



UNIVERSITY
OF TURKU

This is a self-archived – parallel-published version of an original article.

This article has been accepted for publication in Monthly notices of the royal astronomical society ©: 2023 .Published by Oxford University Press on behalf of the Royal Astronomical Society. All rights reserved.

AUTHOR Ertini K., Folatelli G., Martinez L., Bersten M.C., Anderson J.P., Ashall C., Baron E., Bose S., Brown P.J., Burns C., Derkacy J.M., Ferrari L., Galbany L., Hsiao E., Kumar S., Lu J., Mazzali P., Morrell N., Orellana M., Pessi P.J., Phillips M.M., Piro A.L., Polin A., Shahbandeh M., Shappee B.J., Stritzinger M., Suntzeff N.B., Tucker M., Elias-Rosa N., Kuncarayakti H., Gutiérrez C.P., Kozyreva A., Müller-Bravo T.E., Chen T.W., Hinkle J.T., Payne A.V., Székely P., Szalai T., Barna B., Könyves-Tóth R., Bánhidi D., Bíró I.B., Csányi I., Kriskovits L., Pál A., Szabó Z., Szakáts R., Vida K., Vinkó J., Gromadzki M., Harvey L., Nicholl M., Paraskeva E., Young D.R., Englert B.

TITLE SN 2021gno: a calcium-rich transient with double-peaked light curves

YEAR 2023

DOI <https://doi.org/10.1093/mnras/stad2705>

VERSION Publisher's PDF

CITATION K Ertini, G Folatelli, L Martinez, M C Bersten, J P Anderson, C Ashall, E Baron, S Bose, P J Brown, C Burns, J M DerKacy, L Ferrari, L Galbany, E Hsiao, S Kumar, J Lu, P Mazzali, N Morrell, M Orellana, P J Pessi, M M Phillips, A L Piro, A Polin, M Shahbandeh, B J Shappee, M Stritzinger, N B Suntzeff, M Tucker, N Elias-Rosa, H Kuncarayakti, C P Gutiérrez, A Kozyreva, T E Müller-Bravo, T -W Chen, J T Hinkle, A V Payne, P Székely, T Szalai, B Barna, R Könyves-Tóth, D Bánhidi, I B Bíró, I Csányi, L Kriskovits, A Pál, Zs Szabó, R Szakáts, K Vida, J Vinkó, M Gromadzki, L Harvey, M Nicholl, E Paraskeva, D R Young, B Englert, SN 2021gno: a calcium-rich transient with double-peaked light curves, Monthly Notices of the Royal Astronomical Society, Volume 526, Issue 1, November 2023, Pages 279–298, <https://doi.org/10.1093/mnras/stad2705>

SN 2021gno: a calcium-rich transient with double-peaked light curves

K. Ertini ^{1,2}★, G. Folatelli ^{1,2,3}, L. Martinez ^{1,2,4}, M. C. Bersten ^{1,2,3}, J. P. Anderson ^{5,6}, C. Ashall ⁷, E. Baron ^{8,9}, S. Bose ^{10,11}, P. J. Brown ¹², C. Burns ¹³, J. M. DerKacy ^{8,7}, L. Ferrari ^{1,2}, L. Galbany ^{14,15}, E. Hsiao ¹⁶, S. Kumar ¹⁶, J. Lu ¹⁶, P. Mazzali ^{17,18}, N. Morrell ¹⁹, M. Orellana ^{4,20}, P. J. Pessi ¹, M. M. Phillips ¹⁹, A. L. Piro ¹³, A. Polin ^{13,21}, M. Shahbandeh ¹⁶, B. J. Shappee ²², M. Stritzinger ²³, N. B. Suntzeff ¹², M. Tucker ²², N. Elias-Rosa ^{24,14}, H. Kuncarayakti ^{25,26}, C. P. Gutiérrez ^{25,26}, A. Kozyreva ²⁷, T. E. Müller-Bravo ^{14,15}, T. -W. Chen ²⁸, J. T. Hinkle ²², A. V. Payne ²², P. Székely ^{29,30}, T. Szalai ^{29,30,31}, B. Barna ²⁹, R. Könyves-Tóth ^{31,32,33}, D. Bánhidi ³⁴, I. B. Bíró ^{30,34}, I. Csányi ³⁴, L. Kriskovits ^{31,32}, A. Pál ^{31,32,35,36,37}, Zs. Szabó ^{31,32,38,39}, R. Szakáts ^{31,32}, K. Vida ^{31,32}, J. Vinkó ^{31,32,35,40}, M. Gromadzki ⁴¹, L. Harvey ⁴², M. Nicholl ⁴³, E. Paraskeva ⁴⁴, D. R. Young ⁴⁵ and B. Engler ⁴⁶

Affiliations are listed at the end of the paper

Accepted 2023 September 1. Received 2023 August 25; in original form 2023 January 5

ABSTRACT

We present extensive ultraviolet (UV) and optical photometric and optical spectroscopic follow-up of supernova (SN) 2021gno by the ‘Precision Observations of Infant Supernova Explosions’ (POISE) project, starting less than 2 d after the explosion. Given its intermediate luminosity, fast photometric evolution, and quick transition to the nebular phase with spectra dominated by [Ca II] lines, SN 2021gno belongs to the small family of Calcium-rich transients. Moreover, it shows double-peaked light curves, a phenomenon shared with only four other Calcium-rich events. The projected distance from the centre of the host galaxy is not as large as other objects in this family. The initial optical light-curve peaks coincide with a very quick decline of the UV flux, indicating a fast initial cooling phase. Through hydrodynamical modelling of the bolometric light curve and line velocity evolution, we found that the observations are compatible with the explosion of a highly stripped massive star with an ejecta mass of $0.8 M_{\odot}$ and a ^{56}Ni mass of $0.024 M_{\odot}$. The initial cooling phase (first light-curve peak) is explained by the presence of an extended circumstellar material comprising $\sim 10^{-2} M_{\odot}$ with an extension of $1100 R_{\odot}$. We discuss if hydrogen features are present in both maximum-light and nebular spectra, and their implications in terms of the proposed progenitor scenarios for Calcium-rich transients.

Key words: stars: massive – supernovae: general – supernovae: individual: SN 2021gno.

1 INTRODUCTION

In the past decade, the advent of high-cadence, all-sky surveys has drastically increased the number of discovered transients, including supernovae (SNe) which lie outside the traditional classification schemes. Filippenko et al. (2003) first classified ‘Calcium-rich’ (or Ca-rich) SNe after observing four SNe with very strong emission from the Ca II near-infrared triplet, and the [Ca II] doublet near 730 nm. These SNe were first classified as Type Ib or Ic according to their early spectra, but they started to show distinct features as they evolved. Perets et al. (2010) studied an SN with similar properties, SN 2005E, which additionally was sub-luminous (absolute B -band peak magnitude of -14.8 mag) and rapidly evolving.

As more of these transients have been discovered, the ‘Ca-rich’ classification has become more robust. Kasliwal et al. (2012) identified the main properties of these transients as peak luminosity values

that are intermediate between those of novae and SNe (-14 to -16.5 mag), faster photometric evolution than normal SNe, photospheric velocities comparable to normal SNe, an early transition to the nebular phase, and nebular spectra dominated by calcium emission. These events have also been referred to as ‘Calcium-rich gap transients’ because of their location in the luminosity gap between novae and SNe. Their name is reinforced, not only by strong calcium emission but by the high nebular Ca to O line ratio they exhibit. Shen, Quataert & Pakmor (2019) argued that while these transients have such high Ca to O line ratios, this does not necessarily imply a large production of Ca, thus they adopt the name ‘Calcium-strong Transients.’ Furthermore, Polin, Nugent & Kasen (2021) showed that models with only 1 per cent of Ca produce nebular spectra that cool primarily through [Ca II] emission. For simplicity, in this work, we will refer to this class of events as ‘Ca-rich SNe’.

The Ca-rich class is relatively new, with few members, and they present several characteristics that make them a heterogeneous group. The majority of these transients are generally found in elliptical or S0 galaxies, which suggests a relationship with an old stellar

* E-mail: keilaertini@gmail.com

population (Dong et al. 2022). Additionally, they are often found at large offsets from the nuclei of their host galaxies (Kasliwal et al. 2012; Lyman et al. 2016). Foley (2015) studied a sample of thirteen Ca-rich transients and found that approximately one-third of them were located at projected distances greater than 20 kpc from the galaxy nuclei. However, there is a fraction of these transients that are found well within their host galaxies and near star-forming regions, as is typical for core-collapse (CC) SNe. For example, iPTF16hgs (De et al. 2018b) was located at about 6 kpc (projected distance) from the centre of its star-forming host galaxy, iPTF15eqv (Milisavljevic et al. 2017) was offset by 6.5 kpc from the centre of its spiral host galaxy, and SN 2019ehk (Jacobson-Galán et al. 2020b), the closest Ca-rich transient to date, was found in the well-known spiral galaxy NGC 4321 (Messier 100) at 1.8 kpc of projected distance from its nucleus.

Adding to the heterogeneity within this class, there is a small number of Ca-rich transients that present double-peaked light curves. This is the case of iPTF16hgs, SN 2018lqo (De et al. 2020), and SN 2019ehk. There is another object, iPTF14gqr (De et al. 2018a), that has a double-peaked light curve and shares all the photometric properties of a Ca-rich transient, but its spectrum at maximum light is quite distinct, and is thus not considered to be a canonical class member.

The diversity in the observed properties of these transients may suggest different physical origins. The origin of Ca-rich transients is still a matter of active debate, and multiple scenarios have been proposed (see Shen et al. 2019, for a recent review). Given the remote locations where these transients are usually found, the proposed progenitor systems frequently involve white dwarfs (WDs). For example, a WD passing close enough to an intermediate-mass black hole ($M_{BH} < 10^5 M_{\odot}$), or a WD tidally disrupted by a neutron star (NS) (Sell et al. 2015). However, many Ca-rich transients are explained as the detonation of a He shell on the surface of a WD (Perets et al. 2010; Waldman et al. 2011), which could be accreted from a He WD or a He star. There is, however, another explanation for the origin of a minority of Ca-rich transients that are found relatively close to their host-galaxy centres and near star-forming regions. They could be explained by the CC explosion of a massive star, which has been stripped of most of its hydrogen envelope (Kawabata et al. 2010). Another possibility within the CC scenario is a He star in a binary system with an NS, which leads to the complete stripping of the He envelope, leading to what is called an ultra-stripped SN (USSN; Tauris et al. 2013; Tauris, Langer & Podsiadlowski 2015).

Nakaoka et al. (2021) suggested that the double-peaked iPTF16hgs, iPTF14gqr, and SN 2019ehk (and potentially iPTF15eqv, which has no pre-maximum observations) belong to a subpopulation within the Ca-rich class that is associated with the USSN scenario. At the same time, De et al. (2021) suggested that iPTF15eqv and SN 2019ehk belong to a class of CCSNe with low-mass CO cores distinct from the thermonuclear Ca-rich transients found in old environments.

In this paper we present photometry and spectra of SN 2021gno, an SN which is located in the relatively inner regions of its host galaxy NGC 4165. SN 2021gno was initially classified as a Type II SN (Hung et al. 2021), and later re-classified as Type Ib and Type IIb (Dahiwalé & Fremling 2021; Perley 2021, respectively), but its spectral as well as photometric evolution indicates it belongs to the Ca-rich class. We present observations obtained by the Precision Observations of Infant Supernova Explosions (POISE¹; Burns et al. 2021) collaboration, starting 0.8 d after discovery and spanning 25 d

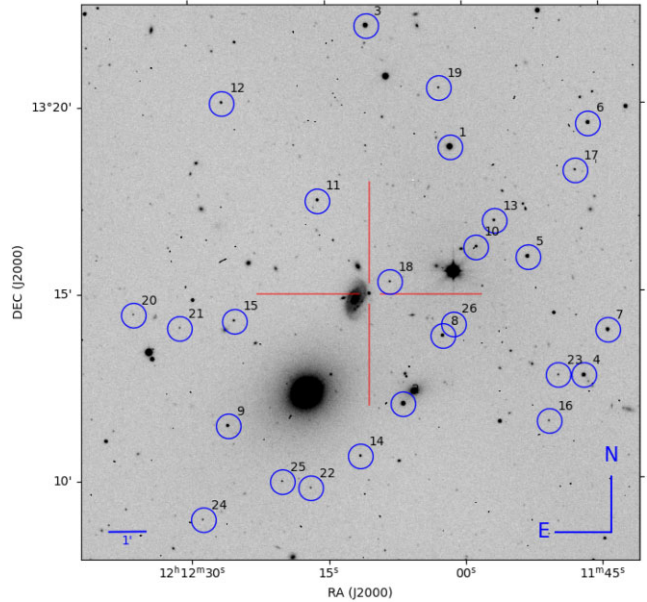


Figure 1. Swope Telescope *r*-band image of SN 2021gno (red cross) in NGC 4165. The blue circles indicate the local sequence stars. North is up and east is to the left.

with ≈ 1 observation per day in all BV*ugri*-bands. Such an early and high-cadence follow-up allows to observe initial light-curve peaks, which have been only seen in a handful of Ca-rich objects. Recently, Jacobson-Galán et al. (2022) presented multiwavelength observations of SN 2021gno and suggested it was produced by the explosion of a WD star, likely due to the merger of a hybrid + CO WD system. Here we present an alternative progenitor scenario for this event: the CC explosion of a highly stripped, massive star.

The paper is organized as follows. In Section 2, we present the observations and data reduction of SN 2021gno. We analyse its photometric and spectroscopic properties in Section 3. In Section 4, we present the bolometric light curve and associated hydrodynamical modelling for our proposed progenitor scenario. Finally, in Section 5 we provide a summary of our results and a discussion regarding the possible progenitor systems of Ca-rich transients.

2 OBSERVATIONAL DATA

SN 2021gno (Fig. 1) was discovered by the Zwicky Transient Facility (ZTF; Masci et al. 2019) on UTC 2021 Mar 20 05:38:54.24 (JD = 2459293.73) with a magnitude of $r = 18.2$ mag. Its coordinates are $\alpha = 12^h 12^m 10^s.294$ and $\delta = +13^\circ 14' 57''.03$. The object was located in the galaxy NGC 4165, at an angular distance of 24.54 arcsec from its centre, which corresponds to a projected physical separation of 4.56 kpc. NGC 4165 has a reported heliocentric redshift of $z = 0.0062$ (Albaret et al. 2017). A distance of 38.4 ± 1.8 Mpc was adopted for NGC 4165 (see Section 3.1 for more details).

The last non-detection was reported by ZTF on JD = 2459291.85, less than 2 d before the detection, with a limiting magnitude of $r = 20.54$ mag. We take the estimated explosion epoch as the midpoint between the last non-detection and the first detection with an uncertainty equal to half the interval between those epochs, therefore at JD = 2459292.79 \pm 0.94. Unless noted otherwise, we will provide all epochs relative to this adopted explosion date in the rest frame of the SN.

¹<https://poise.obs.carnegiescience.edu/>

2.1 Photometry

The POISE collaboration started its follow-up observations of SN 2021gno shortly after its discovery. Optical photometry in the *BVugri* bands was obtained with the 1.0 m Swope Telescope at Las Campanas Observatory (LCO), Chile. The observations were reduced using the POISE photometric pipeline, which closely follows the Carnegie Supernova Project (CSP) pipeline outlined in Contreras et al. (2010) and Krisciunas et al. (2017). Briefly, the raw data are bias and flat-field-corrected based on nightly calibrations. A world coordinate system (WCS) plate solution is computed using the Refcat2 catalogue of standards (Tonry et al. 2018). Being a follow-up program rather than a discovery survey, POISE does not have reference images with which to do host galaxy subtractions. Instead, public data from Pan-STARRS (Kaiser et al. 2002) or SkyMapper (Wolf et al. 2018) are used for preliminary photometry. In the case of SN 2021gno, the host galaxy light is negligible and mismatches between these survey filters and our Swope filters do not introduce significant errors. Photometric calibration is done by observing Landolt (1992) standards for *BV* and Smith et al. (2002) standards for *ugri*. Using these standards, we computed colour terms that transform the instrumental magnitudes to the standard systems (see Krisciunas et al. 2017 for the colour term equations and coefficients as well as the atmospheric extinction corrections). After correcting for atmospheric extinction, the colour terms are used in reverse to transform the Landolt (1992) and Smith et al. (2002) standard magnitudes into Swope natural system magnitudes. These natural magnitudes are then used to calibrate local sequence stars in the field of SN 2021gno (see Fig. 1). Finally, the calibrated natural magnitudes of the local sequence stars are used to calibrate the relative photometry of SN 2021gno as a function of time. The net result is that the POISE light curves are in the CSP natural system for the Swope telescope. Optical photometry in the Swope natural system is listed in Table A1.

Once POISE observations were finished, we followed SN 2021gno as part of the Aarhus-Barcelona cosmic FLOWS project² using the Las Cumbres Observatory Global Telescope (LCOGT; Brown et al. 2013) network of 1 m telescopes equipped with the Sinistro cameras and the *BVgri* filters. These data were calibrated using the *BVgri* local sequence magnitudes in the CSP natural system from the Swope observations. LCOGT data are listed in Table A3.

SN 2021gno was also observed in the *BgVriz* bands with two identical 0.8 m telescopes at the Baja Observatory and Konkoly Observatory, both located in Hungary. Image subtraction was performed using a late-time template image obtained at JD = 2459586.6 (i.e. 293.81 d after the explosion) and the photometric calibration was based on field stars in the Pan-STARRS DR1 catalogue (Chambers et al. 2016). Photometry from Baja and Konkoly Observatories is listed in Table A2.

Additionally, space-based observations were triggered with the Ultraviolet/Optical Telescope (UVOT; Roming et al. 2005) at the *Neil Gehrels Swift Observatory* (Gehrels et al. 2004), starting on 2021 March 20 (JD = 2459294.044). The images were extracted from NASA's High Energy Astrophysics Science Archive Research Center (HEASARC³), in *w2*, *m2*, *w1*, *u*, *b*, and *v* filters. Aperture photometry was performed following the procedures in Brown et al. (2009), with a 3 arcsec aperture, subtracting the galaxy count rate measured in a 3 arcsec aperture on observations from 2021 July 19. The magnitudes were calculated by applying a re-computed aperture

correction and using the zero-points from Breeveld et al. (2011) in the *Swift* system. The final Vega magnitudes are listed in Table A4. The *Neil Gehrels Swift Observatory* also triggered the X-Ray Telescope (Burrows et al. 2005) and detected a bright X-ray emission at ≈ 0.3 d after the discovery (Jacobson-Galán et al. 2022). An analysis of the X-ray observations is beyond the scope of this paper, but see Jacobson-Galán et al. (2022) for more details.

The resulting UV and optical light curves of SN 2021gno are shown in Fig. 2.

2.2 Spectroscopy

Spectroscopic observations started soon after discovery, at approximately 4 d relative to the time of the explosion. The observations continued until 2021 July 17, covering 116 d. The log of spectroscopic observations is listed in Table 1.

Three spectra were taken before the second maximum with the SPectrograph for the Rapid Acquisition of Transients (SPRAT; Piascik et al. 2014) on the 2.0 m Liverpool Telescope at Observatorio del Roque de los Muchachos, and with the Wide-Field Spectrograph (WiFeS; Dopita et al. 2007) on the Australian National University 2.3 m telescope at the Siding Spring Observatory. After maximum light, spectra were taken with the following telescopes/instruments: (1) the ESO Faint Object Spectrograph and Camera (EFOSC2; Buzzoni et al. 1984) on the 3.58 m New Technology Telescope (NTT) at La Silla Observatory, within the Public European Southern Observatory Spectroscopic Survey of Transient Objects (ePESSTO +) collaboration (Smartt et al. 2015); (2) the Alhambra Faint Object Spectrograph and Camera (ALFOSC) on the 2.5 m Nordic Optical Telescope (NOT) at Observatorio del Roque de los Muchachos, which were obtained in collaboration with the Nordic optical telescope Unbiased Transient Survey 2 (NUTS2⁴); (3) the Supernova Integral Field Spectrograph (SNIFS; Lantz et al. 2004) on the University of Hawaii 2.2 m Telescope (UH2.2m) at Mauna Kea, as a part of the SCAT survey (Tucker et al. 2022); (4) the Dual Imaging Spectrograph (DIS) on the 3.5 m Telescope at the Apache Point Observatory⁵; and (5) the Ohio State Multi-Object Spectrograph (OSMOS) on the 2.4 m Hiltner telescope at the MDM Observatory. Two nebular spectra were obtained on July 8 and 17, with the Optical System for Imaging and low-Intermediate-Resolution Integrated Spectroscopy (OSIRIS) on the Gran Telescopio Canarias at Observatorio del Roque de los Muchachos, and with the FOcal Reducer/low dispersion Spectrograph 2 (FORS2; Appenzeller et al. 1998), mounted on the Very Large Telescope at Paranal Observatory, as part of the FORS + Survey of Supernovae in Late Times program (FOSSIL, Kuncarayakti et al. in prep.).

We reduced the EFOSC2 spectra with the PESSTO pipeline (Smartt et al. 2015). DIS spectra were reduced using IRAF (Tody 1986), including standard methods for bias and flat-field corrections, with flux calibrations based on standard stars observed at a similar airmass and during the same night as the SN, and cosmic ray removal performed with via Lacosmic (van Dokkum 2001). The 2D SNIFS frames were pre-processed and extracted into a 3D data cube using the methods of Bacon et al. (2001). Then, a custom Python pipeline performed aperture photometry on the 3D data cube using a wavelength-dependent trace and the extracted spectra were placed on a relative flux scale using observations of spectrophotometric standard stars (Tucker et al. 2022). The OSMOS

²<https://flows.phys.au.dk/>

³<https://heasarc.gsfc.nasa.gov/>

⁴<https://nuts.sn.ie/>

⁵Owned and operated by the Astrophysical Research Consortium.

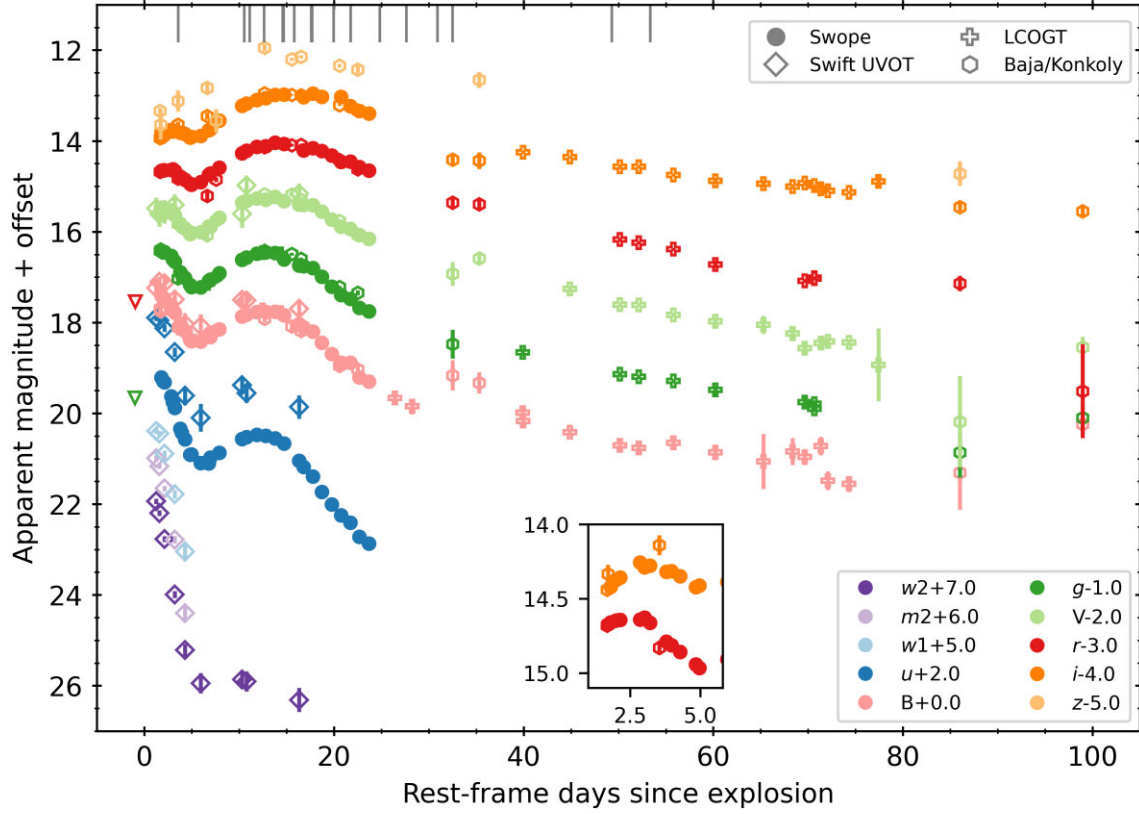


Figure 2. Optical and UV light curves of SN 201gno, corrected for Galactic extinction. For clarity, the light curves are shifted by the offsets indicated in the lower legend. Different instruments are indicated as different markers. The filled symbols are used for Swope photometry and open symbols for other telescopes. Although they are plotted with the same colours, note that the transmission functions of the Swift *ubv* bands, specially *u*, do not match those of the ground-based *uBV* bands. ZTF non-detections are indicated as triangles. The inset plot shows the period around the first peak of the light curve, in *r* and *i* bands. Rest-frame epochs of optical spectra are marked as grey lines along the top axis.

Table 1. Log of spectroscopic observations. The phase is indicated in rest-frame days from explosion.

Date	JD	Telescope	Instrument	Phase
2021 Mar 22	2459296.44	LT	SPRAT	3.6
2021 Mar 29	2459303.437	LT	SPRAT	10.6
2021 Mar 30	2459304.071	ANU 2.3-m	WiFes	11.2
2021 Mar 31	2459305.493	LT	SPRAT	12.6
2021 Mar 31	2459305.540	LT	SPRAT	12.7
2021 Apr 03	2459307.568	NOT	ALFOSC	14.7
2021 Apr 03	2459307.572	NTT	EFOSC	14.7
2021 Apr 03	2459307.602	NTT	EFOSC	14.7
2021 Apr 04	2459308.763	UH88	SNIFS	15.9
2021 Apr 05	2459310.489	NOT	ALFOSC	17.6
2021 Apr 06	2459310.655	APO 3.5-m	DIS	17.8
2021 Apr 08	2459312.917	UH88	SNIFS	20.0
2021 Apr 10	2459314.714	NTT	EFOSC	21.8
2021 Apr 13	2459317.863	APO 3.5-m	DIS	24.8
2021 Apr 16	2459320.672	MDM 2.4-m	OSMOS	27.7
2021 Apr 19	2459323.953	UH88	SNIFS	31.0
2021 Apr 21	2459325.539	NOT	ALFOSC	32.5
2021 May 06	2459342.479	NOT	ALFOSC	49.4
2021 May 11	2459345.718	MDM 2.4-m	OSMOS	52.6
2021 May 12	2459346.584	NTT	EFOSC	53.5
2021 July 08	2459404.425	GTC	OSIRIS	110.9
2021 July 17	2459413.478	VLT	FORS2	119.9

data were reduced with PyRAF-based SimSpec⁶ pipeline. The SPRAT spectra were reduced using the standard liverpool telescope pipeline.⁷ The ALFOSC and GTC spectra were reduced using the dedicated pipeline FOSCGUI⁸, which follows the standard procedures. This includes bias subtraction, flat-field correction, 1D extraction, and wavelength and flux calibration. The FORS2 spectrum was reduced using the ESOReflex (Freudling et al. 2013) pipeline following standard procedures, and using observations of spectrophotometric standard stars taken under the same grism setting.

We have also included three public spectra of SN 201gno available from the Weizmann Interactive Supernova Data Repository (WISeREP⁹; Yaron & Gal-Yam 2012). The spectral sequence of SN 201gno is shown in Fig. 3. The two nebular spectra of SN 201gno are shown in Fig. 4.

⁶<https://astro.subhashbose.com/simspec/>

⁷<https://telescope.livjm.ac.uk/>

⁸FOSCGUI is a Python-based graphic user interface developed by E. Cappellaro and aimed at extracting supernova spectroscopy and photometry obtained with FOSS-like instruments. A package description can be found at <http://sngroup.oapd.inaf.it/foscgui.html>.

⁹<https://wiserep.weizmann.ac.il>

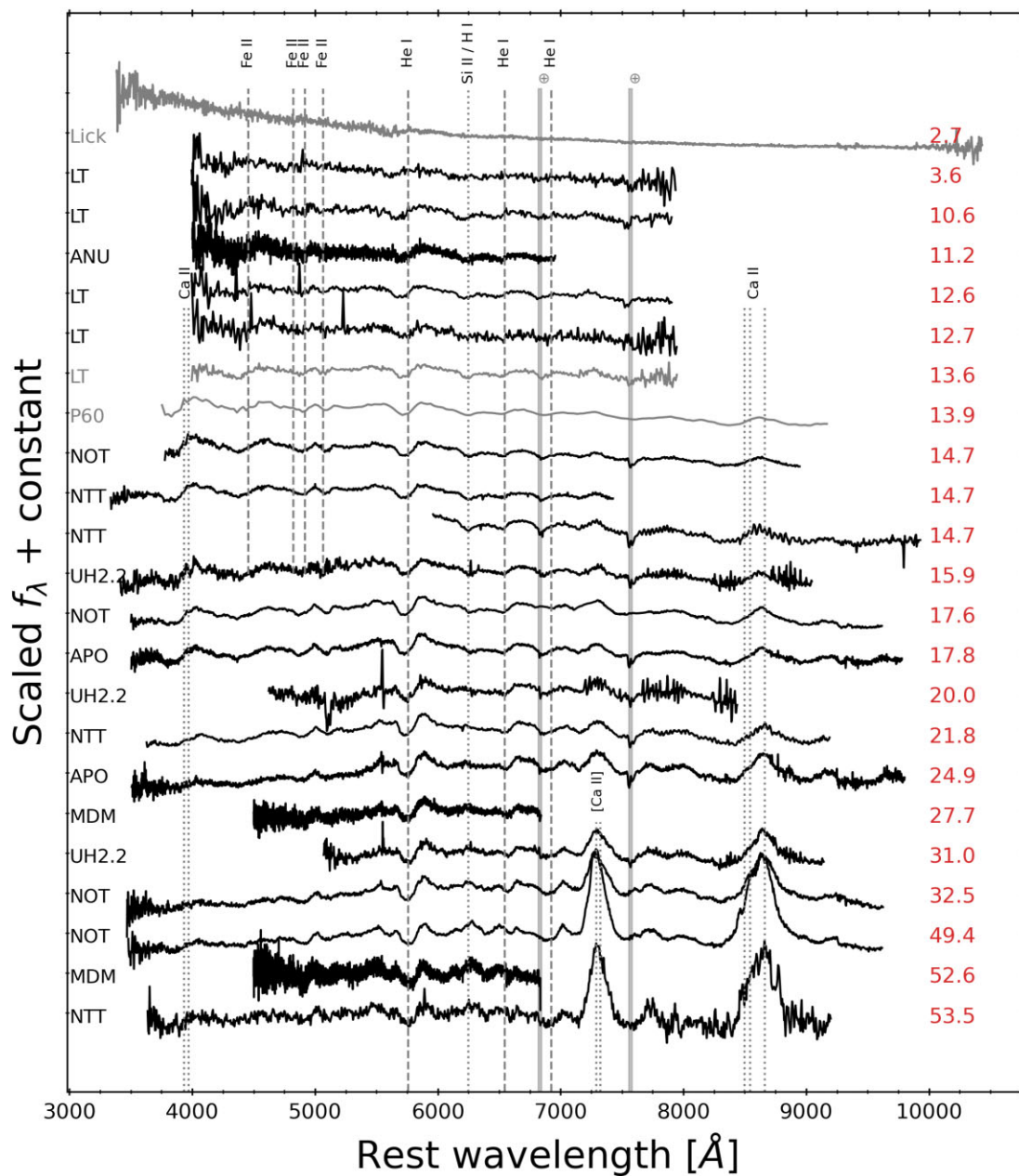


Figure 3. Spectroscopic evolution of SN 2021gno. Spectra taken by the POISE collaboration are plotted in black, and the three public spectra from TNS are plotted in grey. The phase in rest-frame days since explosion is given on the right-hand side of each spectrum, and the labels on the left-hand side indicate the telescope. Main optical lines at a fixed expansion velocity of 6000 km s^{-1} are marked with dashed grey vertical lines, while the grey bands mark the location of telluric absorptions.

3 PROPERTIES

3.1 Distance and extinction

The galaxy NGC 4165 also hosted the Type Ia SN 1971G, enabling a precise redshift-independent distance estimation. This was done by Mueller & Hoefflich (1994), who derived a distance of 34.7 Mpc using models of detonation, delayed detonation, and deflagration. In addition, Hoefflich & Khokhlov (1996) obtained a value of 36 Mpc using three different models of delayed detonation and one of deflagration. The distance was also estimated using the Tully–Fisher method by several authors, as listed in the NASA Extragalactic

Database (NED¹⁰). In particular, Yasuda, Fukugita & Okamura (1997) obtained a value of 57.5 Mpc, and Theureau et al. (2007) obtained 57.6 Mpc using *J*-band data, and 59.9 Mpc using *H*-band data.¹¹ Additionally, NGC 4165 is possible member of a galaxy group that is included in the Cosmicflows-2 catalogue (Tully et al. 2013) whose brightest member is NGC 4168. Tully et al. (2013) give a

¹⁰<https://ned.ipac.caltech.edu/>

¹¹By mistake the NED lists an additional distance from Theureau et al. (2007) with the note ‘mean’, but this is a redshift-dependent measurement and should not be considered as a Tully–Fisher estimation.

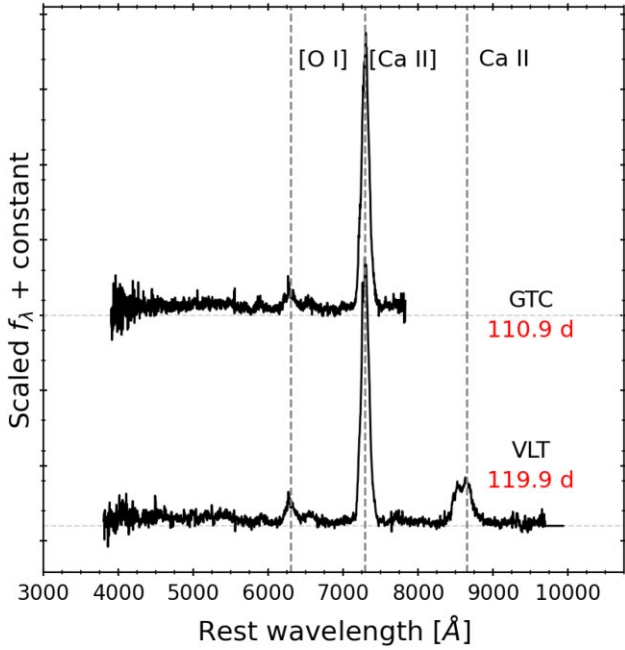


Figure 4. Nebular spectra of SN 2021gno. The labels on the right-hand side of each spectrum indicate the phase with respect to the explosion epoch (in red), and the telescope (in black). Main lines are marked in the rest frame with grey dashed vertical lines. The dashed horizontal lines mark zero flux.

distance to the group of 34.5 ± 1.3 Mpc (for $H_0 = 74.4$ km s $^{-1}$ Mpc $^{-1}$).

Given the large discrepancies among previous distance estimations, we recalculated it using the available *UBV* photometry of SN 1971G (Cadonau & Leibundgut 1990, and references therein) and the Type Ia SN light-curve fitting package *SNOOpy* (Burns et al. 2011), with the `EBV2_model` option. *SNOOpy* adopts $H_0 = 72$ km s $^{-1}$ Mpc $^{-1}$, $\Omega_M = 0.28$, and $\Omega_\Lambda = 0.72$ (Spergel et al. 2007). Although there is pre-maximum *B*-band photometry (as early as -17 d), none of the observations cover the time of maximum light. This, along with the heterogeneous origin of the (pre-CCD) photometry, may be the reason why the *SNOOpy* fits yield a negative $E(B - V)$ colour excess value. We considered such fits as unphysical and decided to fix the extinction to zero (consistent with the analysis below). With such a constraint the resulting distance modulus was $\mu = 32.92 \pm 0.10$ mag, which corresponds to a distance of 38.4 ± 1.8 Mpc. That is the distance value that we adopt throughout this paper. Note that this distance is consistent with the *Cosmicflows-2* catalogue distance for the NGC4168 group.

Regarding the extinction, for the Milky-Way (MW) component we adopted a value of $E(B - V) = 0.03$ mag from the infrared dust maps (Schlafly & Finkbeiner 2011) available from NED, and an extinction law from Cardelli, Clayton & Mathis (1989) with $R_V = 3.1$. To determine any possible host-galaxy extinction, we first examined the spectra for signs of Na I D absorption. We found no evidence of such an absorption at the redshift of the host galaxy in any of the available spectra. This suggests a smaller extinction component from the host than from the MW. Thus we consider the host galaxy extinction to be negligible for the rest of the analysis.

3.2 Light curves

As seen in Fig. 2, SN 2021gno presents a double-peaked optical light curve. This property is shared with a small number of Ca-rich

Table 2. Light-curve properties of SN 2021gno.

Filter	JD _{max} −2459000	t_r^a [d]	Main peak		Decl. rate [mag d $^{-1}$]
			m_{\max} [mag]	M_{\max}^b [mag]	
<i>B</i>	305.59	12.7	17.92 ± 0.09	−14.99	0.142 ± 0.009
<i>V</i>	306.51	13.6	17.32 ± 0.05	−15.60	0.091 ± 0.007
<i>u</i>	305.08	12.2	18.62 ± 0.04	−14.29	0.233 ± 0.008
<i>g</i>	305.92	13.0	17.56 ± 0.09	−15.35	0.121 ± 0.008
<i>r</i>	307.25	14.4	17.14 ± 0.06	−15.77	0.073 ± 0.005
<i>i</i>	307.72	14.8	17.03 ± 0.05	−15.88	0.052 ± 0.008
First peak					
<i>w2</i>	–	–	–	–	1.08 ± 0.02
<i>m2</i>	–	–	–	–	1.11 ± 0.08
<i>w1</i>	–	–	–	–	0.88 ± 0.07
<i>B</i>	–	–	–	–	0.25 ± 0.03
<i>V</i>	–	–	–	–	0.16 ± 0.01
<i>u</i>	–	–	–	–	0.56 ± 0.02
<i>g</i>	–	–	–	–	0.25 ± 0.02
<i>r</i>	295.30	2.5	17.70 ± 0.01	−15.21	0.12 ± 0.02
<i>i</i>	295.50	2.7	17.80 ± 0.01	−15.11	0.08 ± 0.03

Note. ^aThe uncertainty for all rise times is taken as 0.9 d, such as the error in the explosion epoch, since it is the main source of error. ^bThe uncertainty in M_{\max} can be computed as the sum in quadrature of the errors in m_{\max} and the estimated 0.1 mag error in the distance.

transients, the others being iPTF16ghs, SN 2018lqo, SN 2019ehk, and SN 2021inl (De et al. 2018a, 2020; Jacobson-Galán et al. 2020b, 2022). The rise to the first peak in SN 2021gno is observed in the *r* and *i* bands, which is also seen in SN 2019ehk. In the rest of the bands the data show an initial decline, meaning that the first peak possibly occurred in those bands prior to the start of observations.

We first analysed the second peak, which we will call the main peak. We fit the light curves using a low-order polynomial in order to get peak magnitudes and rise times. Due to a gap in observations between 30 and 40 d after the explosion, we used observations taken before 30 d after the explosion for the fits, to avoid large uncertainties introduced by the lack of data. Also, to avoid dispersion in the fits, we excluded the *u*, *b*, and *v* *Swift* bands from our analysis since their effective wavelengths are somewhat different from the standard ones (Poole et al. 2008). The results for all the bands are listed in Table 2. We obtained an absolute peak magnitude in the *r* band of $M_r = -15.77 \pm 0.12$ mag at JD = 2459307.25, giving a rise time of $t_r = 14.46$ d. The uncertainty in the *r*-band absolute peak magnitude is the result of adding the uncertainties in the apparent peak magnitude, and the distance in quadrature. For the *z* band, these parameters were not calculated given the low cadence of the observations. We used the same interpolations of the *r*- and *i*-band light curves in order to estimate the magnitude and rise time of the first peak. The results are listed in the lower panel of Table 2. Our estimated values are compatible with those given by Jacobson-Galán et al. (2022).

We fit straight lines to the light curves (in magnitude scale) in order to obtain initial decline rates after the first peak in all bands. Using the data points between 2.3 and 5 d from explosion, we obtained decline rates for the *r* and *i* band. Assuming the first peak in the other bands takes place before the observations start, we use the data points between the earliest observation and 4.5 d after the explosion in order to fit decline rates for the rest of the bands (excluding the *z* band). The results are given in Table 2. We note that the initial decline is slower the redder the band is. The decline rates after the main peak are also listed in Table 2. The same wavelength dependence as for the initial decline rates is observed.

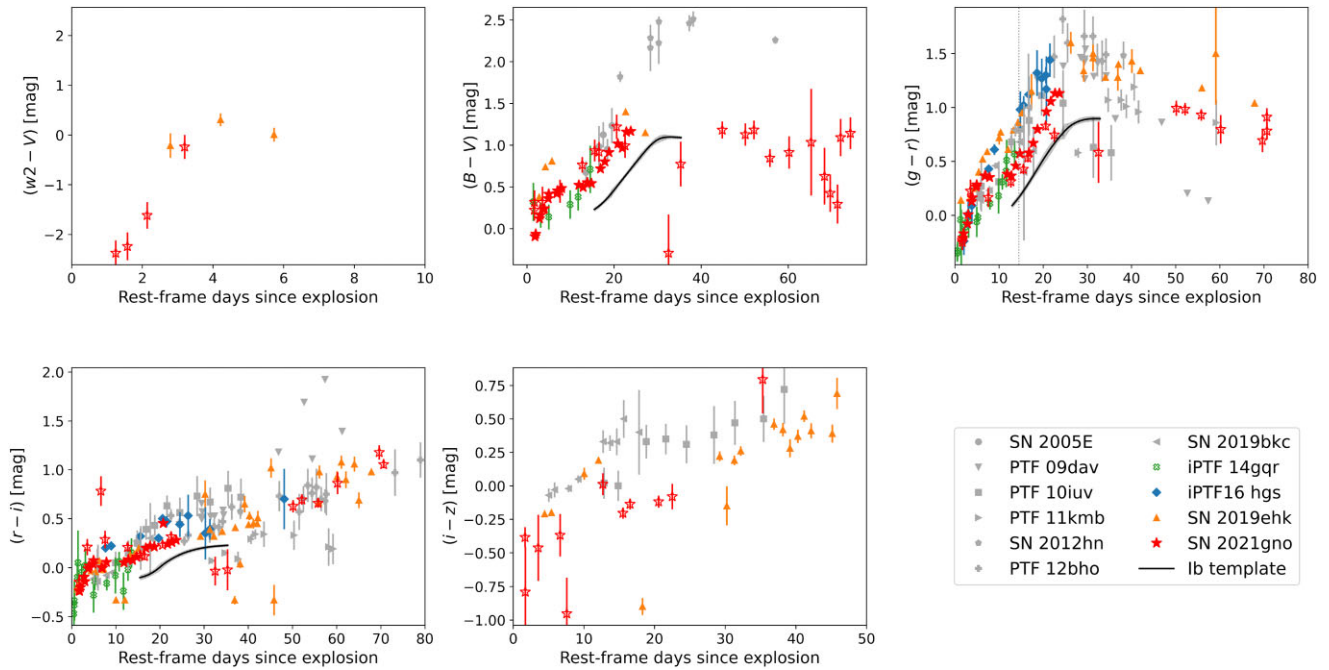


Figure 5. Colour curves of SN 2021gno (red stars) compared to other Ca-rich transients. The filled symbols are used for Swope photometry and open symbols for other telescopes. Ca-rich transients SN 2005E (Perets et al. 2010), PTF09dav (Sullivan et al. 2011), PTF10iuv (Kasliwal et al. 2012), PTF11kmb (Lunnan et al. 2017), SN 2012hn (Valenti et al. 2014), PTF12bho (Lunnan et al. 2017), SN 2016hnk (Jacobson-Galán et al. 2020a), and SN 2019bkc (Prentice et al. 2020) are marked in grey. The double-peaked Ca-rich transients iPTF16hgs (De et al. 2018b), SN 2019ehk (Jacobson-Galán et al. 2020b), and the peculiar double-peaked iPTF14gqr (De et al. 2018a) are marked in blue, orange, and green, respectively. For comparison, SN-Ib colour templates from Stritzinger et al. (2018) are shown as a solid black line. For reference, the epoch of r -band maximum light is marked with dotted lines in the $(g-r)$ panel.

The high cadence in observations of SN 2021gno across all bands allowed us to study its colour evolution compared with a comprehensive sample of Ca-rich transients taken from the literature. The extinction-corrected $(w2 - V)$, $(B - V)$, $(g - r)$, $(r - i)$, and $(i - z)$ colour curves of SN 2021gno are shown in Fig. 5. The comparison sample was constructed by taking all Ca-rich transients with available colours, and their photometry was corrected for Galactic extinction neglecting the host-galaxy component, except for SN 2019ehk, which has substantial extinction from its host (Jacobson-Galán et al. 2020b; Nakaoka et al. 2021). We adopted a value of $E(B - V)_{\text{Host}} = 0.47$ mag from Jacobson-Galán et al. (2020b) for this SN. The colour curves of SN 2021gno are broadly compatible with those of similar objects, although there is a substantial dispersion within the sample. Furthermore, SN 2021gno lies on the ‘blue edge’ of the colour distribution, adding to the conclusion of a negligible host-galaxy extinction component (see Section 3.1).

SN 2021gno quickly evolved from blue to red colours during the first 3 weeks of evolution (through both peaks in the light curve). This suggests a rapid cooling of the ejecta (see Section 4.1). The same behaviour is observed for other double-peaked Ca-rich transients. Afterwards, the colours remained red and nearly constant for the rest of the evolution. As seen from the $(w2 - V)$ panel in Fig. 5, SN 2021gno is the first Ca-rich transient for which this colour is available right after the explosion. The existence of early UV–optical colours is crucial to constrain the temperature during the fast initial evolution (see Section 4.1).

As the majority of Ca-rich transients have spectra similar to type Ib SNe at maximum light, we also show in Fig. 5 the Type Ib intrinsic colour templates of Stritzinger et al. (2018) for $(B - V)$, $(g - r)$, and

$(r - i)$. The colour templates are given for times after maximum light, so we converted them to epochs after explosion using the rise times of SN 2021gno in the corresponding bands. Even if the overall photometric evolution of SN 2021gno is faster than that of typical stripped-envelope (SE) SNe, the rate at which colours change after maximum light is comparable to those of Type Ib SNe, but with a nearly constant shift towards redder colours.

We compare the absolute g - and r -band light curve of SN 2021gno to that of several Ca-rich transients in Fig. 6. Similar to the colour evolution, the criteria for defining the comparison sample was to take the available photometry of all Ca-rich transients from the literature. The main peak of SN 2021gno is comparable in duration and luminosity with those of the other Ca-rich transients. The first peak is similar to those of iPTF16hgs and SN 2019ehk both in terms of duration and slope. The similarity of initial behaviour, both in the light curve and in the colour evolution, among Ca-rich transients with double-peaked light curves may indicate a similar origin for this subgroup of events.

3.3 Spectral evolution

Figs 3 and 4 show the spectral sequence of SN 2021gno, which was obtained between 2.7 and 120 d relative to the time of the explosion. As is typical of Ca-rich SNe, SN 2021gno rapidly evolves to the nebular phase, with forbidden transitions that dominate at about 50 d after explosion (37 d after the main maximum). The last two spectra at 111 and 120 d are clearly nebular, with a very weak continuum and dominated by emission lines mostly from Ca II. In the following sections we describe the properties of the spectra in the photospheric and nebular phases.

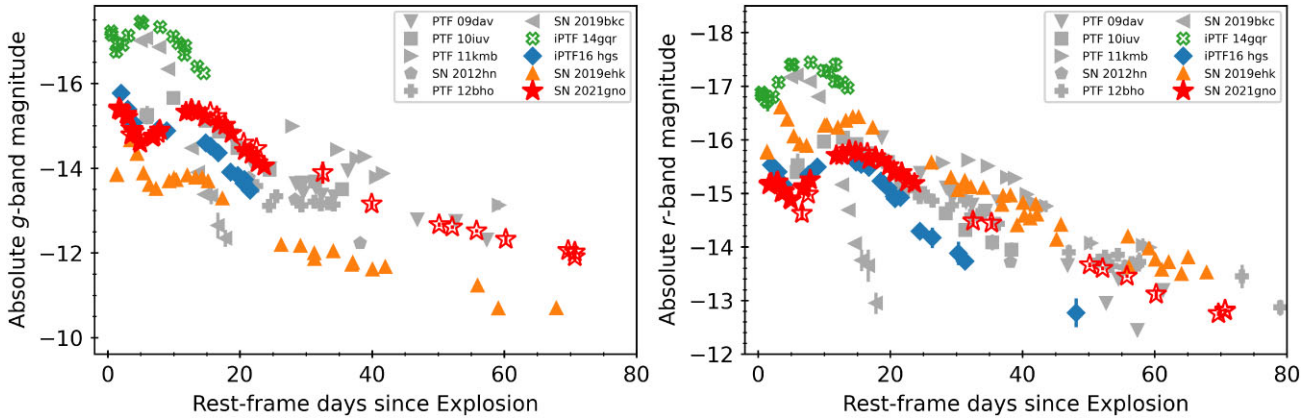


Figure 6. Absolute-magnitude light-curve comparison of SN 2021gno (red stars) and other Ca-rich transients (g band on the left panel, and r band on the right panel). The filled symbols are used for Swope photometry and open symbols for other telescopes. The double-peaked Ca-rich transients iPTF16hgs, SN 2019ehk, and the peculiar double-peaked iPTF14gqr are marked in blue, orange, and green, respectively, whereas objects with singly peaked light curves are shown in grey. References are listed in Fig. 5.

3.3.1 Photospheric phase

After an initial nearly featureless spectrum with a blue continuum at 2.7 d, the spectra of SN 2021gno become dominated by helium lines, most prominently He I λ 5876 (Fig. 3). This continues until about 30 d past explosion. Although this feature could be due, at least in part, to Na I D, the identification as He I is supported by the presence of an absorption that can be identified as He I λ 6678. The He I λ 7065 line may also be present, although it is affected by the telluric H₂O band. The identification is confirmed by our SYNOW analysis (see below).

At times around maximum light, the spectrum resembles those of normal Type Ib SNe. Other spectral features can be identified as due to Ca II $\lambda\lambda$ 8498, 8542, 8662 (the Ca II IR triplet), a set of blended Fe II lines where a weak Fe II λ 5169 absorption is distinguishable, and possibly O I λ 7774. An absorption is also present near 6250 Å, at least until maximum light. This feature usually appears in Type Ib SNe and in Ca-rich transients as well, and it is often associated with Si II λ 6355, but its identification is ambiguous (e.g. see Folatelli et al. 2014).

We measured the expansion velocities of the prominent features in the spectra by fitting a Gaussian to the minimum of the absorption profiles. The velocity evolution of He I λ 5876, He I λ 6678, and Fe II λ 5169 is shown in Fig. 7. For He I λ 5876 the expansion velocity decreases from a value of $\approx 10,500 \pm 120$ km s⁻¹ at 10 d after the explosion to $\approx 8000 \pm 85$ km s⁻¹ at maximum light. After maximum the He I λ 5876 velocity eventually levels off at approximately 6000 ± 90 km s⁻¹ at 25 d from the explosion, whereas the He I λ 6678 continues decreasing. For Fe II λ 5169 the velocities go down from $\approx 5000 \pm 370$ km s⁻¹ around maximum light to $\approx 3700 \pm 90$ km s⁻¹ at 20 d after explosion, when it becomes too weak to be measured.

In order to provide a robust identification of spectral features we calculated synthetic spectra using the SYNOW code (Fisher et al. 1999; Branch et al. 2002) with the aim of reproducing the 14 d after explosion spectrum (i.e. around maximum light). The result is shown in Fig. 8. SYNOW assumes spherical symmetry, homologous expansion, a sharp photosphere at a given expansion velocity v_{ph} ¹² where the continuum emission is defined by a black body (BB)

¹²In this configuration, velocity is a measure of the distance of a given element of ejecta from the explosion centre.

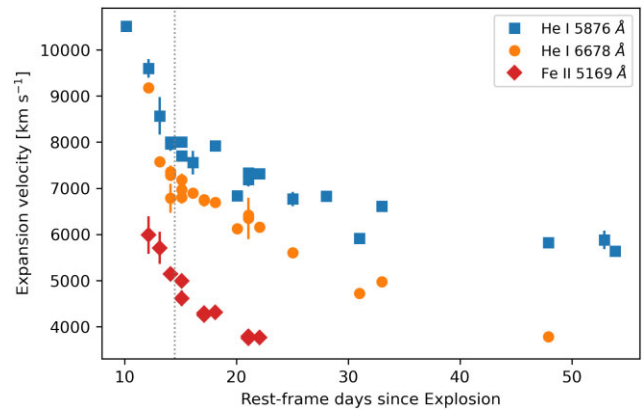


Figure 7. Line velocity evolution of SN 2021gno as derived from the minimum of the absorption components. For reference, the epoch of r -band maximum light is marked with dotted lines.

with temperature T_{BB} , and it treats line formation with the Sobolev approximation. The observed spectrum was best reproduced with $T_{\text{BB}} = 7000$ K and $v_{\text{ph}} = 5000$ km s⁻¹, and including He I, O I, Mg II, Ca II, Ti II, and Fe II in the ejecta (extra species are discussed below to explain the 6250 Å feature). This list of species is not meant to be complete, as other elements or ions may produce identifiable lines. For each species an excitation temperature was adopted close to the assumed value of T_{BB} , except for He I and Ca II, which required temperatures of about 15 000 K in order to reproduce the relative line strengths. Such values should be considered only as a reference and not with strict physical meaning, as deviations from thermal equilibrium and other base assumptions are expected.

The photospheric velocity was suitable to reproduce the location of P-Cygni absorptions of most species assuming a power-law distribution of optical depth as a function of velocity within the ejecta, with power-law indices ranging from 2 to 5 for different species. An exception to this was He I, which was assumed to be distributed in a higher velocity shell detached from the photosphere. This was done by adopting a Gaussian shape for the optical-depth distribution as a function of velocity coordinate, with centre at 7500 km s⁻¹ and width of $\sigma = 3500$ km s⁻¹. Additionally, in order to improve the match on the blue side of the Ca II H&K and IR-triplet features, a high-velocity

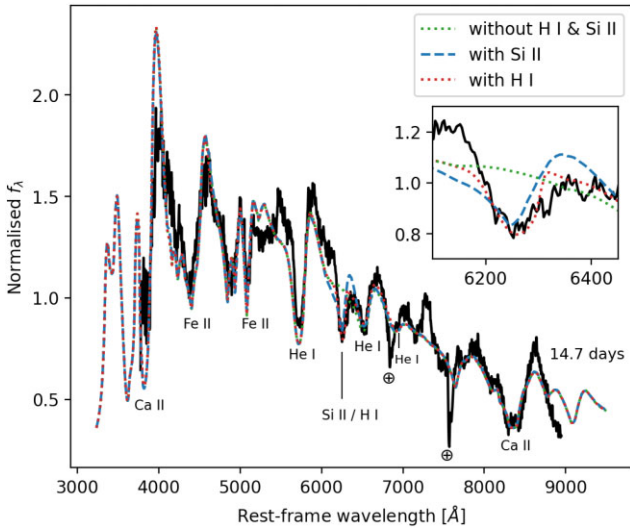


Figure 8. Comparison of the maximum-light spectrum of SN 2021gno at 14.7 d after the explosion (solid black lines) with synthetic spectra calculated with SYNOW. Spectral features are labelled with the main ions that cause them. The SYNOW fit including Si II and without hydrogen is plotted as a dashed blue line. An alternative calculation with hydrogen and without Si II is shown with dotted red lines. For completeness, a model with no H α and no Si II is shown (dotted green lines). The inset panel shows the detail around the observed feature at 6250 Å, in which we indicated the Si II λ 6355 line at photospheric velocity and the H α line detached at 15 000 km s⁻¹.

(HV) component was included for Ca II, with a Gaussian distribution centred at 13 000 km s⁻¹ and with $\sigma = 4000$ km s⁻¹.

With the aid of SYNOW we searched for species that could explain the absorption feature observed near 6250 Å. We found a reasonable match by including Si II with a similar distribution in velocity as the bulk of the elements. However, as shown in the inset of Fig. 8, the Si II absorption at photospheric velocity appears slightly shifted to the blue of the observed line, although the discrepancy (of ≈ 500 km s⁻¹) is not substantial enough to discard Si II. Unfortunately, other Si II transitions are blended and too weak to be unambiguously identified. This prevents a definitive conclusion to be driven as to the presence of Si II.

An alternative identification could be H α arising from a detached distribution of H I with a central velocity of 15 000 km s⁻¹ and a width of $\sigma = 2500$ km s⁻¹ (also shown in the inset of Fig. 8). This is at a slightly higher velocity than that of the HV Ca II component. The subsequent disappearance of the 6250 Å feature would indicate that, if present, the amount and surface abundance of hydrogen in the outer part of the ejecta should be small (H masses of $\sim 10^{-3}$ – $10^{-2} M_{\odot}$; Dessart et al. 2011; Hachinger et al. 2012). Again, it is not possible to unambiguously identify other H I lines in the spectrum. Therefore, the identification of this line remains uncertain, with the possibility of it being due to Si II near the photospheric velocity, H α at high velocity, or a combination of both. In Section 3.3.2, we examine the possible detection of H α emission in the nebular spectra of SN 2021gno, and in Section 5 we further discuss the implications of this possibility for the progenitor scenario.

In Fig. 9 (left panel) the spectrum of SN 2021gno at maximum light is compared with those of other Ca-rich transients at a similar epoch, such as PTF10iuv, and the double-peaked events iPTF16hgs, and SN 2019ehk. As SN 2019ehk has substantial extinction from its host galaxy, we de-reddened its spectrum adopting a value of $E(B - V)_{\text{Host}} = 0.47$ mag from Jacobson-Galán et al. (2020b), and using the

parametrization of Fitzpatrick (1999). Given that Ca-rich transients are similar to type Ib/IIb SNe at this epoch, we have included the Type IIb SN 2011dh and Type Ib SN 2009jf at maximum light as they are representative within their respective classes and they have well-sampled spectra near maximum light. Note, however, that the epoch relative to the explosion in those cases is substantially larger due to their slower evolution. SN 2021gno looks similar to the other Ca-rich transients, particularly in the strength of the He I features, which is not greatly different from standard Type Ib spectra. We note that there is an absorption commonly present at around 6200 Å in Ca-rich transients that may be related with the ≈ 6250 Å feature discussed above.

3.3.2 Pre-nebular and nebular phase

In this section we study the spectra obtained after 30 d relative to the explosion time, when nebular transitions start to appear (see Figs 3 and 4). At these epochs the He I features decrease in strength, while Ca II and [Ca II] emissions become dominant. A weak [O I] $\lambda\lambda$ 6300, 6364 emission feature can also be seen in the last two spectra obtained after 110 d (see Fig. 4). He I λ 5876 may still be detectable in the last two spectra, although it may be blended with the Na I D doublet.

In Fig. 9 (right panel) we compare the spectrum of SN 2021gno entering the nebular phase with those of three other Ca-rich transients observed at similar epochs. SN 2021gno looks similar to the Ca-rich transients both in terms of the Ca II and [Ca II] emission and in the weak [O I] emission. Except for the strength of the [O I] $\lambda\lambda$ 6300, 6364 feature, the ≈ 50 d spectra of the Ca-rich objects resemble those of normal Type Ib and Type IIb SNe at a much later age of ≈ 150 – 200 d, i.e. during the nebular phase.

A weak emission is present at around 6550 Å in the spectra obtained after 110 d (Fig. 4). Its identification is also ambiguous, as it could be produced by H α or by Ca I] λ 6572. [N II] $\lambda\lambda$ 6548, 6583 may also contribute to this feature as proposed by Jerkstrand et al. (2015) (see also Fang & Maeda 2018, for a detailed analysis of this structure). We measured a central wavelength of ≈ 6540 Å, which would correspond to velocity shifts of ≈ -1000 km s⁻¹, ≈ -1500 km s⁻¹ or ≈ -2000 km s⁻¹ if it were due to H α , Ca I], or [N II] (relative to the strongest component at 6583 Å), respectively.¹³ The first shift is similar to what is seen for the [Ca II] $\lambda\lambda$ 7291, 7324 feature (see below), so this favours the identification as H α . On the other hand, the lack of noticeable H β and the presence of strong [Ca II] lines (Fig. 4) provide support to the association with Ca I] λ 6572. In Section 5 we further discuss the implications of the possible presence of hydrogen in the late-time spectra of some Ca-rich transients.

In Fig. 10 we inspect the late-time profiles in velocity space of the [O I] $\lambda\lambda$ 6300, 6364 and the [Ca II] $\lambda\lambda$ 7291, 7324 emissions. The [O I] $\lambda\lambda$ 6300, 6364 feature is compatible with a double-peaked shape (Mazzali et al. 2005), which has been observed in nebular spectra of SE SNe (e.g. Prentice & Mazzali 2017). A slight overall blueshift is seen with respect to the reference wavelength of 6300 Å.¹⁴ The blue peak is centred roughly at -1300 km s⁻¹ while the red peak is centred at ≈ 800 km s⁻¹. The separation between both peaks is substantially smaller than the one expected (of ≈ 3000 km s⁻¹) between the line components. Therefore, the shape of this line is suggestive of an asymmetric distribution of the O-rich material, such

¹³The OSIRIS spectrum in the observed configuration provides a resolution of ≈ 300 km s⁻¹.

¹⁴The shift would be slightly larger if the λ 6364 component influenced the actual reference wavelength of this blend.

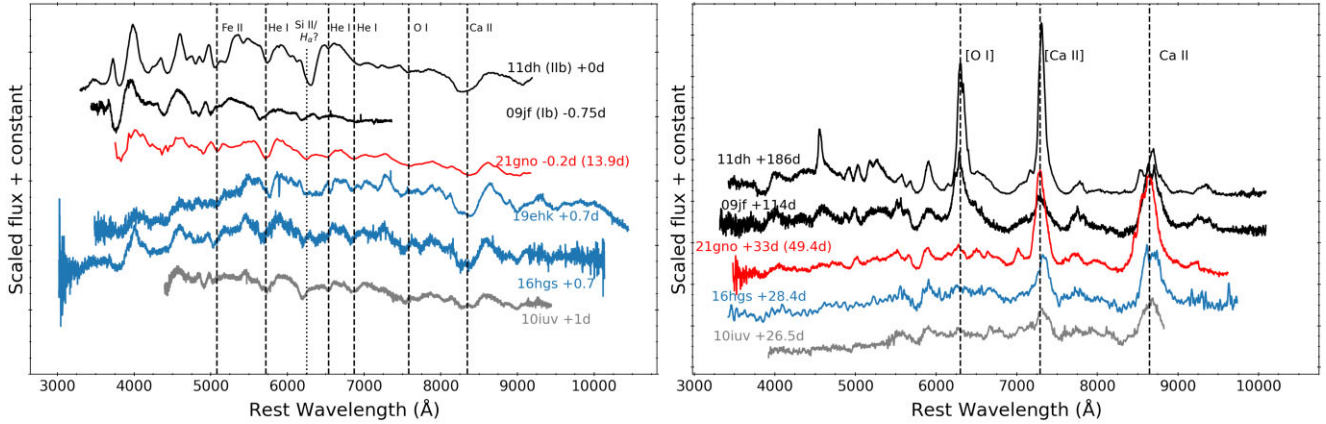


Figure 9. Left: SN 2021gno spectral comparison at maximum light (red). The Ca-rich transient PTF10iuv (Kasliwal et al. 2012) is marked in grey; the double-peaked Ca-rich transients iPTF16hgs (De et al. 2018b), and SN2019ehk (De et al. 2021) are marked in blue. For comparison, the prototypical Type IIb SN2011dh (Ergon et al. 2014) and Type Ib SN 2009jf (Modjaz et al. 2014) are shown in black. Epochs are in rest-frame days since maximum light (for reference they are given relative to the explosion time in parentheses for SN 2021gno). The dashed vertical lines correspond to the main optical lines marked at the absorption minimum in the spectrum of SN 2021gno. The dotted vertical line is located at 6250 Å (see Section 3.3.1). Right: Spectral comparison of SN 2021gno at early nebular phase (in red). PTF10iuv is marked in grey, and the double-peaked Ca-rich transient iPTF16hgs is marked in blue (references are the same as in the left panel). Type IIb 2011dh (Shivvers et al. 2013) and Type Ib SN 2009jf (Shivvers et al. 2019) spectra are marked in black.

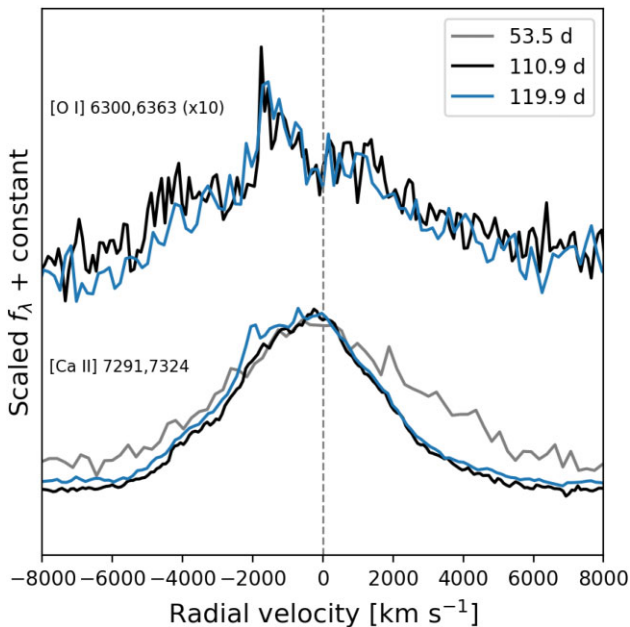


Figure 10. [Ca II] and [O I] emission lines observed in the transitional and nebular spectra of SN 2021gno, plotted in velocity space. The spectra correspond to 53.5 d (grey), 110.9 d (black), and 119.9 d (blue) relative to the explosion epoch. Reference wavelengths were assumed as 6300 Å for the [O I] $\lambda\lambda 6300,6363$ complex, and 7304.4 Å for the [Ca II] $\lambda\lambda 7291,7324$ doublet. The [O I] line of the 53.5-d spectrum is not included since it is not visible at that epoch.

as in a toroidal or a bipolar structure (Maeda et al. 2008; Tanaka et al. 2009). However the S/N is low, so this profile may be compatible with a shell-like structure result of the lack of oxygen in the innermost ejecta (Taubenberger et al. 2009; Mazzali et al. 2005, 2017).

The [Ca II] line is instead singly peaked, although it also appears to be blueshifted. Relative to the adopted effective central wavelength of 7304.4 Å, the blueshift goes from ≈ -50 km s $^{-1}$ to ≈ -500 km s $^{-1}$ between 50 and 120 d past explosion. During the same time lapse, the

line profile becomes asymmetric. As shown in Fig. 10, the red side of the line at 111 and 120 d appears to be suppressed as compared with the 53-d spectrum. On the contrary, the shape of the line on the blue side remains nearly constant.¹⁵

There are several possible explanations for the observed evolution of the [Ca II] line. It could be due to changes in the timing of the emergence of contaminating lines, possibly from iron-peak elements (Maeda et al. 2008; Jerkstrand et al. 2015; Dessart et al. 2021). However, in order to maintain the blue side of the line in a highly coincidental fashion. Thus we consider this to be an unlikely possibility. Another explanation would be the formation of dust at some time between 50 and 110 d (see for example, Silverman et al. 2013). Dust would absorb the light from the far side of the ejecta and thus would reduce the emission on the red side of the line. The amount of dust would increase with time, causing an increasing blueshift of the line, which is observed (Taubenberger et al. 2009). However, one would expect an increase of overall extinction and reddening at late times and therefore an increase of the light-curve decline slopes, especially in the bluest bands. This is not apparent in the observed light curves, at least until ≈ 80 d past explosion. Dust formation is thus unlikely the cause of the [Ca II] line evolution. Nevertheless, some mechanism producing self-absorption in a clumpy ejecta may reduce the flux on the red side of the line without a noticeable effect on the light curves and colours (Wang & Hu 1994). Self-absorption and scattering in the ejecta may also produce blueshifted, asymmetric lines, but the effect should decrease with time as the density decreases (Jerkstrand 2017). The reduction in density is supported by the decrease in flux of permitted features (e.g. the Ca II IR triplet) relative to forbidden transitions ([Ca II] $\lambda\lambda 7291,7324$), as can be seen in Figs 4 and 9. Thus the line should become more symmetric at later epochs, contrary to what we observe. A final alternative would be an asymmetric distribution of the Ca-rich material (Maeda et al. 2007), although this requires a

¹⁵Note that the wings of the line at 53 d are affected by the existence of non-negligible continuum emission.

change in the physical distribution of the material that produces the line occurring some time between 50 and 110 d.

For the last two spectra, at 111 and 120 d after the explosion, we measured the $[\text{Ca II}]/[\text{O I}]$ flux ratio, obtaining a mean value of 10.5 ± 0.1 . This is a typical value for Ca-rich transients, as compared with those of normal Type Ib/c SNe that always lie below 2 (Milisavljevic et al. 2017; Fang et al. 2022).

4 BOLOMETRIC EVOLUTION AND MODELLING

To better understand the physics underlying SN 2021gno, we studied the photometric evolution via black body fits to the spectral energy distribution at each epoch. This allowed us to determine the BB temperature and radius evolution and to construct a bolometric light curve (Section 4.1) that we then employed to compare with hydrodynamic models (Section 4.2).

4.1 Bolometric luminosity

After correcting the observed magnitudes for extinction (see Section 3.1), we converted them to monochromatic fluxes at the effective wavelength of each filter, using the transmission functions of the photometric system, taken from the CSP webpage.¹⁶ If a given epoch lacked observations in a certain band, we made interpolations using a Gaussian Processes method with the Python library GPpy.¹⁷ In order to secure a self-consistent coverage in the optical range, we restricted the calculations to the photometry from the Swope telescope at early times, and the LCOGT at late times. We also interpolated the z -band, since it is necessary to obtain reliable BB fits. Then we integrated the monochromatic fluxes along wavelength for each epoch, using the trapezoidal method in order to obtain a quasi-bolometric flux (F_{qbol}). As F_{qbol} is calculated with all observed bands, and the wavelength coverage may vary between epochs, we also calculated the $F_{u \rightarrow z}$ flux by consistently integrating in the range between the u and z bands. Because of the limited u -band coverage, $F_{u \rightarrow z}$ was calculated between 2 and 24 d after the explosion. Both F_{qbol} and $F_{u \rightarrow z}$ are listed in Table A5.

To allow a comparison with the bolometric luminosities produced by our hydrodynamical models (Section 4.2), we account for the flux outside the wavelength range covered by our broad-band photometry by performing ultraviolet and infrared extrapolations using BB fits to the spectral energy distribution at each epoch. At early epochs the BB distribution represents a good description of the SN emission, but as the ejecta expand and cool the emission at shorter wavelengths starts deviating from the BB model because of line blanketing produced by iron-group elements. Including the bluer bands in those cases worsens the BB fits and biases the temperatures towards lower values, which in turn produces an overestimation of the IR correction (Faran, Nakar & Poznanski 2018; Martinez et al. 2022). With the aim of having reliable fits, we followed a procedure similar to that of Faran et al. (2018). When observations were available in either $w2$, $m2$, $w1$, or u bands and the flux in those bands dropped by more than 1σ below the BB model, we calculated it again excluding those bands. This was done while there were observations available in those bands (≈ 16 d after the explosion in the UV, and ≈ 24 d in u band). After that epoch, the fits included the complete set of optical bands.

We note that after around 30 d from the explosion, the presence of strong emission lines on the red part of the spectrum (see Figs 3 and 4) causes the SED to be overestimated in the i and z bands relative to the continuum. This, in turn, would produce an overestimation of the extrapolated IR flux. In order to estimate the size of this effect, we removed both strong calcium emission lines from the available spectra after 30 d past explosion, and we calculated synthetic photometry. The difference between the original and ‘continuum only’ spectra at ≈ 30 d amounted to 0.1 and 0.4 mag in the i and z bands, respectively. At ≈ 50 d, the differences increased to 0.5 and 0.95 mag in the respective bands. To correct the observed i - and z -band photometry we used a straight-line fit to these differences between 30 and 50 d, and a constant value thereafter (as a conservative lower limit to the correction given the lack of spectra until 110 d). This was done solely to produce corrected SEDs to be fit by BB functions for extrapolation, as the integrated fluxes F_{qbol} and $F_{u \rightarrow z}$ were left unmodified.

The infrared flux beyond the z band, F_{IR} , was estimated by integrating the fitted BB function from the effective wavelength of the z band to infinity. To account for the unobserved ultraviolet flux, F_{UV} , when all bands were available for the BB fits, we integrated the BB from 0 \AA to the bluest band observed at that epoch. In the rest of the cases we integrated a linear extrapolation from the effective wavelength of the bluest band to zero flux at 2000 \AA (Bersten & Hamuy 2009).

The total bolometric fluxes were calculated as $F_{\text{bol}} = F_{\text{UV}} + F_{\text{qbol}} + F_{\text{IR}}$ and then converted to luminosity assuming the distance calculated in Section 3.1. The uncertainty in the luminosity was estimated by considering uncertainties in the distance, photometry, and the estimated errors of the extrapolated fluxes. The resulting bolometric light curve is shown on the top panel of Fig. 11. The luminosity drops rapidly during the first ≈ 5 d past explosion before rising to a local maximum that occurs at ≈ 15 d. Due to the lack of z -band data and the uncertainties in the IR extrapolations after about 25 d past explosion, we consider the bolometric light curve to be less reliable after that epoch than at earlier times.

The BB temperature and radius evolution are shown on the middle panel of Fig. 11. At early epochs, the BB temperature drops rapidly, from over 20 000 to ≈ 7000 K in less than 5 d, which is indicative of a fast cooling of the shocked envelope. This fast cooling coincides with the initial light-curve peaks. We note the significance of UV observations to make reliable estimations of the temperature at early times. After that, the temperature shows a slower decline, reaching ≈ 4000 K at times past maximum light. Both behaviours are consistent with the colour evolution of SN 2021gno seen in Fig. 5. At later times (45 to 75 d past explosion), the temperature remains nearly constant, although we note that the BB approximation breaks down as the SN transitions into an emission-line-dominated spectrum.

Similar to the temperature evolution, the BB radius shows a steep initial rise until about Day 5 (with a velocity of $\approx 5000 \text{ km s}^{-1}$), and then slows down to $\approx 3000 \text{ km s}^{-1}$ until about ≈ 25 d after the explosion. Until that time, despite the aforementioned blanketing effects, the BB approximation remains valid and the BB radius roughly follows the photospheric radius. Later on the derived BB radius recedes to a nearly constant value, although this behaviour probably lacks any physical meaning. Although the methods employed in Jacobson-Galán et al. (2022) to compute the bolometric luminosity are slightly different, all the estimated bolometric quantities are similar to our results.

The fractional contributions of the $\text{UV}/F_{\text{qbol}}/\text{IR}$ components to the total flux are depicted on the bottom panel of Fig. 11. The observed

¹⁶<https://csp.obs.carnegiescience.edu/data/filters>

¹⁷<https://github.com/SheffieldML/GPy>

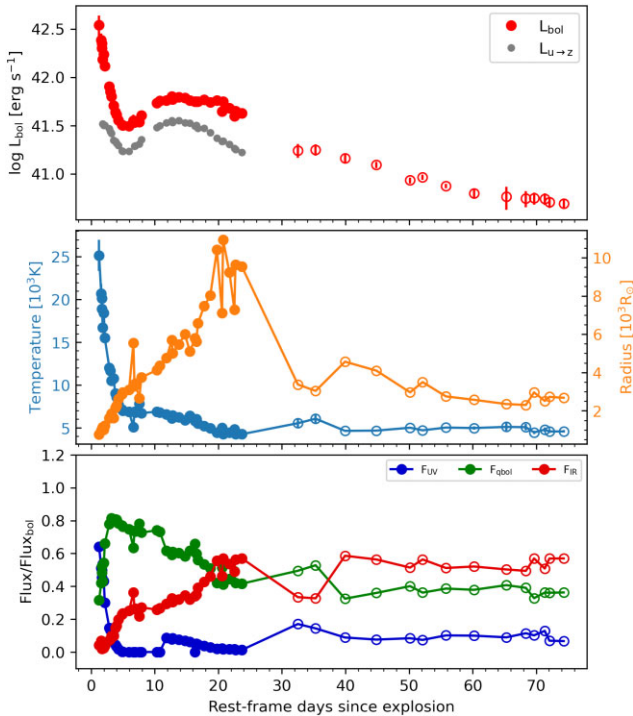


Figure 11. Top: Bolometric light curve of SN 2021gno. The luminosity calculated from $F_{u \rightarrow z}$ is also shown. Middle: Evolution of the fitted BB temperature (blue) and radius (orange) for SN 2021gno. Bottom: Contributions to the total flux. F_{qbol} is the integrated flux in the available bands, while F_{UV} and F_{IR} are the estimated flux from extrapolations to the UV and IR, respectively (see text for details). The open symbols indicate epochs where z band was not available.

contribution (F_{qbol}) dominates between about 3 and 15 d past explosion (i.e. from slightly before the bolometric local minimum until slightly past maximum light). At earlier epochs, when the ejecta are very hot, the extrapolated UV flux is significant and it contributes up to 60 per cent of the total estimated flux. The UV contribution rapidly decreases as the temperature drops. The contribution of the IR extrapolated flux is small at early times but it becomes important soon after and it dominates after maximum light. We estimate it to be ≈ 60 per cent of the total flux at the latest epochs, although as mentioned above, the IR flux may be overestimated.

4.2 Hydrodynamical modelling

We model the bolometric light curve and the line velocity evolution of SN 2021gno with the one-dimensional Lagrangian hydrodynamical code of Bersten, Benvenuto & Hamuy (2011). In the context of a calcium-rich transient, we propose an ad hoc pre-explosion model with a highly stripped star as an initial condition. In order to obtain the initial luminosity peak, the additional presence of circumstellar matter (CSM) is assumed. Even though the model simulation is self-consistent at all epochs, the analysis can be decoupled into two distinct phases: (a) the initial decline due to post-shock cooling of the CSM, and (b) subsequent evolution through the main peak that is powered by radioactive decay. From the former phase we can mainly obtain the mass and extent of the CSM, and from the latter phase one can derive other overall explosion parameters, such as the ejected mass, the explosion energy, and the mass of radioactive material.

As a pre-explosion structure for the main peak of the light curve we construct a highly stripped object by using the public stellar

evolution code MESA¹⁸ version 15140 (Paxton et al. 2011, 2013, 2015, 2018, 2019) with the ‘input files’ of Moriya et al. (2017) modified to comply with the new MESA version and to include a nuclear network of 21 isotopes (approx21.net). The prescription proposed by Moriya et al. (2017) simulates mass-loss from a He star via Roche-lobe overflow (RLOF) in a simplified manner, following the evolutionary calculations by Tauris et al. (2013) of a close binary system comprising a He star that transfers matter to an NS.

According to Tauris et al. (2013), mass transfer by RLOF occurs after core He exhaustion. Therefore, we first evolve an isolated He star of $4 M_{\odot}$ until it finishes burning He in its core. Then, we reduce the mass of the star using the parameter ‘relax mass’, defined by MESA, until the He star reaches a mass of $2.5 M_{\odot}$ and we continue the evolution until core collapse, which is taken as the time when any location inside the stellar model reaches an infall velocity of 1000 km s^{-1} . At the end of its evolution, the star has $2.48 M_{\odot}$ and a CO core mass of $1.89 M_{\odot}$. This core mass would indicate a $M_{\text{ZAMS}} \approx 15\text{--}17 M_{\odot}$ (Ertl et al. 2020), although this may not apply to the atypical evolutionary path considered here. To this compact structure we attach different forms of He-rich CSM with wind-like density distributions of varying mass and radial extent.

To initialize the explosion, some energy is deposited in the form of a thermal bomb at a given mass ‘cut’ within the pre-SN structure. In this case the mass cut was set to $1.7 M_{\odot}$, a value that is somewhat higher than the composition interface between the silicon core and oxygen core (see Morozova, Piro & Valenti 2018), and it was assumed to form a compact remnant. The ejected mass was therefore of $0.8 M_{\odot}$. We varied the explosion energy and ^{56}Ni mass to reproduce the main light-curve peak and the evolution of the Fe II velocity. The resulting preferred model, shown in Fig. 12, has an explosion energy of $E = 0.15 \times 10^{51} \text{ erg}$ and ^{56}Ni mass of $M_{^{56}\text{Ni}} = 0.024 M_{\odot}$. The degree of Ni mixing was fixed to an extension of 95 per cent of the progenitor mass. This parameter affects the timing and depth of the light-curve minimum as more extended nickel produces an earlier rise to the main maximum (Bersten et al. 2012). In order to match the bolometric luminosity after Day 30 we needed to lower the model luminosity by artificially reducing the gamma-ray opacity from the standard value of $\kappa_{\gamma} = 0.03 \text{ cm}^2 \text{ g}^{-1}$ to $\kappa_{\gamma} = 0.01 \text{ cm}^2 \text{ g}^{-1}$. Such an ad hoc modification of the gamma opacity has been used in the literature (Tominaga et al. 2005; Folatelli et al. 2006; Gutiérrez et al. 2021, 2022) to mimic a reduction of the gamma-ray deposition in the ejecta when the nickel mass required to explain the peak luminosity overestimates the tail luminosity. Nevertheless, we note that the late-time bolometric luminosity is less reliable than around maximum light (see Section 4.1). In any case, this modification does not affect the conclusions of our analysis.

We note that the ejected mass derived from the main peak of the light curve is somewhat larger than typical masses assumed for ultra-stripped progenitors, which are $M_{\text{ej}} < 0.2 M_{\odot}$ (Tauris et al. 2015). However, when we adopted lower mass pre-SN models (all starting from the same $4 M_{\odot}$ He star), the resulting light curves were too narrow compared with the observations around the main maximum, as shown in the inset plot of Fig. 12.

Based on the overall explosion parameters found above we focused on fitting the early behaviour as arising from the cooling of the shocked CSM. For this purpose, we modified the external density structure of the progenitor before injecting the explosion energy. We assumed density profiles of the type $\rho \sim r^{-2}$ for the CSM, i.e. due to a steady wind, with different extensions and masses. In Fig. 12,

¹⁸<http://mesa.sourceforge.net/>

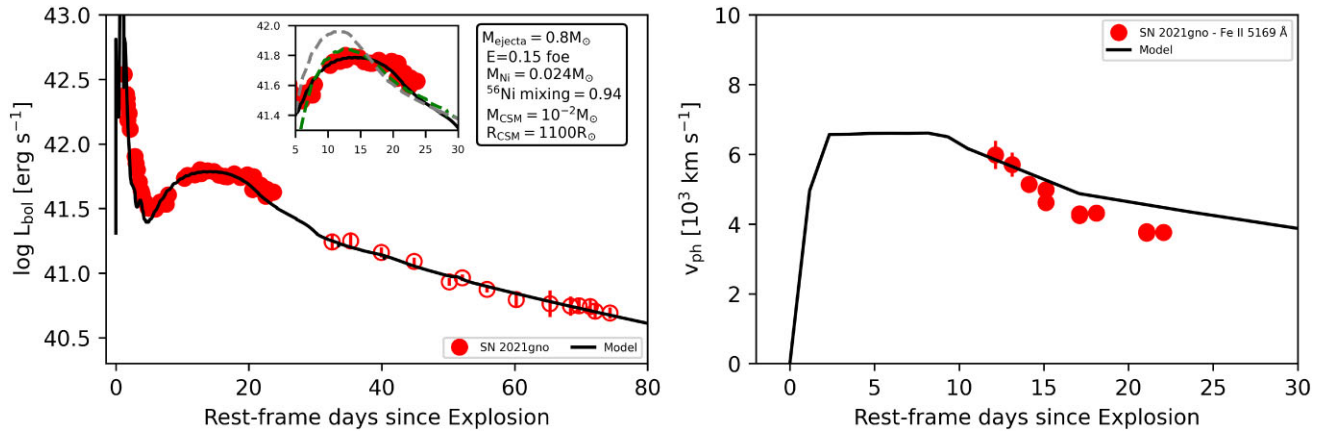


Figure 12. The bolometric light curve of SN 2021gno (red dots, left panel) and Fe II velocity evolution (red dots, right panel), compared with the hydrodynamical model discussed in Section 4.2 (solid black lines). The open symbols indicate epochs where z band was not available. The parameters of the model and the CSM parameters that reproduce the early data are shown in the box. ^{56}Ni mixing is given in fraction of the pre-SN mass. The inset plot in the left panel shows two alternative models (dashed lines) around maximum light for a lower mass pre-SN progenitor of $2.2 M_{\odot}$, with the same conditions as the preferred model (grey), and varying the energy and ^{56}Ni mass (green).

we present a model with a CSM mass of $\sim 10^{-2} M_{\odot}$ and a radius of $1100 R_{\odot}$ that reproduces very well the observations. In the case of a steady wind, this model corresponds to a mass-loss rate of $4 \times 10^{-3} M_{\odot} \text{ yr}^{-1}$. We note that a similar cooling behaviour may be obtained assuming different density distributions of the CSM. For instance, an accelerated wind (Moriya et al. 2017) or an extended envelope in hydrostatic equilibrium were not tested in this case.

The model presented here is not necessarily the only possible scenario for this SN. However, the calculations indicate that this set-up of a highly stripped star with a tenuous material around it is plausible. The parameters quoted here may serve as a comparison for similar objects. The nature of the surrounding material, whether a wind or the result of binary mass transfer or pre-explosion eruptions, is not clear. Nevertheless, the initial decline of the bolometric luminosity and the fast drop of the blackbody temperature (Fig. 11) strongly suggest the presence of such a material.

5 PHYSICAL ORIGIN OF DOUBLE-PEAKED CA-RICH TRANSIENTS

Along with SN 2021gno, other Ca-rich objects that showed double-peaked light curves are iPTF16hgs (De et al. 2018b), SN 2018lqo (De et al. 2020), SN 2019ehk (Jacobson-Galán et al. 2020b; De et al. 2021; Nakaoka et al. 2021), and SN 2020inl (Jacobson-Galán et al. 2022). In Table 3 we summarize the environmental properties of these events. We have included in the table the related transient iPTF14gqr (De et al. 2018a); although it is not generally considered as a Ca-rich event. It should be noted that some Ca-rich transients lack early-time observations, and they may have a first peak that was not observed. Jacobson-Galán et al. (2022) point out that five out of nine (55 per cent) Ca-rich transients discovered less than 3 d after explosion and with < 2 d cadence in the observations show this early peak in the light curve. However, they also open the possibility of marginal detections of very early first peaks for the rest of the sample, meaning that this feature could be present in all Ca-rich transients.

With the exception of SN 2018lqo and SN 2021inl, the rest of the double-peaked Ca-rich transients appeared in environments where ongoing star formation is plausible. This opens the possibility that a subgroup of Ca-rich transients with double-peaked light curves arises

from a younger population than that of WD stars, i.e. this group may originate from CC explosions of massive stars. De et al. (2021) suggested that they can arise from stripped stars at the low-mass end of CCSN progenitors ($8\text{--}10 M_{\odot}$). Nakaoka et al. (2021) also proposed a CC origin for iPTF14gqr, iPTF16hgs, and SN 2019ehk, within the more extreme USSN scenario. For SN 2021gno we presented a progenitor set-up consistent with a highly stripped massive star with an ejecta mass of $0.8 M_{\odot}$ and a ^{56}Ni mass of $M_{^{56}\text{Ni}} = 0.024 M_{\odot}$. Although the ejecta mass is higher than what was estimated for those other objects, the proposed scenario is similar. In conclusion, both WD and massive-star progenitor scenarios may coexist in the case of double-peaked Ca-rich transients.

In the CC scenario the presence of the initial light-curve peaks is explained by the existence of an extended CSM or a thin extended stellar envelope. Furthermore, the detection of X-ray emission in the early light curve of both SN 2019ehk (Jacobson-Galán et al. 2020b) and SN 2021gno (Jacobson-Galán et al. 2022) is consistent with the presence of a shocked CSM. In the case of SN 2021gno, we proposed a CSM containing $\sim 10^{-2} M_{\odot}$ and an extension of $1100 R_{\odot}$.

Double-peaked light curves may also be found in the WD scenario as a consequence of He-shell detonations (Shen et al. 2010). In fact, by modelling X-ray observations of SN 2021gno, and using shock cooling, shock interaction, and radio models, Jacobson-Galán et al. (2022) concluded that the progenitor CSM density is consistent with models for the merger of low-mass, hybrid WDs. They further favour a WD origin for SN 2021gno, contrary to a CC scenario, based on a series of arguments that we address next.

First, Jacobson-Galán et al. (2022) state that ultra-stripped stars produce too low ejecta masses and too little He to explain the observations of SN 2021gno. However, a less extreme progenitor such as the one presented here does produce the necessary ejecta mass (tenths of M_{\odot}) to comply with the light curve and enough He ($\approx 0.5 M_{\odot}$) to match the observed Type-Ib spectrum (see Hachinger et al. 2012). Jacobson-Galán et al. (2022) point out that all of the binary evolution models in the comprehensive set of BPASS (Eldridge et al. 2017) calculations would produce too large ejecta masses for SN 2021gno. Our calculations, however, prove the viability of a low-enough ejecta mass as a consequence of an atypical binary evolution path similar to that of USSN progenitors (Tauris et al. 2013; Moriya & Eldridge 2016).

Table 3. Properties of double-peaked calcium-rich transients.

SN	Host type/ Projected offset from the host nucleus	Star-forming region	Proposed progenitor
iPTF14gqr	Interacting double system/29 kpc	Maybe	-USSN (De et al. 2018a)
iPTF16hgs	Dwarf spiral galaxy ^a /5.9 kpc	Yes	-CCSN from highly stripped massive star in a close binary system (De et al. 2018b) -He-shell detonations on WDs (De et al. 2018b)
SN 2018lqo	E-type galaxy/15.46 kpc	No	-
SN 2019ehk	SAB(s)bc/1.8 kpc	low SFR	-Merger of low-mass WDs (Jacobson-Galán et al. 2020b) USSN from He star + NS system (Nakaoka et al. 2021) -CCSN from low-mass stripped progenitor (De et al. 2021)
SN 2021gno	SBa galaxy/4.5 kpc	low SFR	-Merger of low-mass, hybrid WDs (Jacobson-Galán et al. 2022) -CCSN from stripped progenitor (this work)
SN 2021inl	E/S0 galaxy/23.3 kpc	No	Merger of low-mass, hybrid WDs (Jacobson-Galán et al. 2022)

Note. ^aThe galaxy name is not reported, yet a detailed analysis of its properties can be found in De et al. (2018a).

Another point raised by Jacobson-Galán et al. (2022) is the lack of radio detections, which would be inconsistent with a high mass-loss rate necessary to remove the H-rich envelope of a massive progenitor. We note, however, that in our proposed scenario, the H-rich material is lost upon leaving the main sequence. That is, at least $\sim 10^4$ yr before the explosion. At an assumed expansion velocity of 100 km s^{-1} , this material would lie beyond $\sim 10^{18}$ cm from the progenitor, i.e. far beyond the distance of $\sim 10^{16}$ cm that is probed by the radio constraints.

Lastly, Jacobson-Galán et al. (2022) favour a WD over a CC origin for SN 2021gno also based on the low star formation rate derived from a limit on the local $H\alpha$ luminosity of $L_{H\alpha} < 4.3 \times 10^{36} \text{ erg s}^{-1}$. However, such a limit is still compatible with what is observed for a substantial fraction of normal SESNe, as shown by Kuncarayakti et al. (2018). Therefore, a massive-star origin for SN 2021gno cannot be ruled out based on this star formation rate constraint.

Notwithstanding the above considerations, we note that the massive star progenitor model for SN 2021gno we propose presents some caveats to be taken into account. First, a more detailed binary evolution model should be performed to confirm that a low-mass, low-Ni SN is indeed possible. Secondly, our model has an initial He-core mass that would imply a relatively massive progenitor, which could conflict with the lack of $H\alpha$ emission and thus low SFR at the SN location. This limit is compatible with a low though non-negligible fraction (≈ 20 per cent) of the stripped-envelope SN population.

We have also investigated the possible presence of hydrogen features in the spectra of SN 2021gno. Our SYNOW analysis suggests the possible presence of $H\alpha$ near maximum light, although the identification is not certain. De et al. (2021) claim the detection of H features for SN 2019ehk at very early phases (from flash spectroscopy), near maximum light, and at nebular phases. Interestingly, the nebular phase spectra of SN 2021gno do show a bump that is centred at the location of $H\alpha$, as shown in Fig. 13. The figure shows that similar features are common among double-peaked Ca-rich transients. If they were produced by $H\alpha$, these emissions would typically extend out to approximately $8000\text{--}10\,000 \text{ km s}^{-1}$. In the case of SN 2021gno, the line extends roughly between 6000 and 8000 km s^{-1} , in good agreement with the extent of the [Ca II] and [O I] lines (see Fig. 10). However, a definitive identification for this weak emission is not possible, as it may be due to Ca I $\lambda 6572$, to [N II] $\lambda\lambda 6548, 6583$, or to a combination of both (Milisavljevic et al. 2017).

The presence of hydrogen in Ca-rich transients is a key point since it is directly related with the progenitor system and its ability to retain some H-rich material. If confirmed, such noticeable features in the

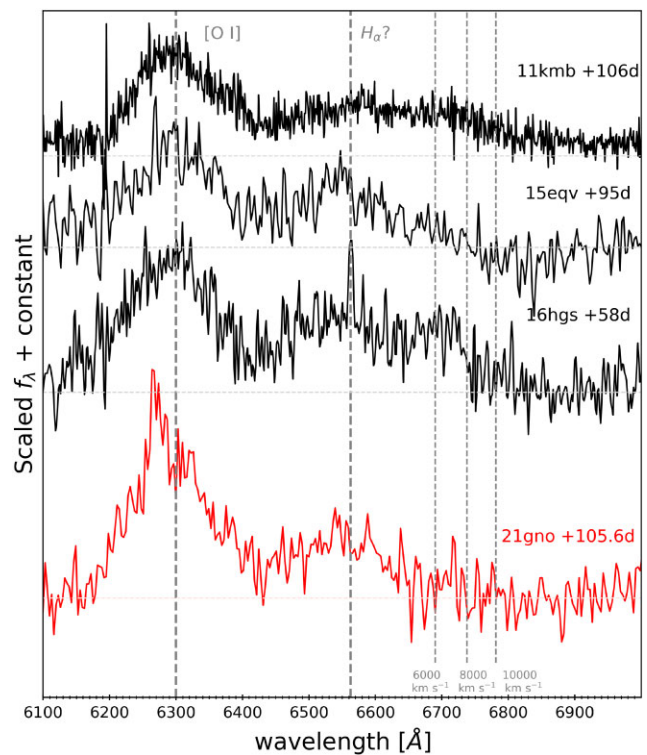


Figure 13. Comparison of the nebular spectrum of SN 2021gno (in red) with those of other Ca-rich transients (in black) around the [O I] $\lambda\lambda 6300, 6363/H\alpha$ region. Transients PTF11kmb (Lunnan et al. 2017), iPTF15eqv (Milisavljevic et al. 2017), and iPTF16hgs (De et al. 2018b) are shown with their epochs relative to maximum light. The dashed horizontal lines indicate the continuum level in the range of $6850\text{--}7000 \text{ \AA}$. The dashed vertical lines indicate the reference wavelengths of 6300 \AA for the [O I] doublet, and 6562.8 \AA for $H\alpha$. The dotted vertical lines show the position of $H\alpha$ at $6000, 8000,$ and $10\,000 \text{ km s}^{-1}$.

maximum-light spectra would imply hydrogen masses of $\sim 10^{-3}\text{--}10^{-2} M_{\odot}$ (Dessart et al. 2011; Hachinger et al. 2012). The proposed scenarios involving the explosion of a WD star would be difficult to reconcile with this observation, as the hydrogen mass should be below $10^{-4} M_{\odot}$ (Zenati, Toonen & Perets 2019). A more natural explanation may be that of a relatively massive progenitor leading to the collapse of a stripped core surrounded by an extended envelope or CSM containing traces of hydrogen. If the observed features are due

to $H\alpha$, then SN 2021gno would belong to a small Type IIb class of Ca-rich transients. Confirmation of the presence of H requires synthetic spectra calculations applied to the explosion scenarios proposed for this type of transients, and specifically in the low ejected mass regime of $M_{ej} \lesssim 1 M_{\odot}$.

It is worth noting that this discussion is valid for the subsample of double-peaked Ca-rich transients we are considering here. Nevertheless, for other Ca-rich events, such as SN 2018lqo and SN 2021inl, WD progenitors are favoured and a massive star origin would not be suitable. As we discussed above, in the case of SN 2021gno, it appears that both formation channels are plausible. While the co-existence of two distinct progenitor channels may introduce additional complexity to the overall scenario, it is worth contemplating both until further observational and modelling efforts are done to shed more light on the matter.

6 CONCLUSIONS

We have presented follow-up observations of SN 2021gno obtained by the POISE project and collaborators. According to its photometric and spectroscopic properties, this event belongs to the calcium-rich transient class. This is determined by its low luminosity, rapid photometric evolution, early evolution to the nebular phase, and strong Ca emission in its nebular spectra. The very early UV and optical data showed double-peaked optical light curves, and extremely fast initial declines in the UV bands. The presence of an initial light-curve maximum is a characteristic that is shared with only a small number of calcium-rich transients. In addition, SN 2021gno appeared well within its host galaxy (in projected location), as opposed to the bulk of the calcium-rich class.

At maximum light, the spectroscopic properties of SN 2021gno are similar to those of Type Ib SNe, as most Ca-rich transients. We calculated synthetic spectra using SYNOW to identify spectral lines. Particularly, we analysed the feature near 6250 Å and discussed the possibility of it being due to $H\alpha$ at high velocity, Si II at photospheric velocity, or a combination of both. Additionally, we analysed the nebular spectra, finding a double-peaked [O I] $\lambda\lambda 6300, 6364$ profile and a blue-shifted [Ca II] $\lambda\lambda 7291, 7324$ profile. There is a weak emission around 6250 Å in the nebular spectra, whose identification is also ambiguous. It could be produced by $H\alpha$, Ca I, or [N II]. We hope future observations of Ca-rich transients will explore the H identification since its presence or lack thereof is crucial to understand the progenitor configuration.

We present a hydrodynamical model that reproduces the bolometric light curve and line velocity evolution. We found that the observations are compatible with the explosion of a highly stripped massive star with an ejecta mass of $0.8 M_{\odot}$ and a ^{56}Ni mass of $0.024 M_{\odot}$. The initial cooling phase is explained by the presence of an extended CSM containing $\sim 10^{-2} M_{\odot}$ with an extension of $1100 R_{\odot}$. This opens the possibility of two different progenitor channels that may coexist to explain the origin of double-peaked Ca-rich transients.

SN 2021gno is yet another example that indicates that high-cadence early observations, as well as deep late-time observations, are crucial for determining the physical origin of Ca-rich transients and the external properties of their progenitors. Continuing such rapid follow-up efforts will help to understand whether or not early peaks in the light curves are intrinsic to this type of transients.

ACKNOWLEDGEMENTS

This work was funded by ANID, Millennium Science Initiative, ICN12_009. LM acknowledges support from UNRN PI2022

40B1039 grant. MO acknowledges support from UNRN PI2022 40B1039 and grant PICT-2020-SERIEA-01141. EB and JD are supported in part by NASA grant 80NSSC20K0538. LG acknowledges financial support from the European Social Fund (ESF) 'Investing in your future' under the 2019 Ramón y Cajal program RYC2019-027683-I. LG and TEMB acknowledges financial support from the Spanish Ministerio de Ciencia e Innovación (MCIN), the Agencia Estatal de Investigación (AEI) 10.13039/501100011033 under the PID2020-115253GA-I00 HOSTFLOWS project, and from Centro Superior de Investigaciones Científicas (CSIC) under the PIE project 20215AT016. LG, TEMB, and NER acknowledge partial support from the program Unidad de Excelencia María de Maeztu CEX2020-001058-M. BJS is supported by NSF grants AST-1920392, AST-1911074, and AST-1911074. MDS is funded by the Independent Research Fund Denmark (IRFD, grant number 10.46540/2032-00022B). The Aarhus supernova group is funded by the Independent Research Fund Denmark (IRFD, grant number 8021-00170B and 10.46540/2032-00022B), and by the VILLUM FONDEN (grant number 28021). NER acknowledges partial support from MIUR, PRIN 2017 (grant 20179ZF5KS), from the Spanish MICINN grant PID2019-108709GB-I00 and FEDER funds. HK was funded by the Academy of Finland projects 324504 and 328898. TEMB acknowledges financial support from Centro Superior de Investigaciones Científicas (CSIC) under the I-LINK 2021 LINKA20409. JTH is supported by NASA award 80NSSC22K0127. LK is supported by the Hungarian National Research, Development and Innovation Office grant PD-134784. TS, LK, and KV are supported by the János Bolyai Research Scholarship of the Hungarian Academy of Sciences. RKT, TS, and KV are supported by the ÚNKP 22-4, and Bolyai + grant 22-5 and ÚNKP-22-5-ELTE-1093 New National Excellence Programs of the Ministry for Culture and Innovation from the source of the National Research, Development and Innovation Fund, respectively. MG is supported by the EU Horizon 2020 research and innovation programme under grant agreement No 101004719. MN is supported by the European Research Council (ERC) under the European Union's Horizon 2020 research and innovation programme (grant agreement No. 948381) and by a Fellowship from the Alan Turing Institute.

The POISE collaboration is grateful to the Carnegie TAC for generous time allocation on the Swope telescope and to Las Campanas technical staff for excellent assistance during the observations. Based on observations collected at the European Southern Observatory under ESO programmes 1103.D-0328 and 0105.D-0511, and on observations made with the Gran Telescopio Canarias (GTC), installed in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias, in the island of La Palma. This project has received funding from the European Union's Horizon 2020 research and innovation programme under grant agreement No 101004719. This project has been supported by the GINOP-2-3-2-15-2016-00033 project of the National Research, Development and Innovation Office of Hungary (NKFIH) funded by the European Union, as well as by NKFIH grants OTKA FK-134432, 2019-2.1.11-TÉT-2019-00056, OTKA K131508 and the Élviol grant KKP-143986. Authors acknowledge the financial support of the Austrian-Hungarian Action Foundation (101öu13). This research has made use of data and/or software provided by the High Energy Astrophysics Science Archive Research Center (HEASARC), which is a service of the Astrophysics Science Division at NASA/GSFC. This work has made use of the NASA/IPAC Extragalactic Database (NED), which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

DATA AVAILABILITY

The photometry presented in this paper is available in Tables A1, A2, A3, and A4 and the spectra via the WISEREP¹⁹ archive (Yaron & Gal-Yam 2012).

REFERENCES

- Albareti F. D. et al., 2017, *ApJS*, 233, 25
- Appenzeller I. et al., 1998, *The Messenger*, 94, 1
- Bacon R. et al., 2001, *MNRAS*, 326, 23
- Bersten M. C., Hamuy M., 2009, *ApJ*, 701, 200
- Bersten M. C., Benvenuto O., Hamuy M., 2011, *ApJ*, 729, 61
- Bersten M. C. et al., 2012, *ApJ*, 757, 31
- Branch D. et al., 2002, *ApJ*, 566, 1005
- Breeveld A. A., Landsman W., Holland S. T., Roming P., Kuin N. P. M., Page M. J., 2011, in McEnery J. E., Racusin J. L., Gehrels N. eds, AIP Conf. Proc. Vol. 1358, Gamma Ray Bursts 2010. Am. Inst. Phys., New York, p. 373
- Brown P. J. et al., 2009, *AJ*, 137, 4517
- Brown T. M. et al., 2013, *PASP*, 125, 1031
- Burns C. R. et al., 2011, *AJ*, 141, 19
- Burns C. et al., 2021, *Astron. Telegram*, 14441, 1
- Burrows D. N. et al., 2005, *Space Sci. Rev.*, 120, 165
- Buzzoni B. et al., 1984, *The Messenger*, 38, 9
- Cadonau R., Leibundgut B., 1990, *A&AS*, 82, 145
- Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245
- Chambers K. C. et al., 2016, preprint (arXiv:1612.05560)
- Contreras C. et al., 2010, *AJ*, 139, 519
- Dahiwal A., Fremling C., 2021, Transient Name Server Classification Report, 2021-1008, 1
- De K. et al., 2018a, *Science*, 362, 201
- De K. et al., 2018b, *ApJ*, 866, 72
- De K. et al., 2020, *ApJ*, 905, 58
- De K., Fremling U. C., Gal-Yam A., Yaron O., Kasliwal M. M., Kulkarni S. R., 2021, *ApJ*, 907, L18
- Dessart L., Hillier D. J., Livne E., Yoon S.-C., Woosley S., Waldman R., Langer N., 2011, *MNRAS*, 414, 2985
- Dessart L., Hillier D. J., Sukhbold T., Woosley S. E., Janka H. T., 2021, *A&A*, 656, A61
- Dong Y. et al., 2022, *ApJ*, 927, 199
- Dopita M., Hart J., McGregor P., Oates P., Bloxham G., Jones D., 2007, *Ap&SS*, 310, 255
- Eldridge J. J., Stanway E. R., Xiao L., McClelland L. A. S., Taylor G., Ng M., Greis S. M. L., Bray J. C., 2017, *Publ. Astron. Soc. Aust.*, 34, e058
- Ergon M. et al., 2014, *A&A*, 562, A17
- Ertl T., Woosley S. E., Sukhbold T., Janka H. T., 2020, *ApJ*, 890, 51
- Fang Q., Maeda K., 2018, *ApJ*, 864, 47
- Fang Q. et al., 2022, *ApJ*, 928, 151
- Faran T., Nakar E., Poznanski D., 2018, *MNRAS*, 473, 513
- Filippenko A. V., Chornock R., Swift B., Modjaz M., Simcoe R., Rauch M., 2003, *IAU Circ.*, 8159, 2
- Fisher A., Branch D., Hatano K., Baron E., 1999, *MNRAS*, 304, 67
- Fitzpatrick E. L., 1999, *PASP*, 111, 63
- Folatelli G. et al., 2006, *ApJ*, 641, 1039
- Folatelli G. et al., 2014, *ApJ*, 792, 7
- Foley R. J., 2015, *MNRAS*, 452, 2463
- Freudling W., Romaniello M., Bramich D. M., Ballester P., Forchi V., García-Dabó C. E., Moehler S., Neeser M. J., 2013, *A&A*, 559, A96
- Gehrels N. et al., 2004, *ApJ*, 611, 1005
- Gutiérrez C. P. et al., 2021, *MNRAS*, 504, 4907
- Gutiérrez C. P. et al., 2022, *MNRAS*, 517, 2056
- Hachinger S., Mazzali P. A., Taubenberger S., Hillebrandt W., Nomoto K., Sauer D. N., 2012, *MNRAS*, 422, 70
- Hoeflich P., Khokhlov A., 1996, *ApJ*, 457, 500
- Hung T., Tinyanont S., Dimitriadis G., Foley R. J., 2021, Transient Name Server Classification Report, 2021-884, 1
- Jacobson-Galán W. V. et al., 2020a, *ApJ*, 896, 165
- Jacobson-Galán W. V. et al., 2020b, *ApJ*, 898, 166
- Jacobson-Galán W. V. et al., 2022, *ApJ*, 932, 58
- Jerkstrand A., 2017, in Alsabti A. W., Murdin P., eds, *Handbook of Supernovae*. Springer International Publishing AG, New York, p. 795
- Jerkstrand A., Ergon M., Smartt S. J., Fransson C., Sollerman J., Taubenberger S., Bersten M., Spyromilio J., 2015, *A&A*, 573, A12
- Kaiser N. et al., 2002, in Tyson J. A., Wolff S. eds, Proc. SPIE Conf. Ser. Vol. 4836, Survey and Other Telescope Technologies and Discoveries. SPIE, Bellingham, p. 154
- Kasliwal M. M. et al., 2012, *ApJ*, 755, 161
- Kawabata K. S. et al., 2010, *Nature*, 465, 326
- Krisciunas K. et al., 2017, *AJ*, 154, 211
- Kuncarayakti H. et al., 2018, *A&A*, 613, A35
- Landolt A. U., 1992, *AJ*, 104, 340
- Lantz B. et al., 2004, in Mazuray L., Rogers P. J., Wartmann R. eds, Proc. SPIE Conf. Ser. Vol. 5249, Optical Design and Engineering. SPIE, Bellingham, p. 146
- Lunnan R. et al., 2017, *ApJ*, 836, 60
- Lyman J. D., Levan A. J., James P. A., Angus C. R., Church R. P., Davies M. B., Tanvir N. R., 2016, *MNRAS*, 458, 1768
- Maeda K. et al., 2007, *ApJ*, 666, 1069
- Maeda K. et al., 2008, *Science*, 319, 1220
- Martinez L. et al., 2022, *A&A*, 660, A40
- Masci F. J. et al., 2019, *PASP*, 131, 018003
- Mazzali P. A. et al., 2005, *Science*, 308, 1284
- Mazzali P. A., Sauer D. N., Pian E., Deng J., Prentice S., Ben Ami S., Taubenberger S., Nomoto K., 2017, *MNRAS*, 469, 2498
- Milisavljevic D. et al., 2017, *ApJ*, 846, 50
- Modjaz M. et al., 2014, *AJ*, 147, 99
- Moriya T. J., Eldridge J. J., 2016, *MNRAS*, 461, 2155
- Moriya T. J. et al., 2017, *MNRAS*, 466, 2085
- Morozova V., Piro A. L., Valenti S., 2018, *ApJ*, 858, 15
- Mueller E., Hoeflich P., 1994, *A&A*, 281, 51
- Nakaoka T. et al., 2021, *ApJ*, 912, 30
- Paxton B., Bildsten L., Dotter A., Herwig F., Lesaffre P., Timmes F., 2011, *ApJS*, 192, 3
- Paxton B. et al., 2013, *ApJS*, 208, 4
- Paxton B. et al., 2015, *ApJS*, 220, 15
- Paxton B. et al., 2018, *ApJS*, 234, 34
- Paxton B. et al., 2019, *ApJS*, 243, 10
- Perets H. B. et al., 2010, *Nature*, 465, 322
- Perley D. A., 2021, Transient Name Server Classification Report, 2021-1009, 1
- Piascik S., Steele I. A., Bates S. D., Mottram C. J., Smith R. J., Barnsley R. M., Bolton B., 2014, in Ramsay S. K., McLean I. S., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 9147, Ground-based and Airborne Instrumentation for Astronomy V. SPIE, Bellingham, p. 91478H
- Polin A., Nugent P., Kasen D., 2021, *ApJ*, 906, 65
- Poole T. S. et al., 2008, *MNRAS*, 383, 627
- Prentice S. J., Mazzali P. A., 2017, *MNRAS*, 469, 2672
- Prentice S. J. et al., 2020, *A&A*, 635, A186
- Roming P. W. A. et al., 2005, *Space Sci. Rev.*, 120, 95
- Schlaflly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Sell P. H., Maccarone T. J., Kotak R., Knigge C., Sand D. J., 2015, *MNRAS*, 450, 4198
- Shen K. J., Kasen D., Weinberg N. N., Bildsten L., Scannapieco E., 2010, *ApJ*, 715, 767
- Shen K. J., Quataert E., Pakmor R., 2019, *ApJ*, 887, 180
- Shivvers I. et al., 2013, *MNRAS*, 436, 3614
- Shivvers I. et al., 2019, *MNRAS*, 482, 1545
- Silverman J. M. et al., 2013, *ApJS*, 207, 3
- Smartt S. J. et al., 2015, *A&A*, 579, A40
- Smith J. A. et al., 2002, *AJ*, 123, 2121
- Spergel D. N. et al., 2007, *ApJS*, 170, 377
- Stritzinger M. D. et al., 2018, *A&A*, 609, A135
- Sullivan M. et al., 2011, *ApJ*, 732, 118

¹⁹<https://www.wiserep.org/>

- Tanaka M. et al., 2009, *ApJ*, 700, 1680
 Taubenberger S. et al., 2009, *MNRAS*, 397, 677
 Tauris T. M., Langer N., Moriya T. J., Podsiadlowski P., Yoon S. C., Blinnikov S. I., 2013, *ApJ*, 778, L23
 Tauris T. M., Langer N., Podsiadlowski P., 2015, *MNRAS*, 451, 2123
 Theureau G., Hanski M. O., Coudreau N., Hallet N., Martin J. M., 2007, *A&A*, 465, 71
 Tody D., 1986, in Crawford D. L.ed., Proc. SPIE Conf. Ser. Vol. 627, Instrumentation in Astronomy VI. SPIE, Bellingham, p. 733
 Tominaga N. et al., 2005, *ApJ*, 633, L97
 Tonry J. L. et al., 2018, *ApJ*, 867, 105
 Tucker M. A. et al., 2022, *PASP*, 134, 124502
 Tully R. B. et al., 2013, *AJ*, 146, 86
 Valenti S. et al., 2014, *MNRAS*, 437, 1519
 van Dokkum P. G., 2001, *PASP*, 113, 1420
 Waldman R., Sauer D., Livne E., Perets H., Glasner A., Mazzali P., Truran J. W., Gal-Yam A., 2011, *ApJ*, 738, 21
 Wang L., Hu J., 1994, *Nature*, 369, 380
 Wolf C. et al., 2018, *Publ. Astron. Soc. Aust.*, 35, e010
 Yaron O., Gal-Yam A., 2012, *PASP*, 124, 668
 Yasuda N., Fukugita M., Okamura S., 1997, *ApJS*, 108, 417
 Zenati Y., Toonen S., Perets H. B., 2019, *MNRAS*, 482, 1135

APPENDIX: TABLES

Table A1. Optical photometry of SN 2021gno with Swope Telescope.

Date	JD	<i>u</i>	<i>B</i>	<i>g</i>	<i>V</i>	<i>r</i>	<i>i</i>
2021 Mar 21	2459294.60	17.35 ± 0.03	17.53 ± 0.01	17.52 ± 0.01	17.62 ± 0.02	17.74 ± 0.01	17.97 ± 0.01
2021 Mar 21	2459294.79	17.43 ± 0.03	17.59 ± 0.01	17.55 ± 0.01	17.66 ± 0.01	17.72 ± 0.02	17.93 ± 0.01
2021 Mar 22	2459295.66	17.76 ± 0.03	17.74 ± 0.01	17.64 ± 0.01	17.62 ± 0.02	17.72 ± 0.01	17.81 ± 0.01
2021 Mar 22	2459295.82	17.89 ± 0.03	17.86 ± 0.01	17.71 ± 0.01	17.69 ± 0.01	17.70 ± 0.02	17.84 ± 0.01
2021 Mar 23	2459296.60	18.48 ± 0.04	18.15 ± 0.01	17.99 ± 0.01	17.91 ± 0.03	17.87 ± 0.01	17.87 ± 0.01
2021 Mar 23	2459296.78	18.58 ± 0.03	18.22 ± 0.01	18.05 ± 0.01	17.98 ± 0.02	17.89 ± 0.01	17.87 ± 0.01
2021 Mar 24	2459297.66	19.05 ± 0.04	18.50 ± 0.01	18.28 ± 0.01	18.14 ± 0.01	18.02 ± 0.01	17.98 ± 0.01
2021 Mar 24	2459297.78	19.05 ± 0.04	18.53 ± 0.01	18.32 ± 0.01	18.12 ± 0.01	18.04 ± 0.01	17.97 ± 0.01
2021 Mar 26	2459299.64	19.24 ± 0.07	18.39 ± 0.02	18.19 ± 0.01	17.97 ± 0.01	17.82 ± 0.01	17.82 ± 0.01
2021 Mar 26	2459299.78	19.12 ± 0.06	18.44 ± 0.02	18.16 ± 0.01	17.98 ± 0.01	17.79 ± 0.01	17.80 ± 0.01
2021 Mar 27	2459300.73	19.01 ± 0.08	18.27 ± 0.02	18.02 ± 0.02	17.78 ± 0.01	17.67 ± 0.01	17.61 ± 0.01
2021 Mar 31	2459304.71	18.62 ± 0.05	17.89 ± 0.01	17.59 ± 0.01	17.36 ± 0.01	17.21 ± 0.01	17.16 ± 0.01
2021 Apr 1	2459305.67	18.64 ± 0.04	17.88 ± 0.01	17.56 ± 0.01	17.39 ± 0.02	17.19 ± 0.01	17.11 ± 0.01
2021 Apr 2	2459306.67	18.69 ± 0.04	17.88 ± 0.01	17.57 ± 0.01	17.32 ± 0.01	17.12 ± 0.01	17.04 ± 0.01
2021 Apr 3	2459307.62	18.81 ± 0.05	17.96 ± 0.01	17.71 ± 0.01	17.41 ± 0.01	17.14 ± 0.01	17.03 ± 0.01
2021 Apr 5	2459309.69	19.32 ± 0.04	18.22 ± 0.01	17.86 ± 0.01	17.50 ± 0.02	17.28 ± 0.01	17.09 ± 0.01
2021 Apr 6	2459310.68	19.53 ± 0.04	18.32 ± 0.01	17.90 ± 0.01	17.51 ± 0.01	17.23 ± 0.01	17.01 ± 0.01
2021 Apr 7	2459311.64	19.88 ± 0.05	18.57 ± 0.01	18.09 ± 0.01	17.65 ± 0.01	17.29 ± 0.01	17.08 ± 0.01
2021 Apr 8	2459312.69	–	–	–	–	17.40 ± 0.03	–
2021 Apr 9	2459313.66	20.39 ± 0.07	19.01 ± 0.01	18.50 ± 0.01	17.98 ± 0.01	17.54 ± 0.01	17.08 ± 0.01
2021 Apr 10	2459314.67	20.55 ± 0.07	19.01 ± 0.01	18.59 ± 0.01	18.03 ± 0.02	17.53 ± 0.01	17.29 ± 0.01
2021 Apr 11	2459315.62	20.86 ± 0.10	19.33 ± 0.02	18.78 ± 0.01	18.16 ± 0.03	17.65 ± 0.01	17.39 ± 0.01
2021 Apr 12	2459316.62	21.01 ± 0.10	19.42 ± 0.02	18.86 ± 0.01	18.25 ± 0.03	17.73 ± 0.01	17.45 ± 0.01

Table A2. Optical photometry of SN 2021gno with 0.8m Telescope at Baja and Konkoly Observatories.

Date	JD	<i>B</i>	<i>g</i>	<i>V</i>	<i>r</i>	<i>i</i>	<i>z</i>
2021 Mar 20	2459294.46	17.81 ± 0.12	17.53 ± 0.02	17.59 ± 0.03	17.76 ± 0.02	17.99 ± 0.02	18.38 ± 0.07
2021 Mar 21	2459294.51	17.87 ± 0.12	17.49 ± 0.05	17.56 ± 0.06	17.74 ± 0.04	17.89 ± 0.06	18.68 ± 0.32
2021 Mar 22	2459296.34	18.23 ± 0.17	18.13 ± 0.09	17.91 ± 0.07	17.91 ± 0.04	17.70 ± 0.06	18.16 ± 0.23
2021 Mar 25	2459299.46	–	–	18.17 ± 0.16	18.29 ± 0.13	17.50 ± 0.07	17.87 ± 0.13
2021 Mar 26	2459300.40	18.29 ± 0.12	18.10 ± 0.07	17.84 ± 0.06	17.93 ± 0.05	17.64 ± 0.07	18.59 ± 0.25
2021 Apr 1	2459305.55	18.03 ± 0.08	17.52 ± 0.04	17.27 ± 0.04	17.21 ± 0.03	17.00 ± 0.02	16.99 ± 0.07
2021 Apr 3	2459308.41	18.21 ± 0.12	17.60 ± 0.03	17.27 ± 0.03	17.17 ± 0.03	17.04 ± 0.01	17.24 ± 0.03
2021 Apr 4	2459309.41	18.31 ± 0.11	17.70 ± 0.02	17.39 ± 0.04	17.16 ± 0.03	17.05 ± 0.02	17.19 ± 0.03
2021 Apr 9	2459313.5	19.07 ± 0.14	18.32 ± 0.03	17.85 ± 0.03	17.49 ± 0.02	17.26 ± 0.01	17.38 ± 0.04
2021 Apr 10	2459315.44	19.15 ± 0.12	18.45 ± 0.05	18.16 ± 0.06	17.70 ± 0.03	17.39 ± 0.04	17.46 ± 0.08
2021 Apr 21	2459325.50	19.28 ± 0.33	19.58 ± 0.31	19.02 ± 0.26	18.43 ± 0.11	18.47 ± 0.09	–
2021 Apr 24	2459328.31	19.45 ± 0.23	–	18.68 ± 0.12	18.47 ± 0.09	18.49 ± 0.18	17.70 ± 0.17
2021 June 13	245937933	21.43 ± 0.80	21.97 ± 0.55	22.27 ± 1.00	20.21 ± 0.17	19.51 ± 0.08	19.76 ± 0.27
2021 June 26	2459392.33	20.36 ± 0.21	21.20 ± 0.36	20.63 ± 0.23	22.59 ± 1.03	19.60 ± 0.09	–

Table A3. Optical photometry of SN 2021gno with LCOGT.

Date	JD	<i>B</i>	<i>g</i>	<i>V</i>	<i>r</i>	<i>i</i>
2021 Mar 24	2459297.61	18.44 ± 0.03	–	–	–	–
2021 Mar 26	2459299.52	18.41 ± 0.04	18.24 ± 0.03	–	–	–
2021 Mar 31	2459305.32	17.84 ± 0.02	17.56 ± 0.02	–	–	–
2021 Apr 2	2459307.28	17.90 ± 0.02	17.60 ± 0.01	–	–	–
2021 Apr 14	2459319.37	19.78 ± 0.10	–	–	–	–
2021 Apr 16	2459321.22	19.97 ± 0.17	–	–	–	–
2021 Apr 28	2459332.96	20.20 ± 0.12	19.76 ± 0.08	–	–	18.30 ± 0.03
2021 May 3	2459337.93	20.54 ± 0.07	–	19.35 ± 0.06	–	18.41 ± 0.03
2021 May 8	2459343.25	20.82 ± 0.12	20.24 ± 0.05	19.70 ± 0.05	19.25 ± 0.04	18.62 ± 0.04
2021 May 10	2459345.23	20.88 ± 0.10	20.30 ± 0.04	19.70 ± 0.04	19.32 ± 0.03	18.62 ± 0.03
2021 May 13	2459348.93	20.77 ± 0.09	20.39 ± 0.04	19.92 ± 0.04	19.46 ± 0.03	18.80 ± 0.03
2021 May 18	2459353.38	20.98 ± 0.17	20.59 ± 0.10	20.07 ± 0.08	19.79 ± 0.08	18.93 ± 0.07
2021 May 23	2459358.48	21.18 ± 0.60	–	20.14 ± 0.19	–	19.00 ± 0.07
2021 May 27	2459361.57	20.96 ± 0.29	–	20.33 ± 0.16	–	19.06 ± 0.05
2021 May 28	2459362.87	21.08 ± 0.17	20.86 ± 0.10	20.66 ± 0.13	20.16 ± 0.04	18.98 ± 0.06
2021 May 29	2459363.89	–	20.95 ± 0.07	–	20.10 ± 0.03	19.04 ± 0.01
2021 May 30	2459364.57	20.84 ± 0.20	–	20.55 ± 0.11	–	19.11 ± 0.03
2021 May 30	2459365.30	21.60 ± 0.20	–	20.51 ± 0.08	–	19.15 ± 0.04
2021 June 2	2459367.57	21.67 ± 0.17	–	20.53 ± 0.07	–	19.19 ± 0.03
2021 June 5	2459370.68	–	–	21.02 ± 0.79	–	18.95 ± 0.14

Table A4. Swift-UVOT photometry of SN 2021gno.

Date	JD	<i>w2</i>	<i>m2</i>	<i>w1</i>	<i>u</i>	<i>b</i>	<i>v</i>
2021 Mar 20	2459294.044	15.21 ± 0.06	15.27 ± 0.06	15.58 ± 0.06	16.04 ± 0.07	17.36 ± 0.12	17.58 ± 0.24
2021 Mar 20	2459294.379	15.47 ± 0.06	15.44 ± 0.05	15.64 ± 0.06	16.09 ± 0.07	17.23 ± 0.11	17.71 ± 0.27
2021 Mar 21	2459294.940	16.04 ± 0.06	15.94 ± 0.06	16.08 ± 0.07	16.27 ± 0.07	17.25 ± 0.11	17.65 ± 0.26
2021 Mar 22	2459296.010	17.26 ± 0.08	17.06 ± 0.08	16.97 ± 0.09	16.80 ± 0.10	17.62 ± 0.15	17.50 ± 0.22
2021 Mar 23	2459297.096	18.48 ± 0.13	18.68 ± 0.18	18.24 ± 0.21	17.76 ± 0.20	18.16 ± 0.24	–
2021 Mar 25	2459298.783	19.21 ± 0.22	–	–	18.24 ± 0.30	18.20 ± 0.26	–
2021 Mar 29	2459303.165	19.13 ± 0.20	–	–	17.52 ± 0.18	17.62 ± 0.16	17.70 ± 0.30
2021 Mar 30	2459303.643	19.17 ± 0.22	–	–	17.70 ± 0.21	17.64 ± 0.17	17.07 ± 0.19
2021 Apr 04	2459309.216	19.58 ± 0.26	–	–	18.00 ± 0.25	17.82 ± 0.19	17.25 ± 0.21
2021 Apr 21	2459326.145	19.63 ± 0.33	–	–	–	–	–

Table A5. Bolometric luminosity of SN 2021gno.

Date	JD	Phase [d]	$F_{u \rightarrow z}$ [$\times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$]	F_{qbol} [$\times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$]	F_{bol} [$\times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$]	L_{bol} [$\times 10^{41}$ erg s $^{-1}$]
2021 Mar 20	2459294.04	1.24	–	5.90 ± 0.18	18.6 ± 4.78	34.69 ± 8.91
2021 Mar 20	2459294.37	1.57	–	5.47 ± 0.17	13.0 ± 2.23	24.24 ± 4.16
2021 Mar 20	2459294.46	1.66	–	5.83 ± 0.14	12.0 ± 0.41	22.49 ± 0.76
2021 Mar 21	2459294.51	1.71	–	5.62 ± 0.12	10.7 ± 0.30	20.04 ± 0.55
2021 Mar 21	2459294.6	1.80	1.87 ± 0.02	3.54 ± 0.08	8.20 ± 0.37	15.28 ± 0.69
2021 Mar 21	2459294.79	1.99	1.81 ± 0.02	5.05 ± 0.11	9.29 ± 0.13	17.29 ± 0.24
2021 Mar 21	2459294.94	2.14	1.81 ± 0.02	4.63 ± 0.10	7.00 ± 0.24	13.04 ± 0.45
2021 Mar 22	2459295.66	2.85	1.66 ± 0.01	3.36 ± 0.07	4.31 ± 0.07	8.02 ± 0.14
2021 Mar 22	2459295.82	3.01	1.55 ± 0.01	2.99 ± 0.05	3.79 ± 0.06	7.07 ± 0.11
2021 Mar 22	2459296.01	3.20	1.48 ± 0.02	2.77 ± 0.07	3.39 ± 0.08	6.32 ± 0.15
2021 Mar 22	2459296.34	3.53	1.26 ± 0.04	2.18 ± 0.04	2.72 ± 0.05	5.07 ± 0.10
2021 Mar 23	2459296.6	3.79	1.23 ± 0.01	1.88 ± 0.03	2.33 ± 0.03	4.33 ± 0.06
2021 Mar 23	2459296.78	3.96	1.18 ± 0.01	1.76 ± 0.02	2.19 ± 0.03	4.07 ± 0.05
2021 Mar 23	2459297.09	4.27	1.11 ± 0.01	1.51 ± 0.05	1.93 ± 0.05	3.60 ± 0.10
2021 Mar 24	2459297.66	4.84	0.97 ± 0.01	1.32 ± 0.02	1.71 ± 0.03	3.19 ± 0.05
2021 Mar 24	2459297.78	4.96	0.97 ± 0.01	1.30 ± 0.02	1.70 ± 0.03	3.17 ± 0.05
2021 Mar 25	2459298.78	5.95	0.97 ± 0.02	1.26 ± 0.06	1.68 ± 0.08	3.13 ± 0.15
2021 Mar 25	2459299.46	6.63	–	1.21 ± 0.05	1.90 ± 0.28	3.55 ± 0.52
2021 Mar 26	2459299.64	6.81	1.09 ± 0.01	1.34 ± 0.02	1.83 ± 0.03	3.40 ± 0.06
2021 Mar 26	2459299.78	6.95	1.11 ± 0.01	1.38 ± 0.02	1.88 ± 0.03	3.40 ± 0.06
2021 Mar 26	2459300.4	7.56	1.15 ± 0.03	1.43 ± 0.04	1.84 ± 0.09	3.42 ± 0.18

Table A5 – continued

Date	JD	Phase [d]	$F_{u \rightarrow z}$ [$\times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$]	F_{qbol} [$\times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$]	F_{bol} [$\times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$]	L_{bol} [$\times 10^{41}$ erg s $^{-1}$]
2021 Mar 27	2459300.73	7.89	1.28 ± 0.02	1.57 ± 0.03	2.16 ± 0.04	4.03 ± 0.08
2021 Mar 29	2459303.16	10.31	1.71 ± 0.07	2.16 ± 0.09	2.91 ± 0.21	5.42 ± 0.39
2021 Mar 30	2459303.64	10.78	1.78 ± 0.06	2.24 ± 0.09	3.06 ± 0.19	5.70 ± 0.36
2021 Mar 31	2459304.71	11.85	1.91 ± 0.03	1.91 ± 0.03	3.09 ± 0.04	5.76 ± 0.08
2021 Apr 1	2459305.55	12.68	2.01 ± 0.03	2.01 ± 0.03	3.39 ± 0.11	6.32 ± 0.20
2021 Apr 1	2459305.67	12.80	1.93 ± 0.03	1.93 ± 0.03	3.17 ± 0.04	5.91 ± 0.08
2021 Apr 2	2459306.67	13.79	2.01 ± 0.03	2.01 ± 0.03	3.33 ± 0.04	6.21 ± 0.08
2021 Apr 3	2459307.62	14.74	1.91 ± 0.03	1.91 ± 0.03	3.28 ± 0.04	6.11 ± 0.08
2021 Apr 3	2459308.41	15.52	1.89 ± 0.03	1.89 ± 0.03	3.08 ± 0.07	5.73 ± 0.14
2021 Apr 4	2459309.21	16.32	1.77 ± 0.05	1.98 ± 0.12	3.01 ± 0.38	5.61 ± 0.71
2021 Apr 4	2459309.41	16.52	1.80 ± 0.03	1.80 ± 0.03	3.00 ± 0.07	5.60 ± 0.13
2021 Apr 5	2459309.69	16.79	1.68 ± 0.03	1.68 ± 0.03	3.00 ± 0.04	5.59 ± 0.08
2021 Apr 6	2459310.68	17.78	1.68 ± 0.03	1.68 ± 0.03	3.15 ± 0.04	5.87 ± 0.09
2021 Apr 7	2459311.64	18.73	1.51 ± 0.03	1.51 ± 0.03	2.96 ± 0.04	5.51 ± 0.08
2021 Apr 8	2459312.69	19.78	1.32 ± 0.05	1.30 ± 0.07	3.10 ± 0.29	5.77 ± 0.55
2021 Apr 9	2459313.5	20.58	1.22 ± 0.01	1.22 ± 0.01	2.38 ± 0.07	4.43 ± 0.13
2021 Apr 9	2459313.66	20.74	1.23 ± 0.03	1.23 ± 0.03	3.00 ± 0.05	5.59 ± 0.09
2021 Apr 10	2459314.67	21.74	1.15 ± 0.03	1.15 ± 0.03	2.58 ± 0.04	4.80 ± 0.08
2021 Apr 10	2459315.44	22.51	1.04 ± 0.02	1.04 ± 0.02	2.12 ± 0.12	3.95 ± 0.23
2021 Apr 11	2459315.62	22.69	1.01 ± 0.03	1.01 ± 0.03	2.40 ± 0.04	4.47 ± 0.08
2021 Apr 12	2459316.62	23.68	0.94 ± 0.03	0.94 ± 0.03	2.27 ± 0.04	4.24 ± 0.08
2021 Apr 21	2459325.5	32.51	–	0.43 ± 0.03	0.90 ± 0.15	1.74 ± 0.28
2021 Apr 24	2459328.31	35.30	–	0.46 ± 0.02	0.91 ± 0.12	1.77 ± 0.23
2021 Apr 28	2459332.96	39.92	–	0.22 ± 0.01	0.74 ± 0.07	1.44 ± 0.13
2021 May 3	2459337.93	44.86	–	0.20 ± 0.01	0.63 ± 0.04	1.23 ± 0.08
2021 May 8	2459343.25	50.15	–	0.16 ± 0.01	0.43 ± 0.03	0.85 ± 0.06
2021 May 10	2459345.23	52.12	–	0.15 ± 0.01	0.47 ± 0.03	0.91 ± 0.05
2021 May 13	2459348.93	55.79	–	0.13 ± 0.01	0.38 ± 0.02	0.74 ± 0.04
2021 May 18	2459353.38	60.22	–	0.11 ± 0.01	0.32 ± 0.04	0.62 ± 0.07
2021 May 23	2459358.48	65.28	–	0.10 ± 0.01	0.28 ± 0.08	0.58 ± 0.15
2021 May 27	2459361.57	68.36	–	0.09 ± 0.01	0.27 ± 0.05	0.55 ± 0.10
2021 May 28	2459362.87	69.65	–	0.08 ± 0.01	0.28 ± 0.04	0.55 ± 0.07
2021 May 30	2459364.57	71.34	–	0.08 ± 0.01	0.27 ± 0.04	0.55 ± 0.07
2021 May 30	2459365.3	72.06	–	0.07 ± 0.01	0.25 ± 0.03	0.50 ± 0.06
2021 June 2	2459367.57	74.32	–	0.07 ± 0.01	0.24 ± 0.02	0.49 ± 0.05

¹Facultad de Ciencias Astronómicas y Geofísicas, Universidad Nacional de La Plata, Paseo del Bosque S/N, B1900FWA La Plata, Argentina²Instituto de Astrofísica de La Plata (IALP), CCT-CONICET-UNLP, Paseo del Bosque S/N, B1900FWA La Plata, Argentina³Kavli Institute for the Physics and Mathematics of the Universe (WPI), The University of Tokyo, 5-1-5 Kashiwanoha, Kashiwa, Chiba 277-8583, Japan⁴Universidad Nacional de Río Negro. Sede Andina, Mitre 630, 8400 Bariloche, Argentina⁵European Southern Observatory, Alonso de Córdova 3107, Casilla 19, 8320000 Vitacura, Santiago, Chile⁶Millennium Institute of Astrophysics MAS, Nuncio Monsenor Sotero Sanz 100, Off. 104, 8320000 Providencia, Santiago, Chile⁷Department of Physics, Virginia Tech, Blacksburg, VA 24061, USA⁸Homer L. Dodge Department of Physics and Astronomy, University of Oklahoma, 440 W. Brooks, Norman, OK 73019-2061, USA⁹Hamburger Sternwarte, Gojenbergsweg 112, D-21029 Hamburg, Germany¹⁰Department of Astronomy, The Ohio State University, 140 W. 18th Avenue, Columbus, OH 43210, USA¹¹Center for Cosmology and AstroParticle Physics (CCAPP), The Ohio State University, 191 W. Woodruff Avenue, Columbus, OH 43210, USA¹²Mitchell Institute for Fundamental Physics and Astronomy, Department of Physics and Astronomy, Texas A&M University 4242 TAMU, College Station, TX 77845, USA¹³The Observatories of the Carnegie Institution for Science, 813 Santa Barbara St, Pasadena, CA 91101, USA¹⁴Institute of Space Sciences (ICE, CSIC), Campus UAB, Carrer de Can Magrans, s/n, E-08193 Barcelona, Spain¹⁵Institut d'Estudis Espacials de Catalunya (IEEC), E-08034 Barcelona, Spain¹⁶Department of Physics, Florida State University, 77 Chieftain Way, Tallahassee, FL 32306, USA¹⁷Astrophysics Research Institute, Liverpool John Moores University, IC2, Liverpool Science Park, 146 Brownlow Hill, Liverpool L3 5RF, UK¹⁸Max-Planck-Institut für Astrophysik, Karl-Schwarzschild Str 1, D-85748 Garching, Germany¹⁹Carnegie Observatories, Las Campanas Observatory, Casilla 601, 1700000 La Serena, Chile²⁰Consejo Nacional de Investigaciones Científicas y Técnicas (CONICET), C1425FQD Capital Federal, Argentina²¹TAPIR, Walter Burke Institute for Theoretical Physics, 350-17, Caltech, Pasadena, CA 91125, USA²²Institute for Astronomy, University of Hawai'i at Manoa, 2680 Woodlawn Dr, Honolulu, HI 96822, USA²³Department of Physics and Astronomy, Aarhus University, Ny Munkegade 120, DK-8000 Aarhus C, Denmark²⁴INAF – Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, I-35122 Padova, Italy²⁵Tuorla Observatory, Department of Physics and Astronomy, University of Turku, FI-20014 Turku, Finland²⁶Finnish Centre for Astronomy with ESO (FINCA), University of Turku, FI-20014 Turku, Finland

²⁷Heidelberger Institut für Theoretische Studien, Schloss-Wolfsbrunnengasse 35, D-69118 Heidelberg, Germany

²⁸The Oskar Klein Centre, Department of Astronomy, Stockholm University, AlbaNova, SE-10691 Stockholm, Sweden

²⁹Department of Experimental Physics, Institute of Physics, University of Szeged, Dóm tér 9, H-6720 Szeged, Hungary

³⁰ELKH-SZTE Stellar Astrophysics Research Group, Szegedi út, Kt. 766, H-6500 Baja, Hungary

³¹Konkoly Observatory, Research Centre for Astronomy and Earth Sciences, Konkoly Thege Miklós út 15-17, H-1121 Budapest, Hungary

³²CSFK, MTA Centre of Excellence, Konkoly Thege Miklós út 15-17, H-1121 Budapest, Hungary

³³Gothard Astrophysical Observatory, ELTE Eötvös Loránd University, H-9400 Szombathely, Hungary

³⁴Baja Astronomical Observatory of University of Szeged, Szegedi út, Kt. 766, H-6500 Baja, Hungary

³⁵Eötvös Loránd University, Department of Astronomy, Pázmány Péter sétány 1/A, H-1117 Budapest, Hungary

³⁶Institute of Physics, Eötvös Loránd University, Pázmány Péter sétány 1/A, H-1117 Budapest, Hungary

³⁷MIT Kavli Institute for Astrophysics and Space Research, 70 Vassar Street, Cambridge, MA 02109, USA

³⁸Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, D-53121 Bonn, Germany

³⁹Scottish Universities Physics Alliance (SUPA), School of Physics and Astronomy, University of St Andrews, North Haugh, St Andrews KY16 9SS, UK

⁴⁰Department of Optics & Quantum Electronics, Institute of Physics, University of Szeged, Dóm tér 9, H-6720 Szeged, Hungary

⁴¹Astronomical Observatory, University of Warsaw, Al. Ujazdowskie 4, PL-00-478 Warszawa, Poland

⁴²School of Physics, Trinity College Dublin, The University of Dublin, Dublin 2, Ireland

⁴³Birmingham Institute for Gravitational Wave Astronomy and School of Physics and Astronomy, University of Birmingham, Birmingham B15 2TT, UK

⁴⁴Department of Physics and Astronomy, University of California, Davis, One Shields Avenue, CA 95616, USA

⁴⁵Astrophysics Research Centre, School of Mathematics and Physics, Queen's University Belfast, Belfast BT7 1NN, UK

⁴⁶Comparte de los Pozos 1028, C1222AAL Ciudad Autónoma de Buenos Aires, Argentina

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.