars. The complete set of *TIPTOP* parameters is summarized in Table A1, which lists in the first block the parameters we kept fixed (first block) and those we fit (second block), with the assumptions we will describe below.

The three PSF stars in the field have a distinct elongation at ~ 45 deg: this is expected given their relative position to the AO guide star. The AO guide star PSF also has a small amount of elongation, about a factor of 1.2, along the 45 degree direction. This is probably due to the tip/tilt oscillation caused by some vibrations in the telescope path. Moreover, there is a triangular shape near the core of the AO guide star PSF, which is caused by a non-common path aberration between SOUL and LUCI. We neglected such features in

the PSF determination because they have a small effect on the off-axis PSF.

We used the AO–star FWHM to estimate the *TIPTOP* parameters associated with the magnitude of the oscillations, corresponding to a 2D Gaussian kernel with an FWHM of 0.068 and 0.056 arcsec on the major and minor axis, respectively. Note that these values are compatible but slightly smaller than the residual tip/tilt jitter estimated by the AO telemetry (0.03 arcsec RMS corresponds to an FWHM of 0.07 arcsec). They allow us to obtain the correct FWHM of the AO guide star (~ 0.10 and ~ 0.08 arcsec for major and minor axis, respectively) in the range of seeing values considered for the fit (~ 0.9 arcsec). In fact, the AO guide star FWHM depends mainly on the seeing value and on the vibrations level, it is marginally affected by the outer scale, the wind speed and direction (the wind profile affects the halo of the PSF as can be seen in Cantalloube et al. 2020) and it is not affected by the vertical turbulence (C_n^2) profile, which, instead, strongly affects off-axis stars/objects.

We decided to simplify the estimation of the atmospheric profiles choosing a 4-layer configuration with heights 103, 725, 2637, and 11068 m as in Agapito et al. (2014) and we changed the vertical turbulence (C_n^2) profile to produce different isoplanatic angles. The wind speed profile was chosen using as the ground wind speed the value measured by the telescope anemometers, 4 m s^{-1} . For the wind speed in the higher layers, we explored a range of values close to those produced by the ALTA forecasts (Turchi, Masciadri & Veillet 2022): the jet stream speed during the night was predicted to be between 40 and 60 m s^{-1} . Finally, we decided to fix the outer scale at a value of 25 m, which is compatible with the statistics of this parameter estimated by the Adaptive Optics Facility at the ESO VLT between 2019 and 2022 (Haguenaer, Agapito & Neichel 2022).

The results in terms of FWHM of the major axis of the PSFs are summarized in Fig. A1. The first set of PSFs is taken from the LUCI images where field stars are present, while the second set of PSFs is the one generated by *TIPTOP* at the location of the field star and of the galaxy. The FWHM values are computed using a fit to a Moffat function. Note that the major axis is aligned with the lines connecting the AO star and the field stars.

We decided to not use the FWHM of the minor axis of the field stars in this fit. This is because it has a lower sensitivity to the distance from the AO star than the major axis, and a number of bad pixels prevent us from getting an accurate estimate of this FWHM on one of the three field stars. However, this decision is not significant for the result, as the FWHM of the minor axis has a strong correlation with the major axis and therefore does not provide independent elements to the fit (Britton 2006; Beltramo-Martin et al. 2018).

Table A1. PSF model parameters.

Type	Parameter	Value
Fixed	Telescope elevation [deg]	57 ± 1
	DIMM seeing [arcsec]	0.9 ± 0.1
	Outer scale [m]	25
	Ground wind speed [m s^{-1}]	4 ± 1
	Guide star magnitude ($R + I$)	16
	Pupil sampling [pixel]	13×13
	Loop frame rate [Hz]	358
	Loop gain (tilt)	0.6
	Loop gain (high order modes)	0.25
	Fitting	Seeing at zenith [arcsec]
Isoplanatic angle [arcsec]		2.01
Average wind speed [m/s]		19

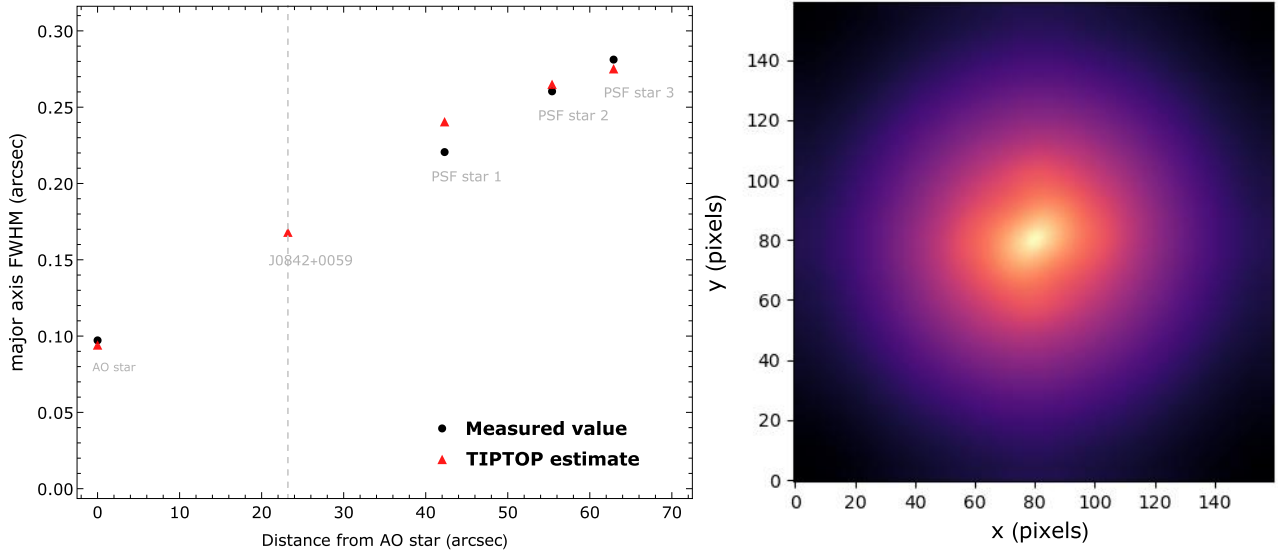


Figure A1. *Left.* FWHM of the PSFs at the position of the four stars and the galaxy for the major axis only. Comparison between the values estimated from LUCI images (black dots) and the TIPTOP estimate (red triangles). Only the TIPTOP estimate is available at 23.2 arcsec, where the galaxy is located (dashed line). *Right.* Image of the PSF used in the structural parameter derivation procedure.

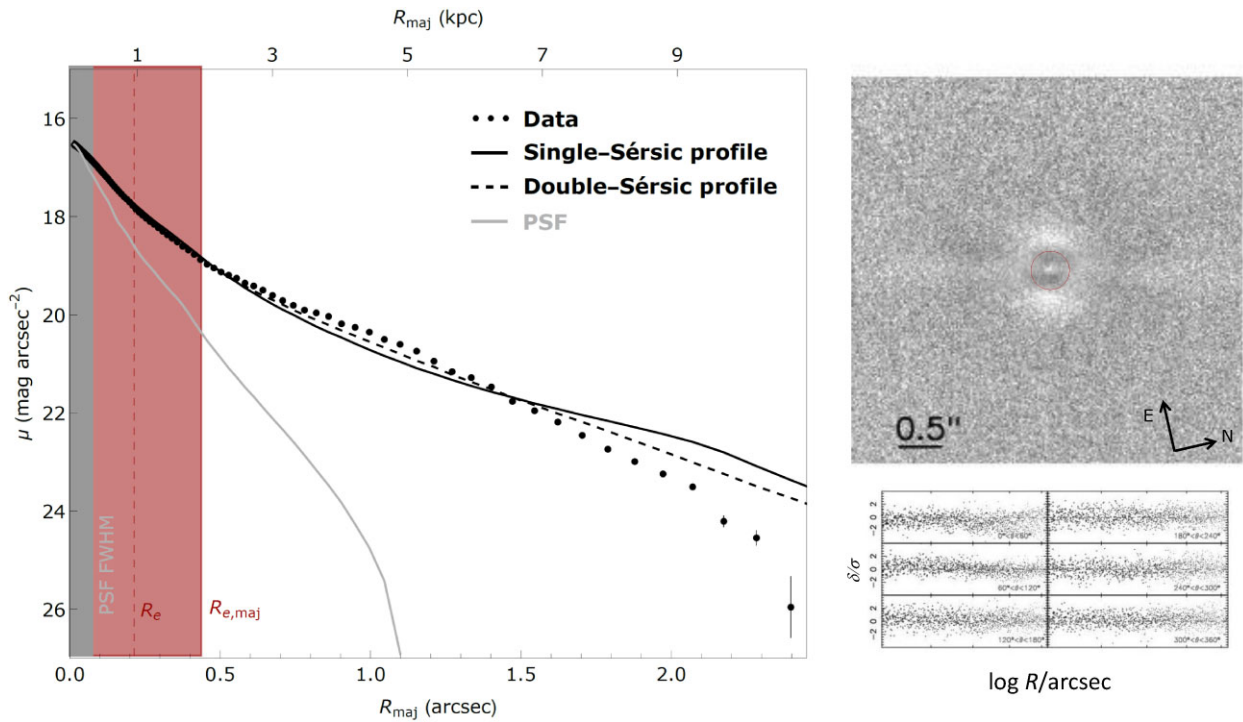


Figure A2. *Left.* The same as Fig. 2, but with in addition the best-fitting PSF-convolved model obtained by fitting a double-Sersic profile using 2DPHOT, plotted as dashed black line. *Right.* Residual image generated by 2DPHOT for the double-Sersic fit. The red circle marks the circularized R_e of the single-Sersic model. See Fig. 2 for more details.

The PSF image is shown in the right panel of Fig. A1. This is clearly stretched in the upper-right direction in the SOUL tile, which is in the direction of the AO star.

Finally, we note that the estimation of the PSF does not take into account the temporal variations of the telescope elevation, the atmospheric parameters (seeing, outer scale, vertical turbulence, and wind profile) and the structural vibrations, since we have considered average values.

APPENDIX B: DOUBLE-SÉRCIC FIT

We have also fitted to the data a combination of a Sérsic and an exponential profile. The best-fitting model and the residual are shown in Fig. A2.

The fit gives $n = 3.82$, $q = 0.58$, and $R_{e,\text{circ}} = 0.05$ arcsec for the innermost component and $q = 0.138$ and $R_{e,\text{circ}} = 0.31$ arcsec for the second one, for which n is fixed to 1. This fit returns a $\chi^2 = 1.2$,

slightly lower than the value obtained for the single-Sérsic fit. For simplicity, we refer to the more concentrated component as a bulge, and we call the $n = 1$ component a disc. This allows us to define a bulge-to-total parameter, as the ratio of the total flux in the bulge and total flux in the galaxy. The bulge component is dominating over the disc with a $B/T = 0.68$. Accordingly to the smaller χ^2 , the residuals shown in the right panel of Fig. A2 exhibit a slight improvement compared to the single-Sérsic case. However, since the central component primarily fits the inner SB profile, where the PSF has the greatest influence, we consider this solution less meaningful and tend to exclude it.

¹INAF – Osservatorio Astronomico di Capodimonte, Salita Moiarriello 16, I-80131 Napoli, Italy

²Max-Planck-Institut für Extraterrestrische Physik (MPE), Giessenbachstraße 1, D-85748 Garching, Germany

³INAF – Osservatorio Astrofisico di Arcetri, Largo Enrico Fermi 5, I-50125 Firenze, Italy

⁴Sub-Department of Astrophysics, Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH, UK

⁵Institute for Astrophysics, School of Physics, Zhengzhou University, Zhengzhou 450001, China

⁶Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK

⁷Department of Physics E. Pancini, University Federico II, Via Cinthia 6, I-80126 Naples, Italy

⁸European Southern Observatory, Karl-Schwarzschild-Straße 2, D-85748 Garching, Germany

⁹INAF – Osservatorio Astronomico di Brera, via Brera 28, I-20121 Milano, Italy

¹⁰DiSAT, Università degli Studi dell’Insubria, via Valleggio 11, I-22100 Como, Italy

¹¹Instituto de Astrofísica de Canarias, Vía Láctea s/n, E-38205 La Laguna, Tenerife, Spain

¹²Departamento de Astrofísica, Universidad de La Laguna, E-38200 La Laguna, Tenerife, Spain

¹³Finnish Centre for Astronomy with ESO (FINCA), University of Turku, FI-20014 Turku, Finland

¹⁴Tuorla Observatory, Department of Physics and Astronomy, University of Turku, FI-20014 Turku, Finland

¹⁵School of Physics, H.H. Wills Physics Laboratory, Tyndall Avenue, University of Bristol, Bristol BS8 1TL, UK

¹⁶Jet Propulsion Laboratory, California Institute of Technology, 4800, Oak Grove Drive, Pasadena, CA 91109, USA

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.